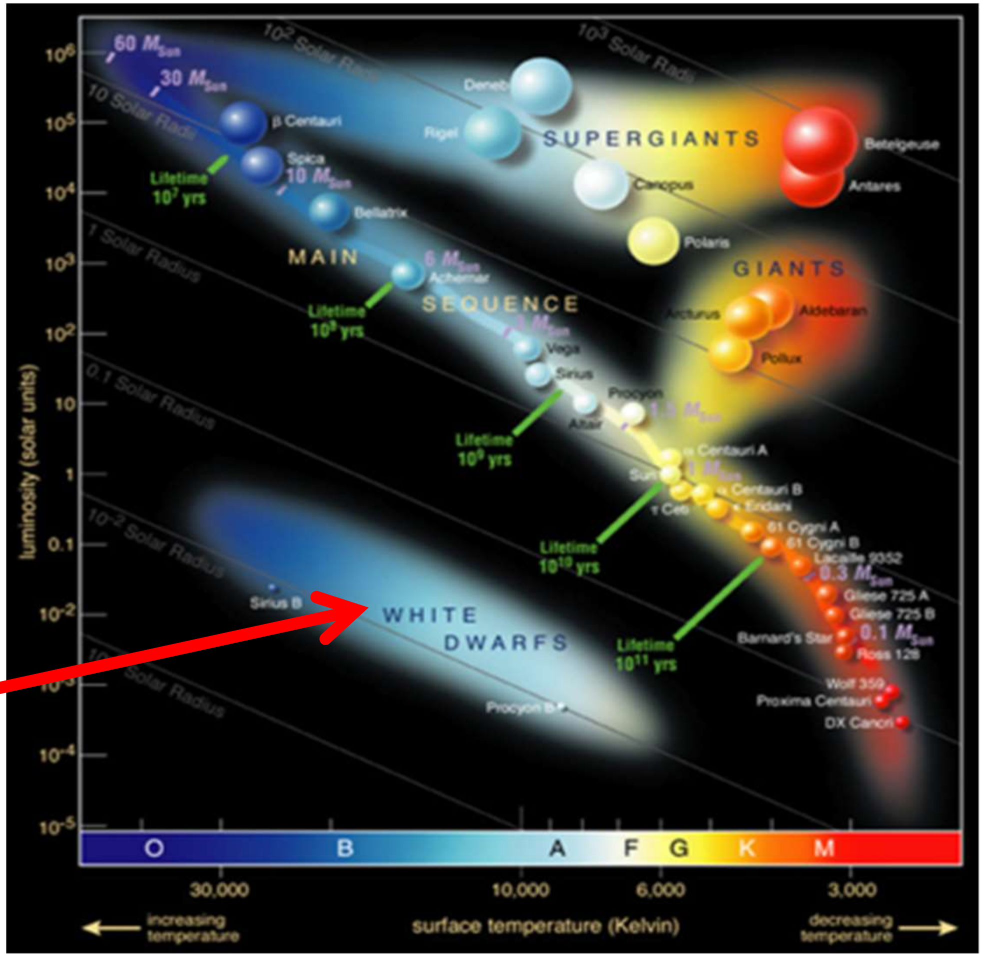
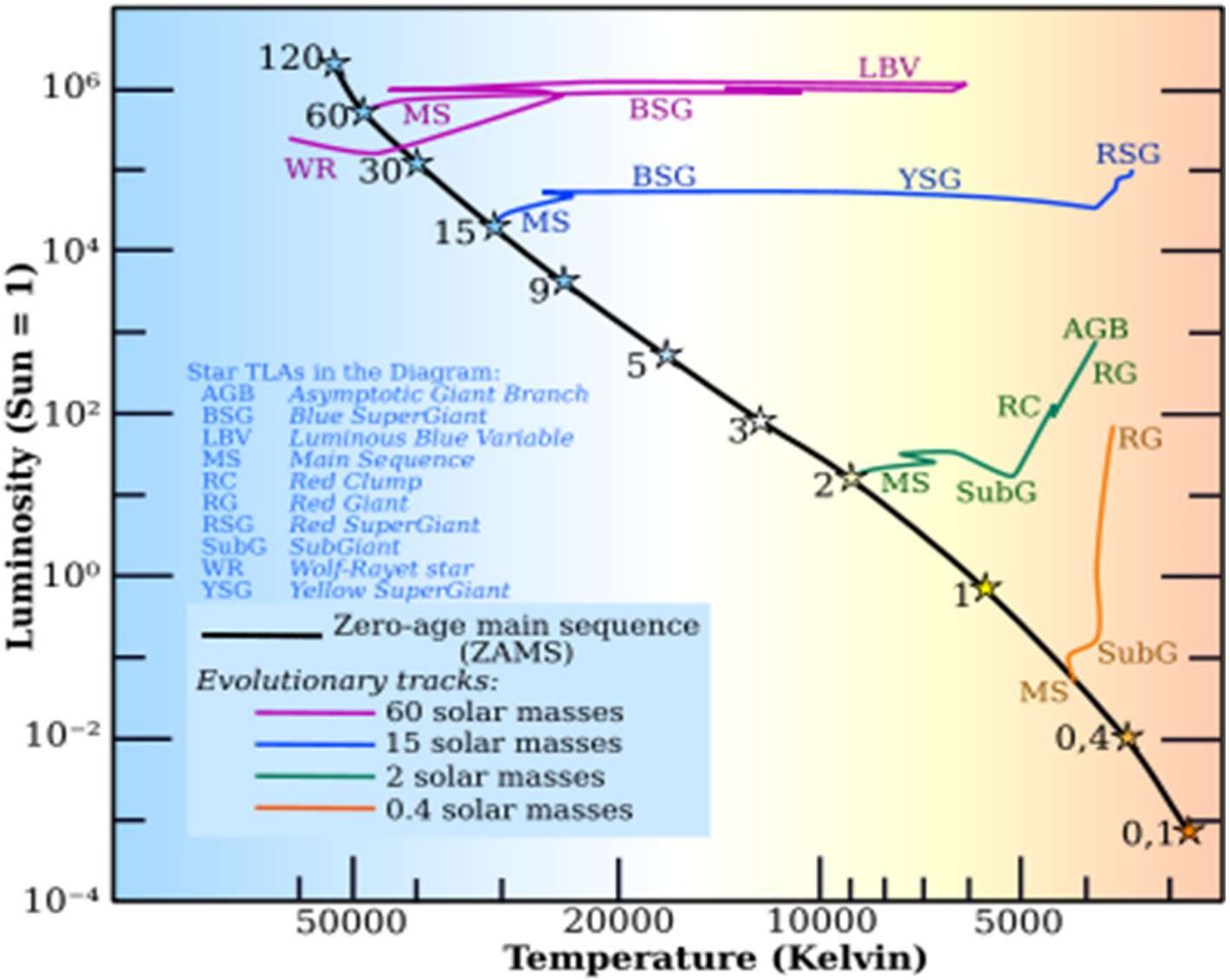
A tour de modeling and analysis of stellar atmospheres throughout the HRD




except for
white dwarfs

## Some different types of stars…

Hot luminous stars:
Massive,
main-sequence (MS)
or evolved, ~10 Rsun.
Strong, fast stellar
winds

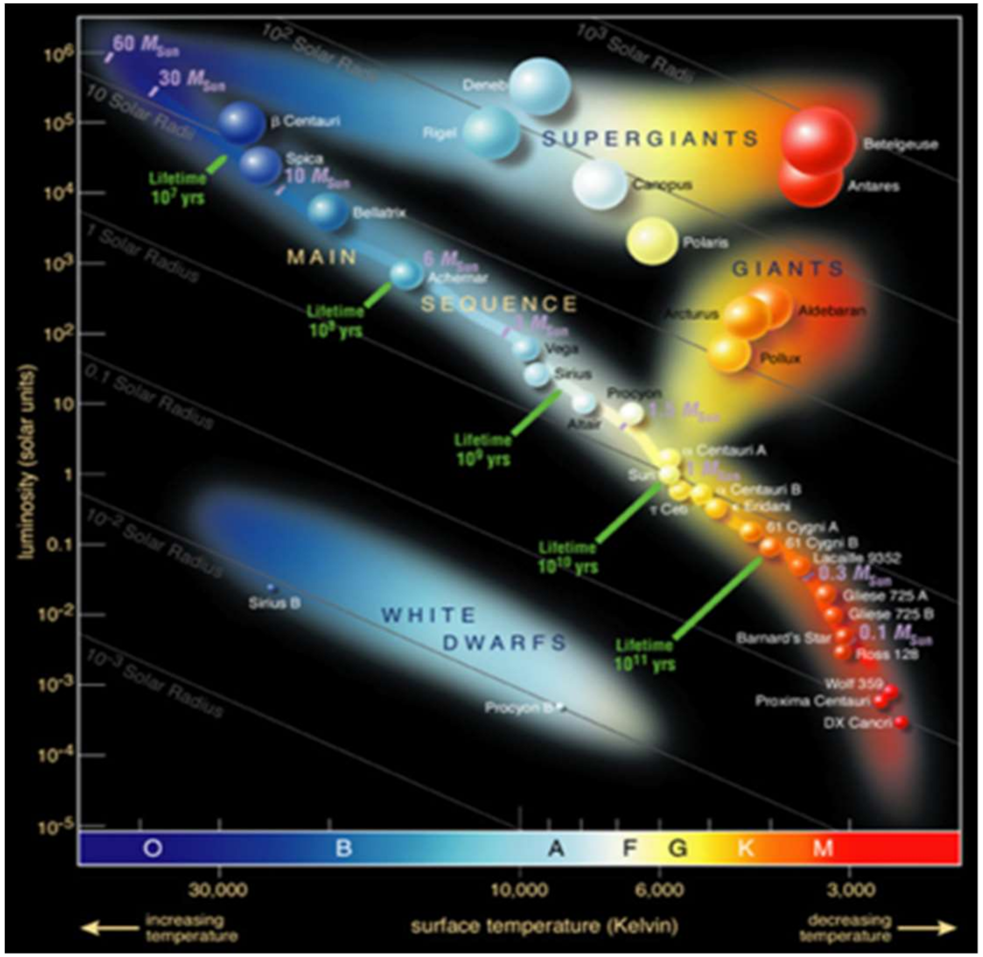



Cool, luminous stars
(RSG, AGB):
Massive or low/interme-
diate mass, evolved,
several 100 (!) Rsun.
Strong, slow stellar winds

Solar-type stars:
Low-mass, on or near MS,
hot surrounding coronae,
weak stellar wind
(e.g. solar wind)

A tour de modeling and analysis of stellar atmospheres throughout the HRD

Different regimes require different key input physics and assumptions

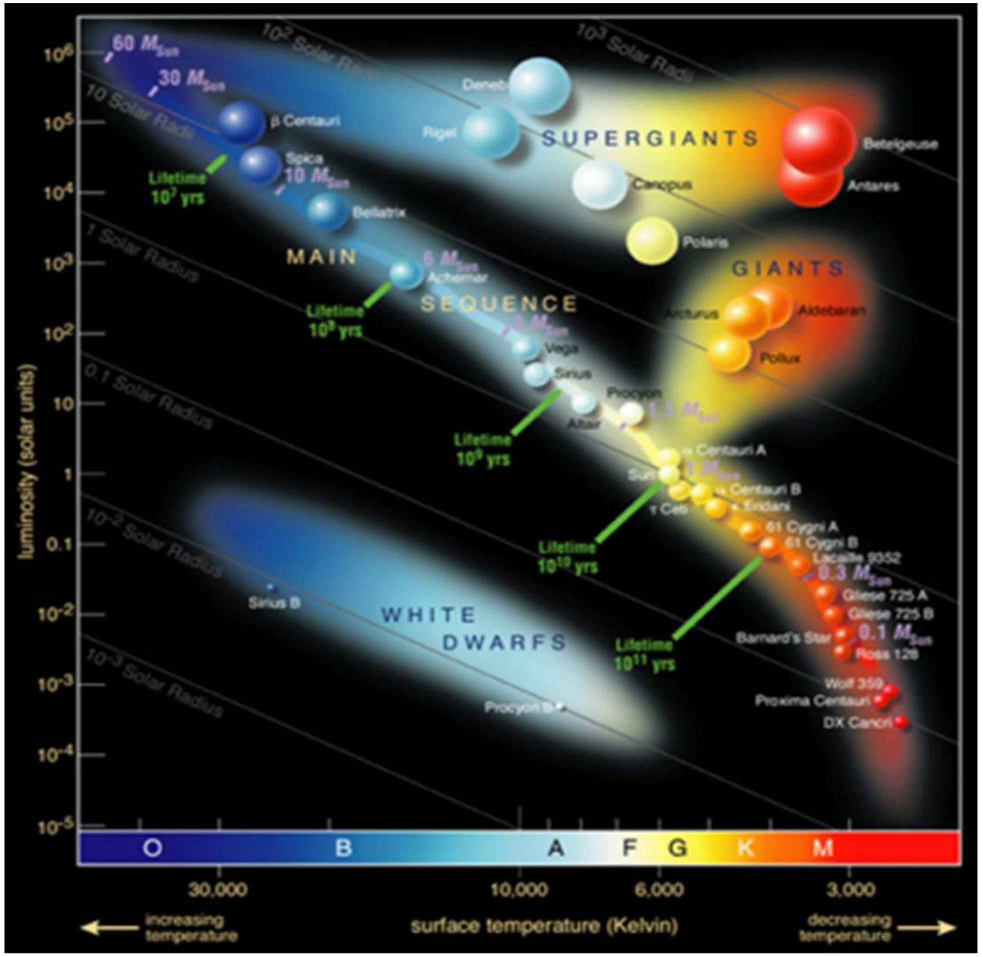



- LTE or NLTE
- Spectral line blocking/blanketing
- (sub-) Surface convection
- Geometry and dimensionality
- Velocity fields and outflows

## Spectroscopy and Photometry

ALSO:
Analysis
of different
WAVELENGTH
BANDS
is different

(X-ray, UV,
optical, infra-
red…)




Depends on where in atmosphere light escapes from

Question: Why is this “formation depth” different for different wavebands and diagnostics?

# Spectroscopy and Photometry (see Chap. 2)

## …gives insight into and understanding of our cosmos

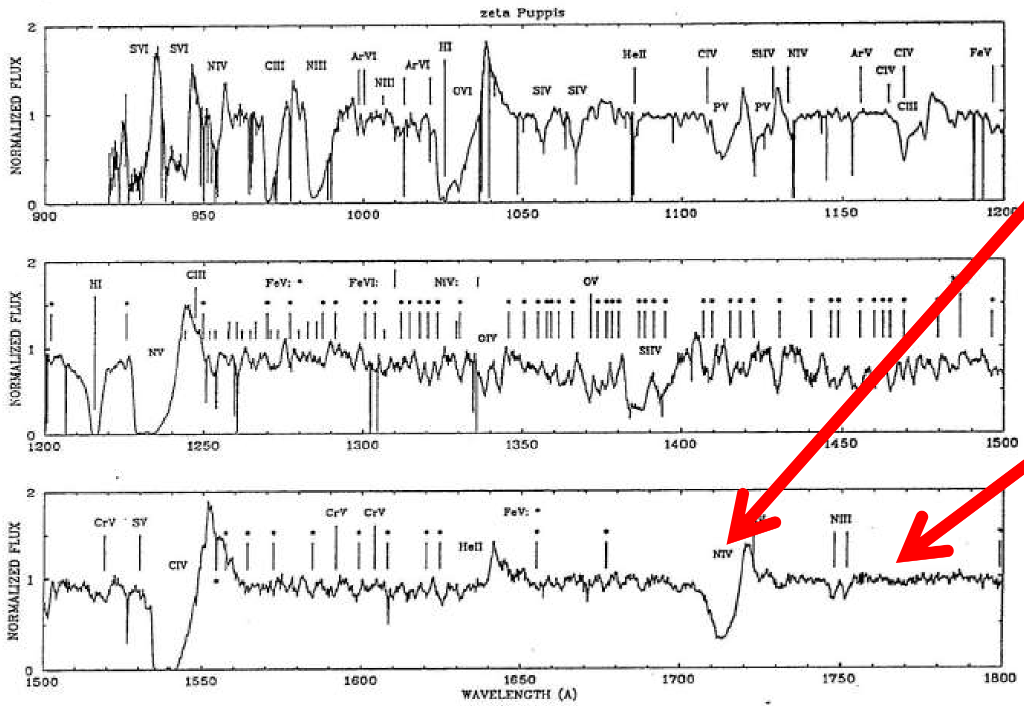
- provides
  - stellar properties, mass, radius, luminosity, energy production, chemical composition,properties of outflows
  - properties of (inter) stellar plasmas, temperature, density, excitation, chemical comp., magnetic fields

- INPUT for stellar, galactic and cosmologic evolution and for stellar and galactic structure

- requires
  - plasma physics, plasma is "normal" state of atmospheres and interstellar matter (plasma diagnostics, line broadening, influence of magnetic fields,…)
  - atomic physics/quantum mechanics, interaction light/matter (micro quantities)
  - radiative transfer, interaction light/matter (macroscopic description)
  - thermodynamics, thermodynamic equilibria: TE, LTE (local), NLTE (non-local)
  - hydrodynamics, atmospheric structure, velocity fields, shockwaves,…

## Spectroscopy (see Chap. 2)

### UV spectrum of the O4I(f) supergiant ζ Pup

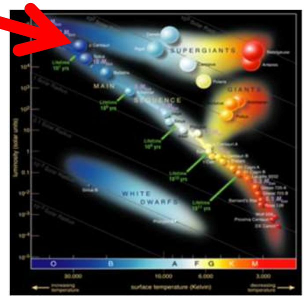



montage of Copernicus ($\lambda < 1500$ Å, high res. mode, $\Delta\lambda \approx 0.05$ Å, Morton & Underhill 1977) and IUE ($\Delta\lambda \approx 0.1$ Å) observations

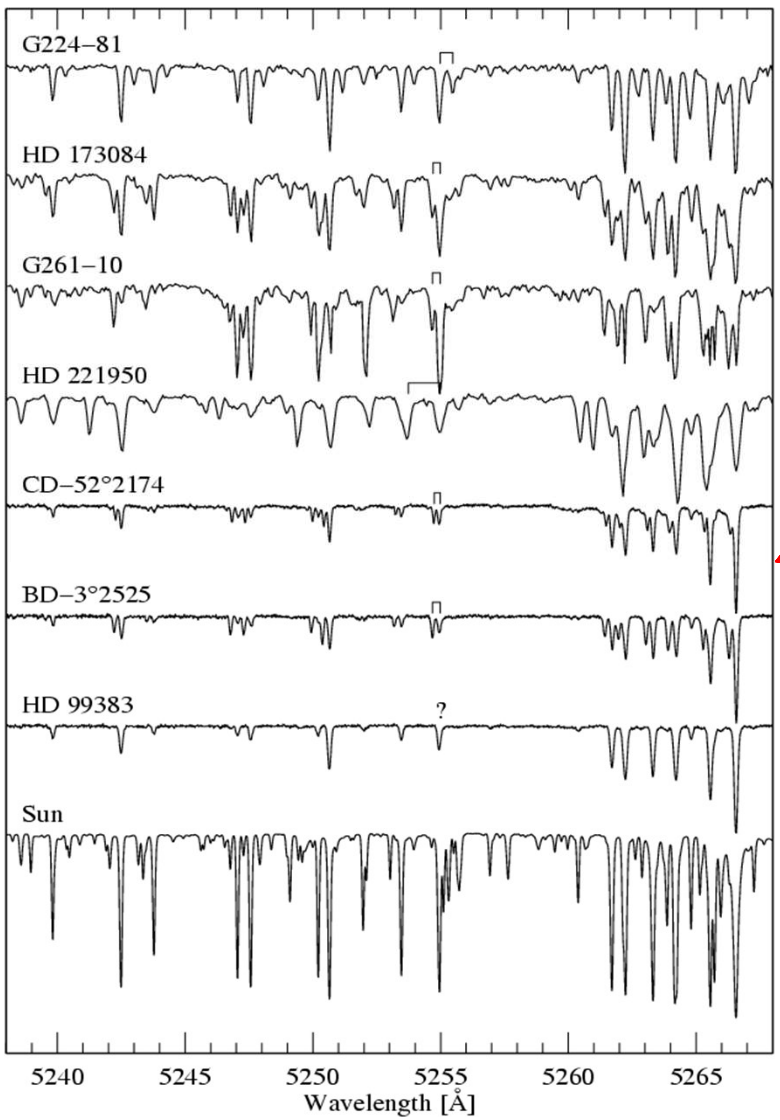
31

UV "P-Cygni" lines formed in rapidly accel-erating, hot stellar winds

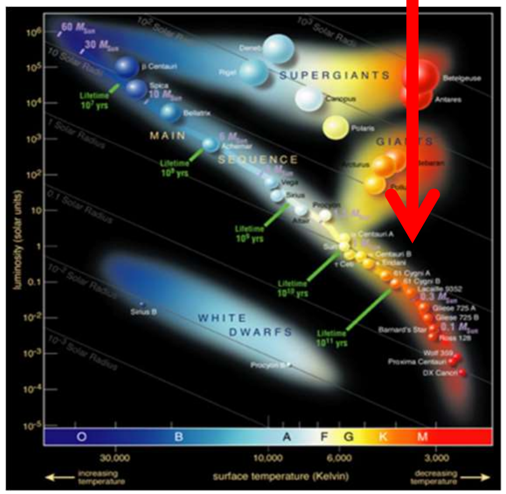
(quasi-) Continuum formed in (quasi-) hydrostatic photosphere



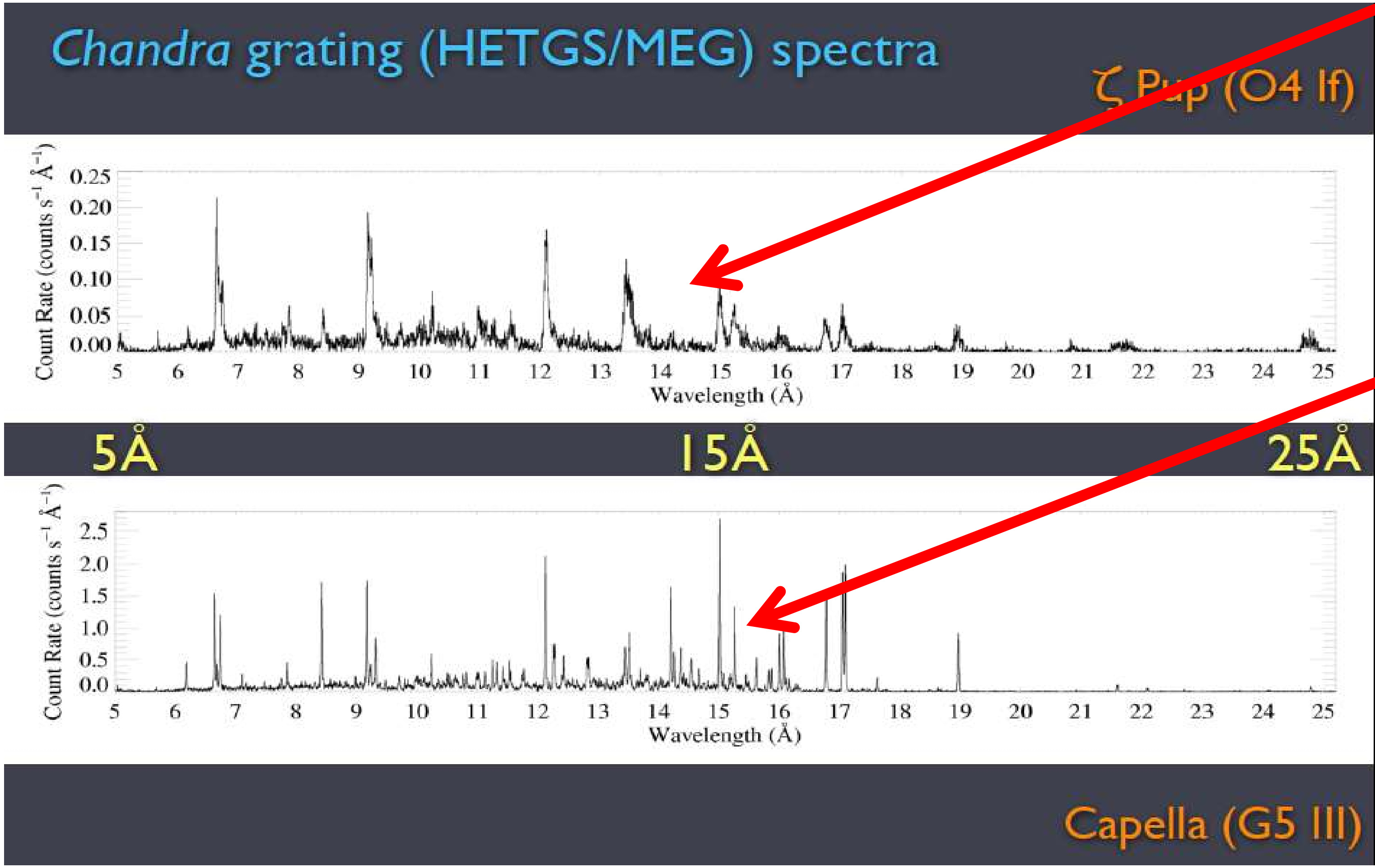
Spectroscopy




Lines and continuum in the optical around 5200 Å, in cool solar-type stars, formed in the photosphere



Spectroscopy




X-rays from hot stars, formed in shocks in stellar wind

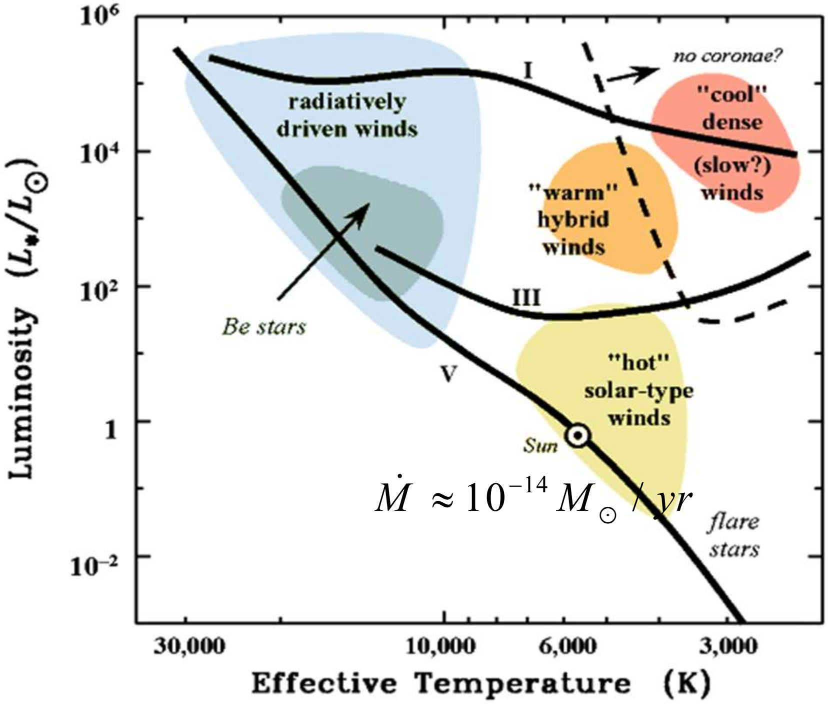
X-rays from cool stars, formed in hot corona

A tour de modeling and analysis of stellar atmospheres throughout the HRD

## Stellar Winds (see Chap. 8/9)

KEY QUESTION: What provides the force able to overcome gravity?

$$\dot{M} \approx 10^{-4} \ldots 10^{-8} M_{\odot} / yr$$




- LTE or NLTE
- Spectral line blocking/blanketing
- (sub-) Surface convection
- Geometry and dimensionality
- Velocity fields and outflows

A tour de modeling and analysis of stellar atmospheres throughout the HRD

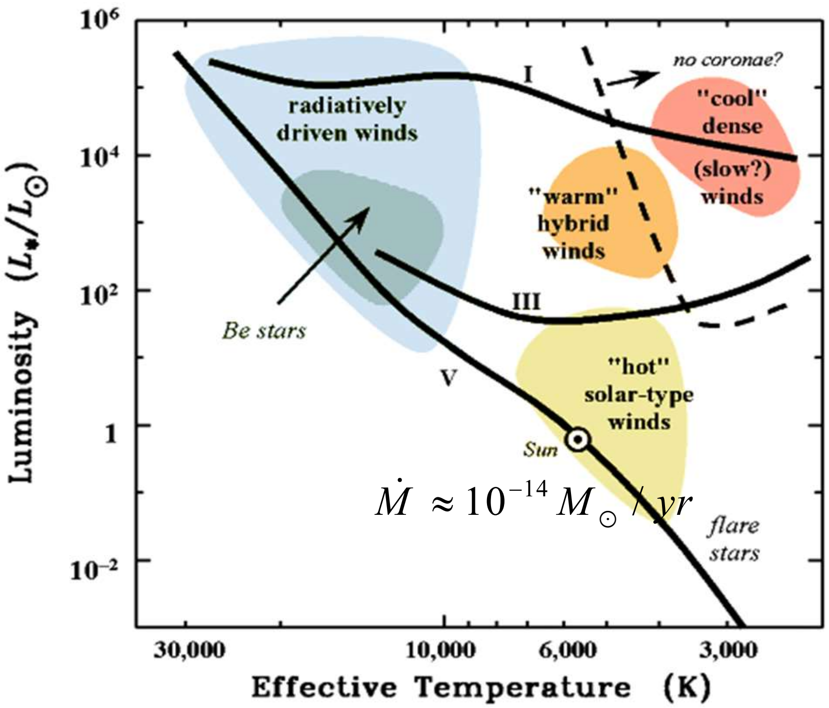
KEY QUESTION: What provides the force able to overcome gravity?

Pressure gradient
in hot coronae of solar-type stars

Radiation force:
Dust scattering (in pulsation-levitated material, see Chap. 8) in cool AGB stars (S. Höffner and colleagues)

Same mechanism in cool RSGs?

$$\dot{M} \approx 10^{-4} \ldots 10^{-8} M_\odot / yr$$




- LTE or NLTE
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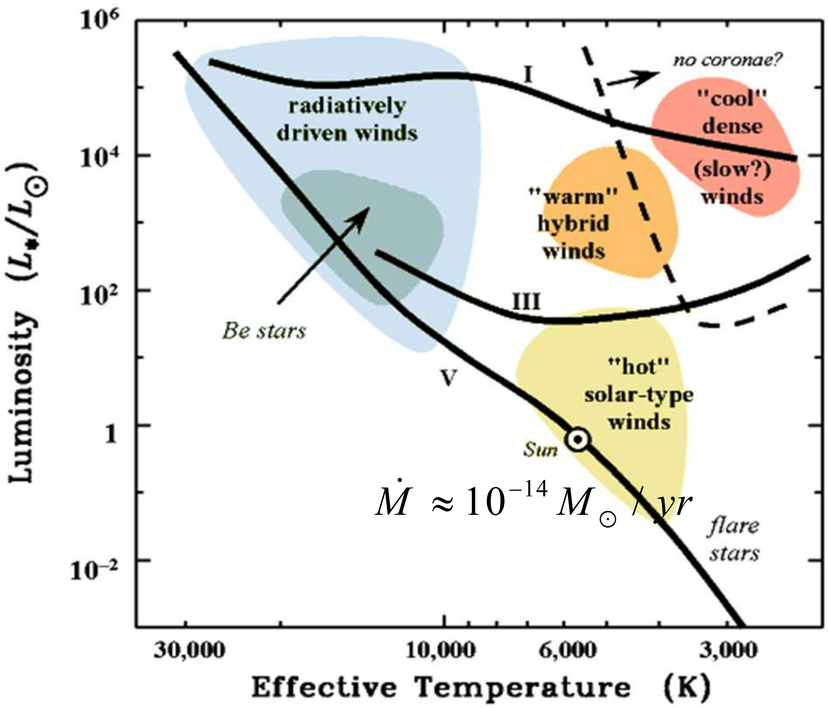
- Geometry and dimensionality
- Velocity fields and outflows

A tour de modeling and analysis of stellar atmospheres throughout the HRD

KEY QUESTION: What provides the force able to overcome gravity?

Radiation force:
line scattering in hot, luminous stars
→ done at USM, more to follow in Chap. 8/9

$$\dot{M} \approx 10^{-4} \ldots 10^{-8} M_{\odot} / yr$$




- LTE or NLTE
- Spectral line blocking/blanketing
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- Geometry and dimensionality
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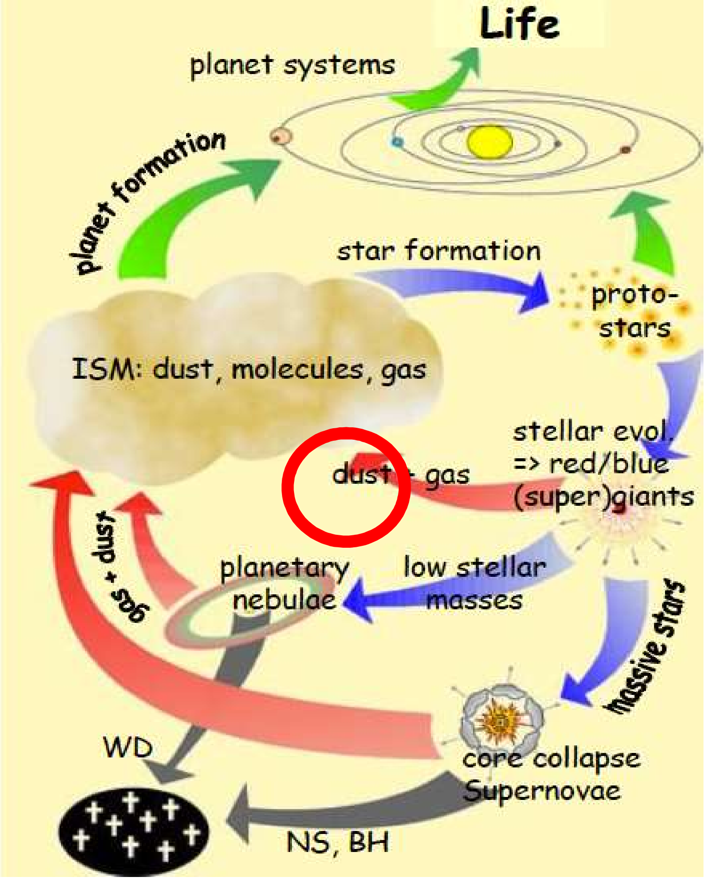
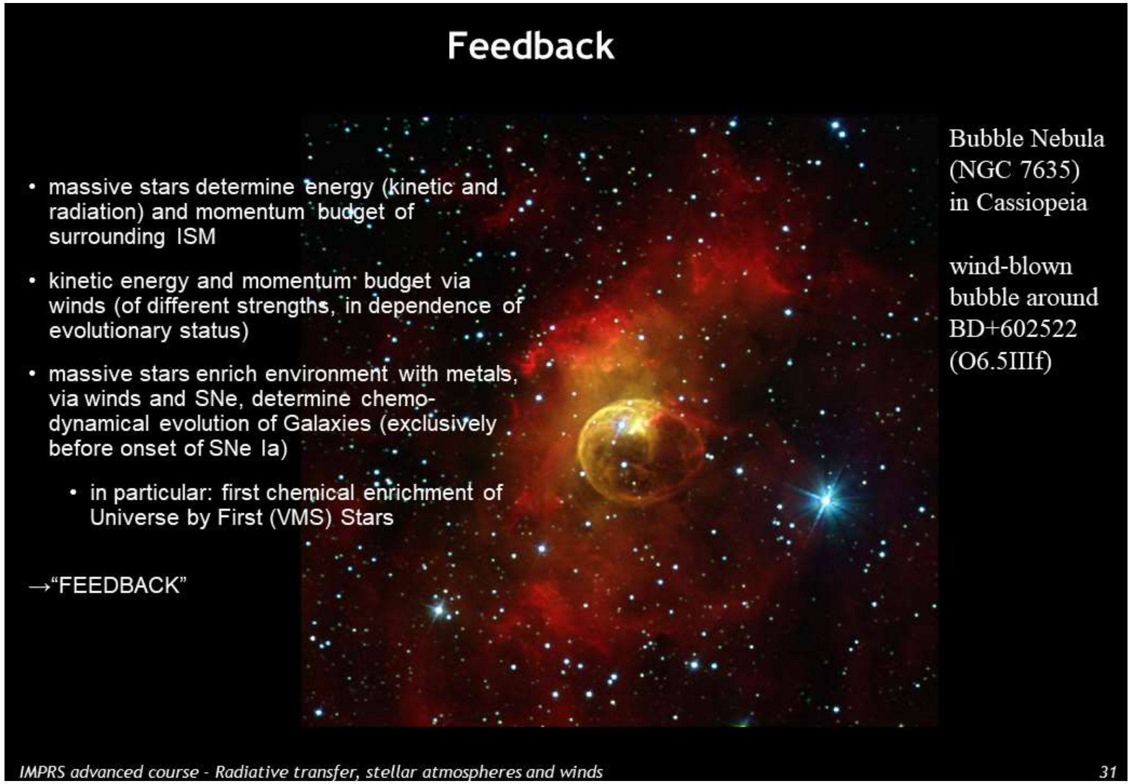
Question: How do you think the high mass loss of stars with high luminosities affects the evolution of the star and its surroundings?

from introductory slides …

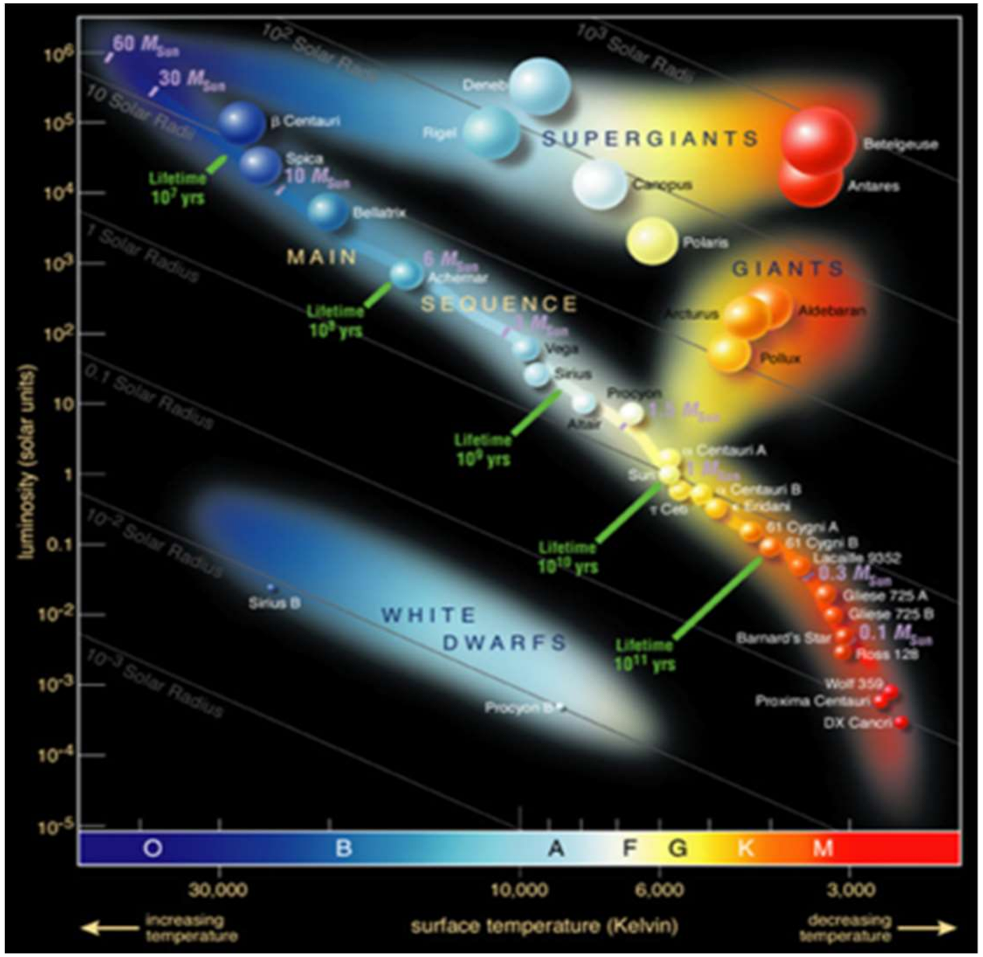



Stellar Winds from hot/evolved cool stars control evolution/late evolution, and feed the ISM with nuclear processed material

A tour de modeling and analysis of stellar atmospheres throughout the HRD

In the following,
we focus on stellar
photospheres



# Stellar Atmospheres in practice

From Chap. 6

## Summary: stellar atmospheres - the solution principle

THUS problem of stellar atmospheres solved (in principle, without convection, p-p geometry, static)

Given $\log g_*$, $T_{eff}$, abundances

(A) hydrostatic equilibrium

$$\frac{dp_{gas}}{dz} = -\rho\,(g_* - g_{rad}); \quad g_{rad} = \frac{4\pi}{c\rho}\int_0^\infty \chi_\nu H_\nu \, d\nu = \frac{4\pi}{c\rho}\left(\sigma^{TH} H(z) + \int_0^\infty \chi_\nu^{rest} H_\nu \, d\nu\right)$$

$$H = \frac{1}{4\pi}\sigma_B T_{eff}^4 \quad \left(= \frac{1}{4\pi}\mathcal{F}\right)$$

$$\Rightarrow \frac{dp_{gas}}{dz} = -\rho\, g_* + \sigma^{TH}\frac{\sigma_B T_{eff}^4}{c} + \frac{4\pi}{c}\int_0^\infty \chi_\nu^{rest} H_\nu \, d\nu$$

(B) equation of rad. transfer

$$\mu\frac{dI_\nu}{dz} = \chi_\nu(S_\nu - I_\nu) \quad \forall \nu,\mu \quad \Rightarrow \quad J_\nu = \frac{1}{2}\int_{-1}^{+1} I_\nu(\mu)\,d\mu; \quad H_\nu = \frac{1}{2}\int_{-1}^{+1} I_\nu(\mu)\,\mu\, d\mu$$

(C) a) radiative equilibrium

scattering terms cancel, since conservative

$$\int_0^\infty (\eta_\nu - \kappa_\nu J_\nu)\,d\nu = \int_0^\infty \left\{(\sigma^{TH} J_\nu + \chi_\nu^{rest} S_\nu^{rest}) - (\sigma^{TH} + \chi_\nu^{rest})J_\nu\right\} d\nu = \int_0^\infty \chi_\nu^{rest}(S_\nu^{rest} - J_\nu)\,d\nu \overset{?}{=} 0$$

b) flux-conservation: $4\pi\int_0^\infty H_\nu(z)\,d\nu = 4\pi H(z) \overset{?}{=} \sigma_B T_{eff}^4$

$\Rightarrow \Delta T(z)$

$\rightarrow \Delta\chi_\nu(z)$ etc

(D) equation of state $p_{gas}(z) = \frac{k_B}{\mu m_H}\rho(z)\,T(z)$

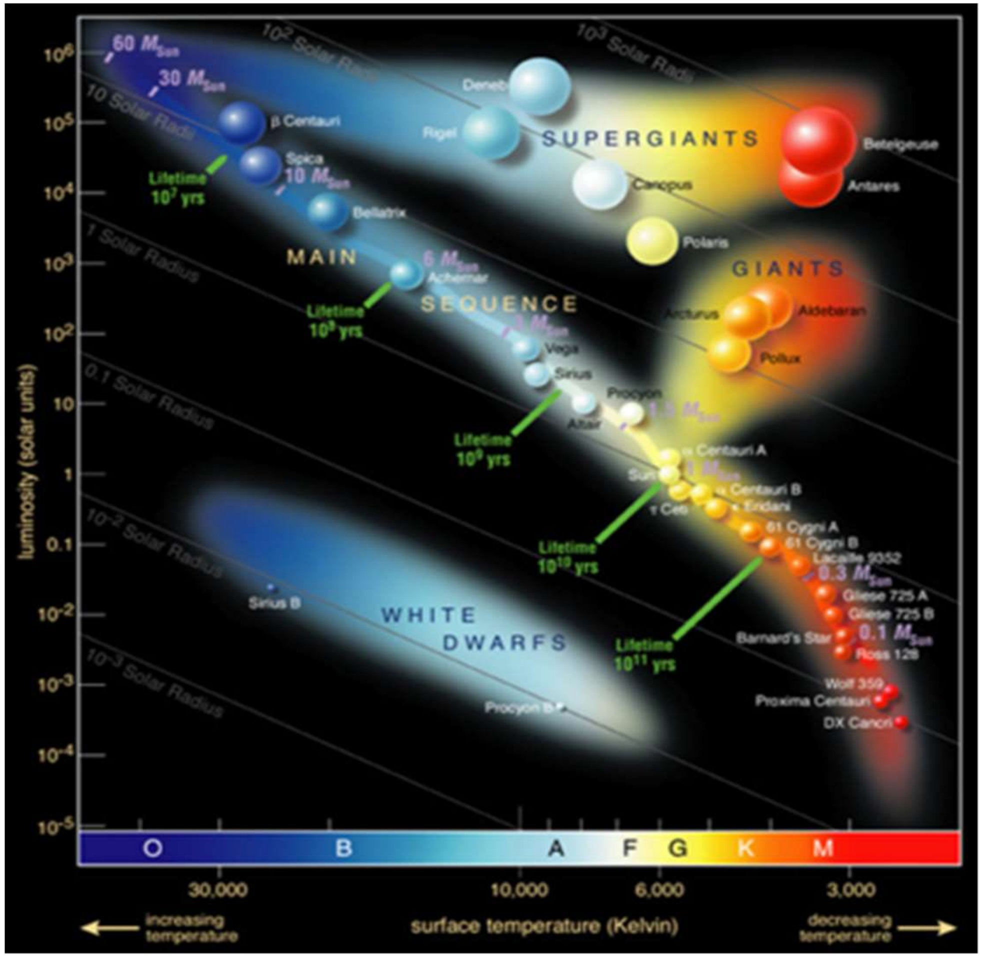
solution by iteration!

! OBSERVATIONS!!

Solution of differential equations A and B by discretization
differential operators => finite differences
all quantities have to be evaluated on suitable grid

Eq. of radiative transfer (B)
usually solved by the so-called
Feautrier and/or Rybicki scheme

A tour de modeling and analysis of stellar atmospheres throughout the HRD




- LTE or NLTE
- Spectral line blocking/blanketing
- (sub-) Surface convection
- Geometry and dimensionality
- Velocity fields and outflows

# LTE or NLTE? (see part 1)

**When is LTE valid???**

roughly: electron collisions $\propto n_e T^{1/2}$ >> photoabsorption rates $\propto I_\nu(T) \propto T^x,\ x \geq 1$

LTE: T low, $n_e$ high — dwarfs (giants), late B and cooler
NLTE: T high, $n_e$ low — all supergiants + rest

however: NLTE-effects also in cooler stars, e.g.. iron in sun

## HOT STARS:

Complete model atmosphere and synthetic spectrum must be calculated in NLTE

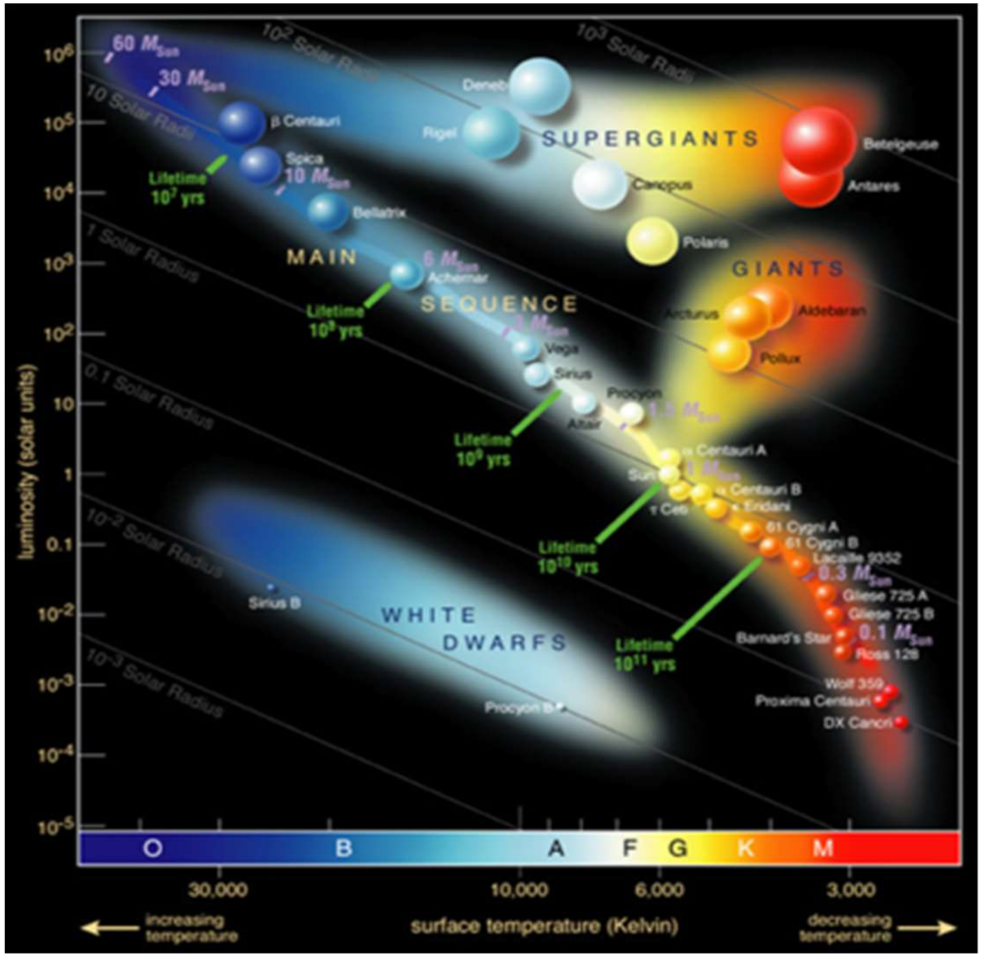
NLTE calculations for various applications (including Supernovae remnants) within the expertise of USM

## COOL STARS:

Standard to neglect NLTE-effects on atmospheric structure, might be included when calculating line spectra for individual “trace” elements (typically used for chemical abundance determinations)

BUT: See work by Phoenix-team (Hauschildt et al.)
ALSO: RSGs still somewhat open question

A tour de modeling and analysis of stellar atmospheres throughout the HRD




- LTE or NLTE
- Spectral line blocking/blanketing
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## Spectral line blocking/blanketing

- Effects of numerous -- literally millions -- of (primarily metal) spectral lines upon the atmospheric structure and flux distribution
- Q: Why is this tricky business?

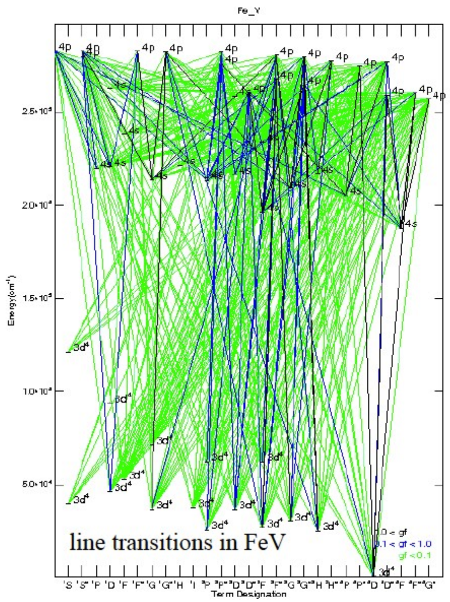
## Spectral line blocking/blanketing

- Effects of numerous -- literally millions -- of (primarily metal) spectral lines upon the atmospheric structure and flux distribution
- Q: Why is this tricky business?

- Lots of atomic data required (thus atomic physics and/or experiments)
- LTE or NLTE?
- What lines are relevant?
  (i.e., what ionization stages? Are there molecules present?)

### Techniques:

Opacity Distribution Functions
Opacity-Sampling
Direct line by line calculations

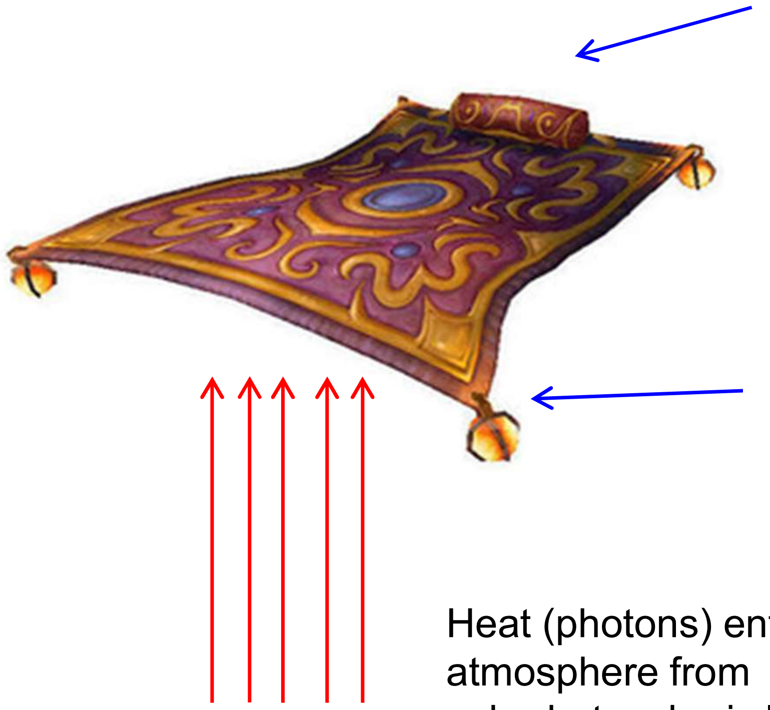


## Spectral line blocking/blanketing

## Back-warming (and surface-cooling)

Numerous absorption lines
 “block” (E)UV radiation flux
Total flux conservation
 demands these photons be
 emitted elsewhere → redistributed to optical/infra-red
Lines act as “blanket”, whereby
 back-scattered line photons are (partly) thermalized and thus heat up deeper layers

“Blanket” typically cools uppermost layers



“Blanket” warms deep layers

Heat (photons) enters atmosphere from sub-photospheric layers

## Spectral line blocking/blanketing

# Back-warming and flux redistribution

…occur in stars of all spectral types

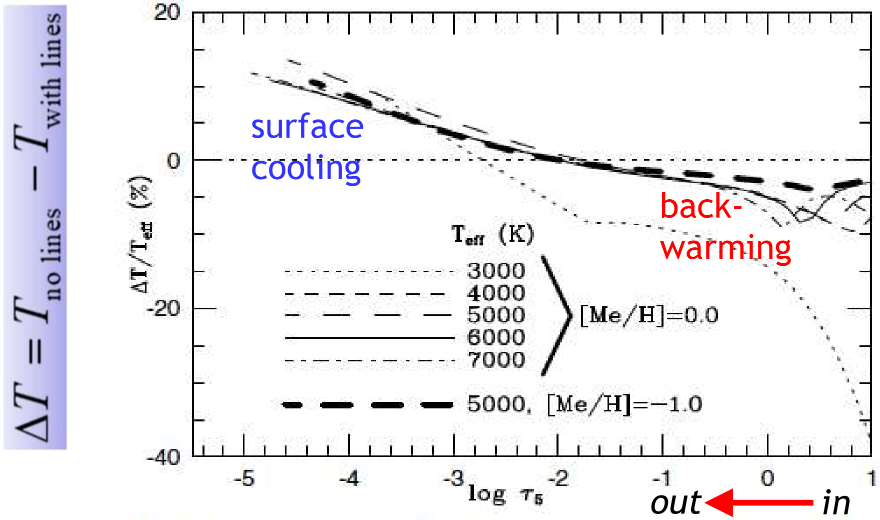



**Fig. 4.** The effects of switching off line absorption on the temperature structure of a sequence of models with $\log g = 3.0$ and solar metallicity. Note that $\Delta T \equiv T(\text{nolines}) - T(\text{lines})$. It is seen that the blanketing effects are fairly independent of effective temperature for models with $T_{\text{eff}} \geq 4000$.

Back warming in cool stars
(from Gustafsson et al. 2008)

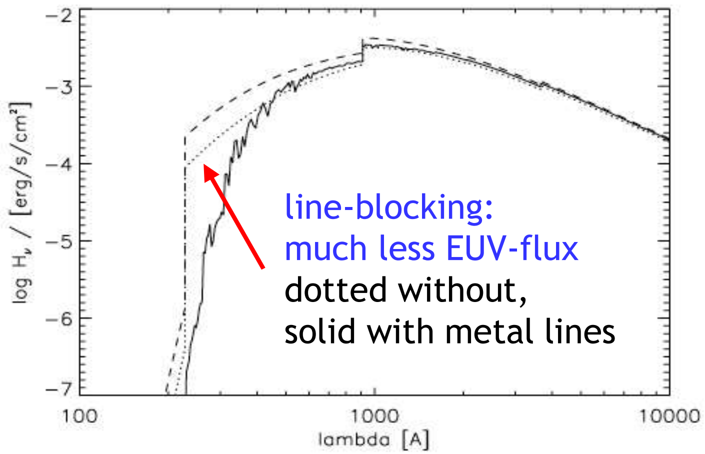



**Fig. 10.** Emergent Eddington flux $H_\nu$ as function of wavelength. Solid line: Current model of HD 15629 (O5V((f)) with parameters from Table 1 ($T_{\text{eff}} = 40\,500$ K, $\log g = 3.7$, "model 1"). Dotted: Pure H/He model without line-blocking/blanketing and negligible wind, at same $T_{\text{eff}}$ and $\log g$ ("model 2"). Dashed: Pure H/He model, but with $T_{\text{eff}} = 45\,000$ K and $\log g = 3.9$ ("model 3").

UV to optical flux redistribution in hot stars
(from Repolust, Puls & Hererro 2004)

## Spectral line blocking/blanketing

# Back-warming and flux redistribution

…occur in stars of all spectral types

From Puls et al. 2008

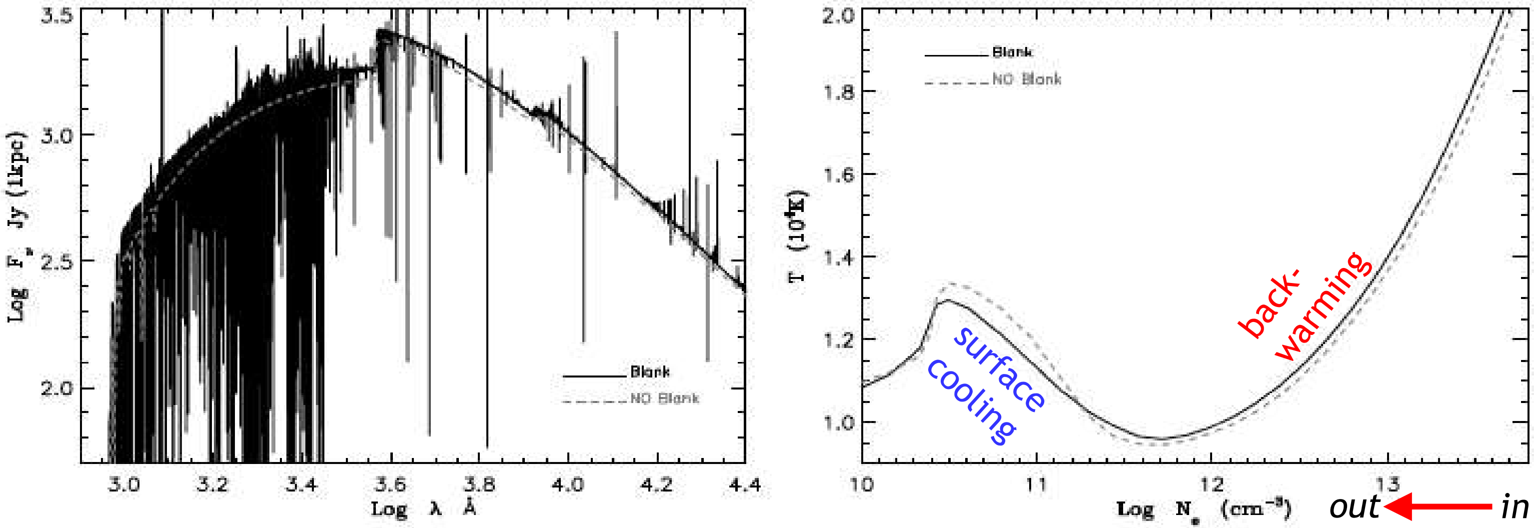



**Fig. 9** Effects of line blanketing (solid) vs. unblanketed models (dashed) on the flux distribution ($\log F_\nu$ (Jansky) vs. $\log \lambda$ (Å), left panel) and temperature structure ($T(10^4$ K) vs. $\log n_e$, right panel) in the atmosphere of a late B-hypergiant. Blanketing blocks flux in the UV, redistributes it towards longer wavelengths and causes back-warming.

# Spectral line blocking/blanketing

in line/continuum forming regions, blanketed models at a certain $T_{eff}$ have a plasma temperature corresponding to an unblanketed model with higher $T'_{eff}$

## Back-warming – effect on effective temperature

RECALL: $T_{eff}$ -- or total flux (plane-parallel) -- fundamental input parameter in model atmosphere!

$$F = \sigma_B T_{\rm eff}^4$$

$T_{eff}$ in cool stars derived, e.g., by optical photometry

From Gustafsson et al. 2008: Estimate effect by assuming a blanketed model with $T_{eff}$ such that the deeper layers correspond to an unblanketed model with effective temperature $T'_{eff} > T_{eff}$

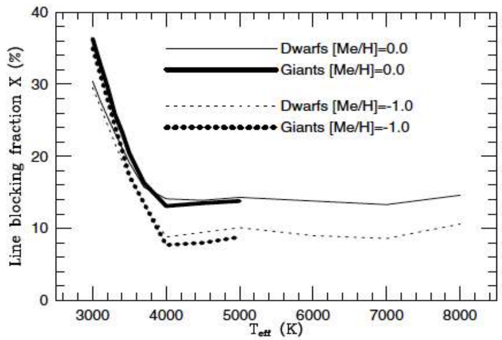



**Fig. 3.** The blocking fraction $X$ in percent for models in the grid with two different metallicities. The dwarf models all have $\log g = 4.5$ while the giant models have $\log g$ values increasing with temperature, from $\log g = 0.0$ at $T_{eff} = 3000$ K to $\log g = 3.0$ at $T_{eff} = 5000$ K.

Question: Why does the line blocking fraction increase for very cool stars?

$$T'_{\rm eff} = (1 - X)^{-\frac{1}{4}} \cdot T_{\rm eff}, \quad (35)$$

where $X$ is the fraction of the integrated continuous flux blocked out by spectral lines,

$$X = \frac{\int_0^\infty (F_{\rm cont} - F_\lambda) d\lambda}{\int_0^\infty F_{\rm cont} d\lambda}. \quad (36)$$

## Spectral line blocking/blanketing

### Back-warming – effect on effective temperature

RECALL: $T_{eff}$ -- or total flux (plane-parallel) -- fundamental input parameter in model atmosphere!

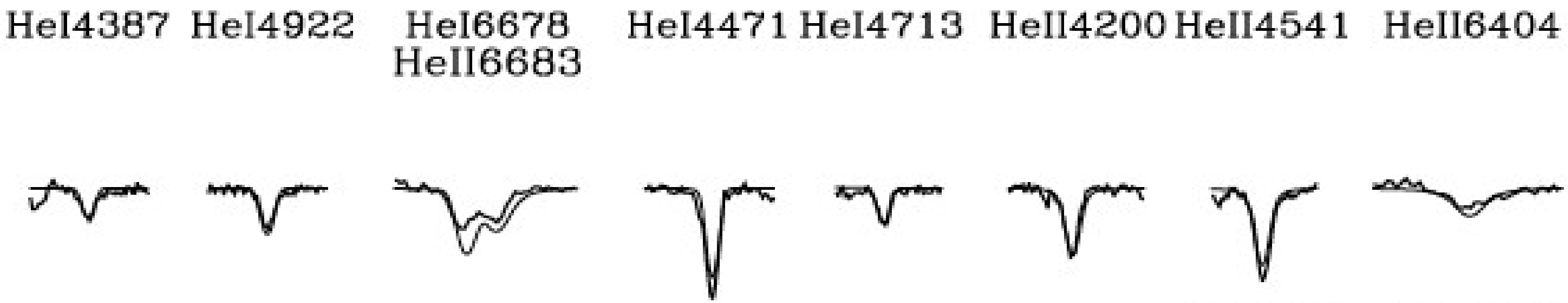
Previous slide were LTE models. In hot stars, everything has to be done in NLTE…

$$F = \sigma_B T_{eff}^4$$

Question: Why is optical photometry generally NOT well suited to derive Teff in hot stars?

# Spectral line blocking/blanketing

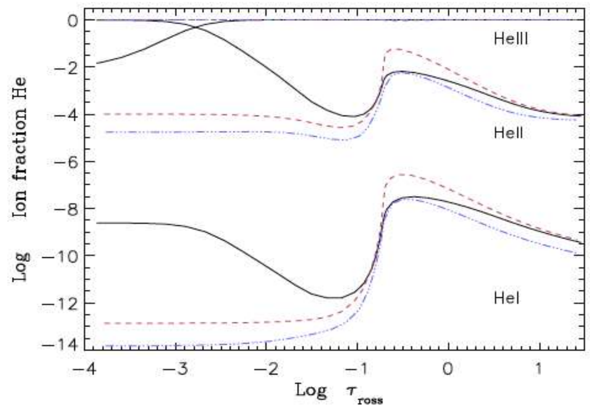
Instead, He ionization-balance is typically used
(or N for the very hottest stars, or, e.g., Si for B-stars)




Simultaneous fits to observed HeI and HeII lines
–– from Repolust, Puls, Hererro (2004)

- Back-warming shifts ionization balance toward more completely ionized Helium in blanketed models

→ thus fitting the same observed spectrum requires lower $T_{eff}$ than in unblanketed models



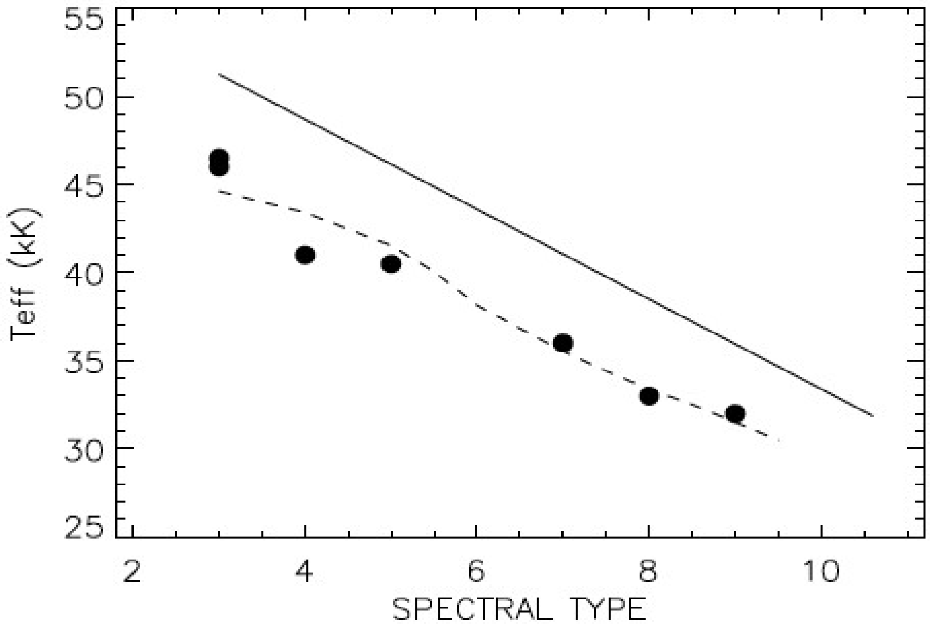

- black – blanketed Teff=45 kK
- red – unblanketed Teff=45 kK
- blue – unblanketed Teff= 50 kK

black and blue have similar (low) HeI/II ionization fractions in weak-line forming region, thus similar line profiles

## Spectral line blocking/blanketing

Instead, He ionization-balance is typically used
(or N for the very hottest stars, or, e.g., Si for B-stars)

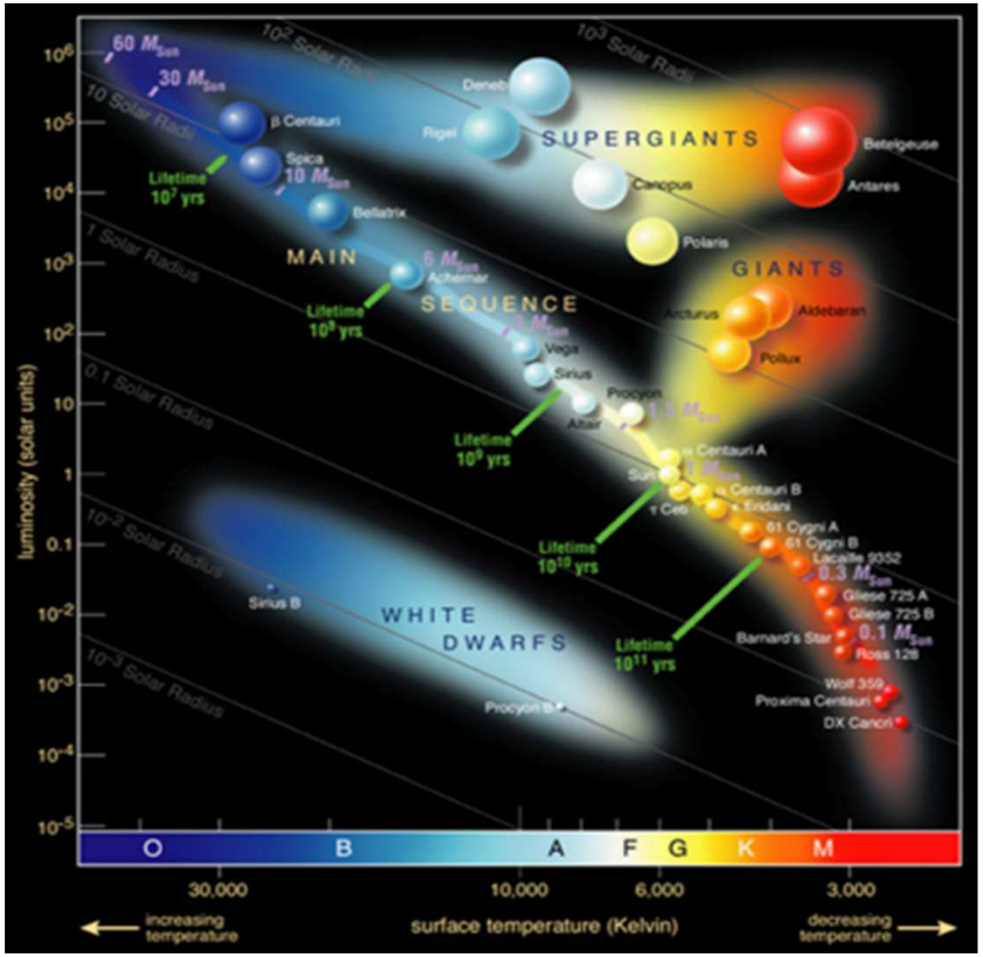
Result: In hot O-stars with Teff~40,000 K, back-warming can lower the derived $T_{eff}$ as compared to unblanketed models by several thousand degrees! (~ 10 %)




New $T_{eff}$ scale for O-dwarf stars. Solid line – unblanketed models. Dashed – blanketed calibration, dots – observed blanketed values (from Puls et al. 2008)

# Stellar Atmospheres in practice

A tour de modeling and analysis of stellar atmospheres throughout the HRD




- LTE or NLTE
- Spectral line blocking/blanketing
- (sub-) Surface convection
- Geometry and dimensionality
- Velocity fields and outflows

## Surface Convection

from part 1

### Convection

energy transport not only by radiation,
however also by

- waves
- heat conduction
- convection

$\left.\begin{array}{l} \text{waves} \\ \text{heat conduction} \end{array}\right\}$ not efficient in typical stellar atmospheres, but ... coronae, chromospheres white dwarfs

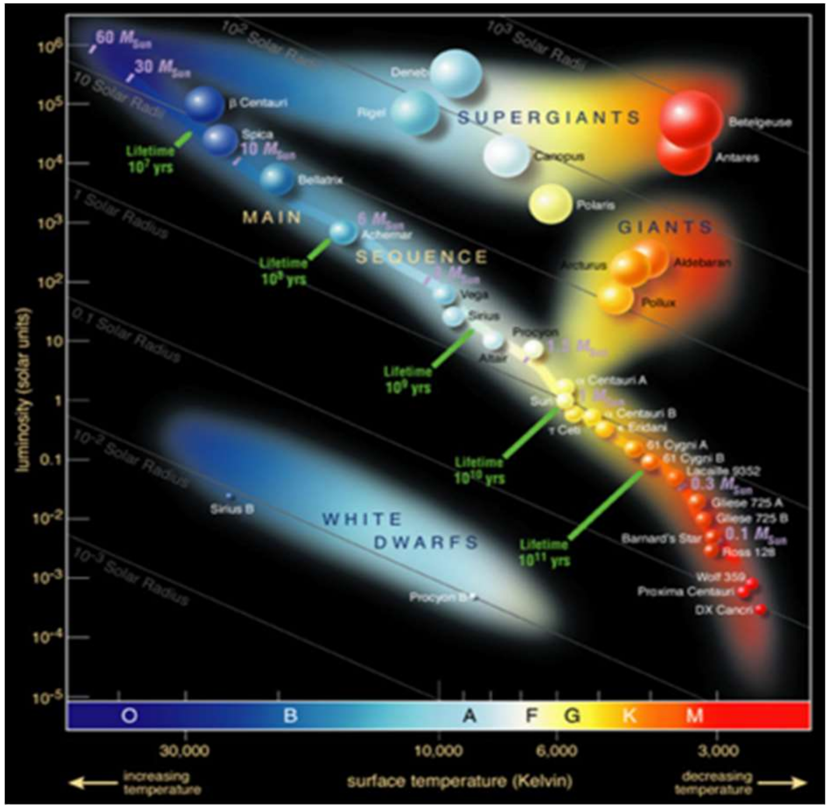
Thus

$$\nabla \cdot \left( \mathcal{F}^{\mathrm{rad}} + \mathcal{F}^{\mathrm{conv}} \right) \overset{\substack{\text{total flux = const} \\ \downarrow}}{=} 0$$

(in quasi-hydrostatic atmospheres)

# Surface Convection

OBSERVATIONS:
“Sub-surface” convection in layers T~160,000 K (due to iron-opacity peak) currently discussed also in hot stars




- H/He recombines in atmospheres of cool stars

→ Provides MUCH opacity

→ Convective Energy transport

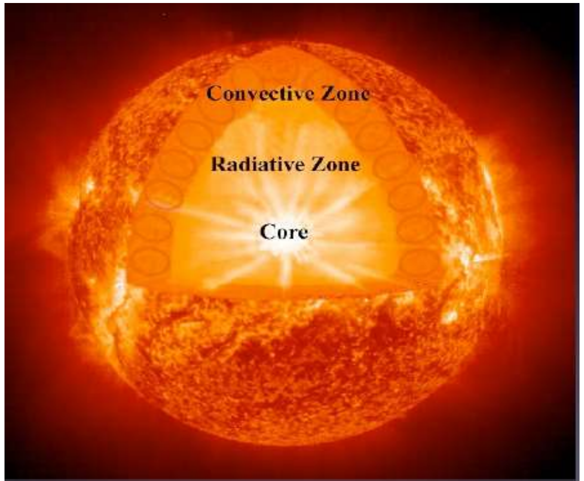


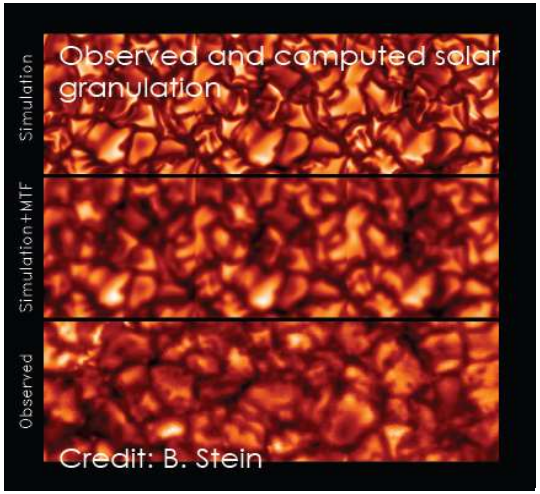
Image: SOHO (ESA/NASA)

# Surface Convection

Traditionally accounted for by rudimentary "mixing-length theory" (see Chap. 6) in 1-D atmosphere codes

BUT:
- Solar observations show very dynamic structure
- Granulation and lateral inhomogeneity

→ Need for full 3-D radiation-hydrodynamics simulations in which convective motions occur spontaneously if required conditions fulfilled (all physics of convection 'naturally' included)

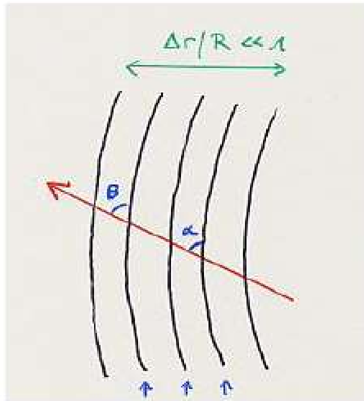



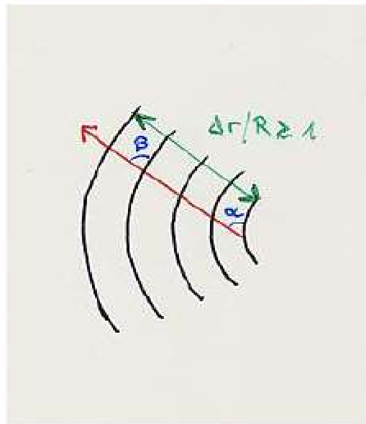
# Surface Convection

Solar-type stars:
Photospheric extent << stellar radius
Small granulation patterns

as long as $\Delta r / R << 1$ => plane-parallel symmetry

light ray through atmosphere

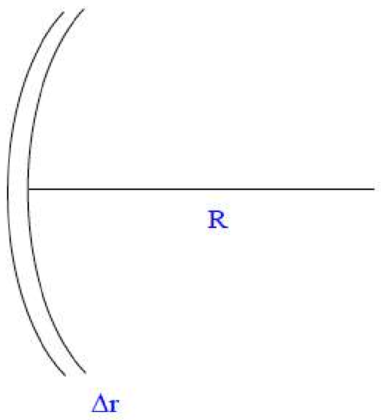



lines of constant temperature and density (isocontours)

curvature of atmosphere insignificant for photons' path : $\alpha = \beta$

significant curvature : $\alpha \neq \beta$, spherical symmetry




**example:** the sun

$R_{sun} \approx 700{,}000$ km
$\Delta r$ (photo) $\approx 300$ km

=> $\Delta r / R \approx 4\ 10^{-4}$

BUT corona
$\Delta r / R$ (corona) $\approx 3$

**examples**

| | |
|---|---|
| solar photosphere / cromosphere | solar corona |
| atmospheres of | atmospheres of |
| main sequence stars | supergiants |
| white dwarfs | expanding envelopes (stellar winds) |
| giants (partly) | of OBA stars, M-giants and supergiants |

from Chap. 3

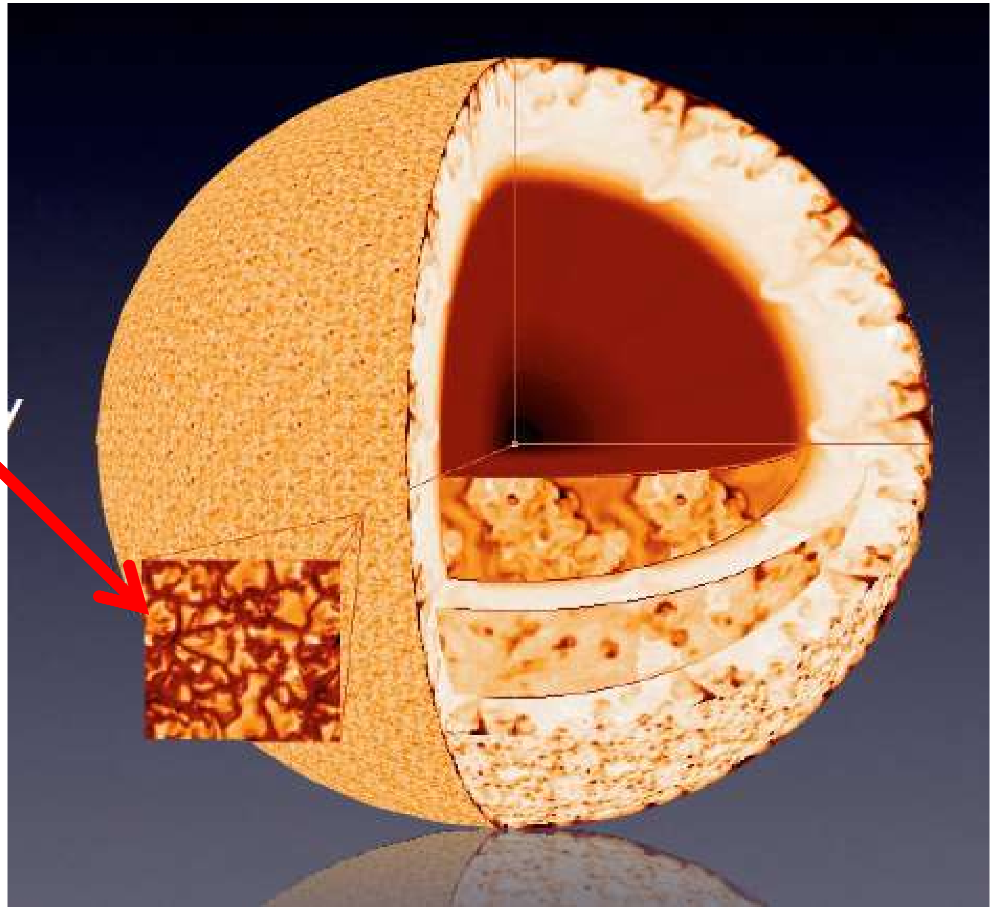
## Surface Convection

Solar-type stars:
Atmospheric extent << stellar radius
Small granulation patterns

→

# Box-in-a-star Simulations

(cmp. plane-parallel approximation)



From Wolfgang Hayek

## Surface Convection

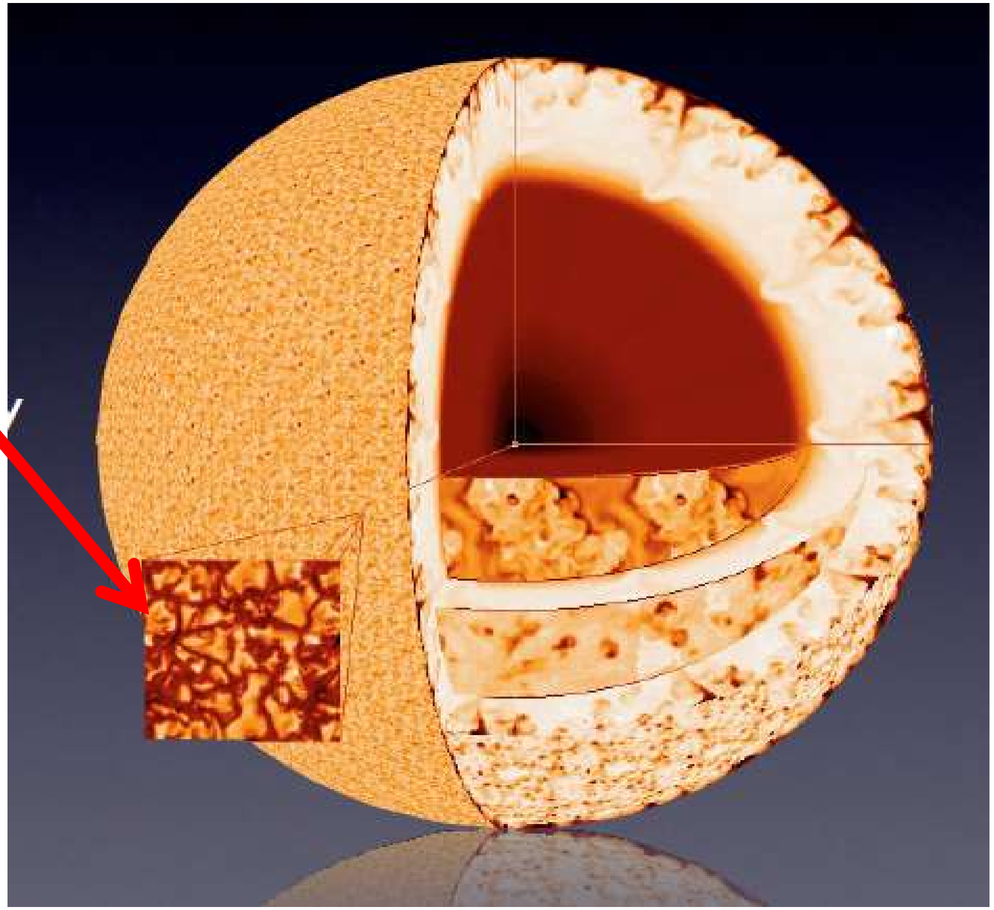
Approach
(teams by Nordlund, Steffen):

Solve radiation-hydrodynamical conservation equations of mass, momentum, and energy (closed by equation of state).

3-D radiative transfer included to calculate net radiative heating/cooling $q_{rad}$ in energy equation, typically assuming LTE and a very simplified treatment of line-blanketing

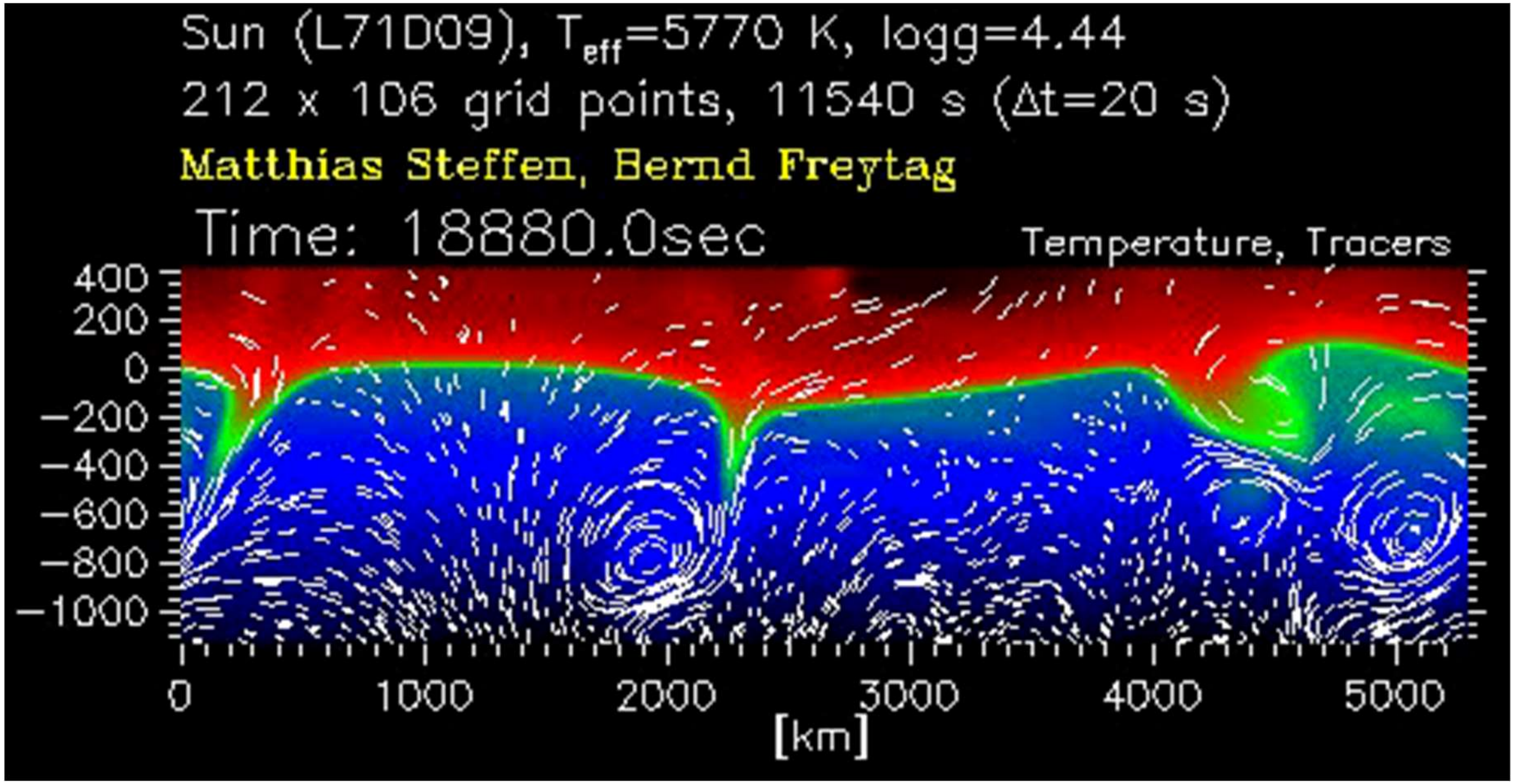
$$q_{\rm rad} = 4\pi\rho \int_\lambda \kappa_\lambda (J_\lambda - S_\lambda)\, d\lambda,$$

(= 0 in case of radiative equilibrium)



From Wolfgang Hayek

## Surface Convection




From Berndt Freytag's homepage: http://www.astro.uu.se/~bf/

# Surface Convection

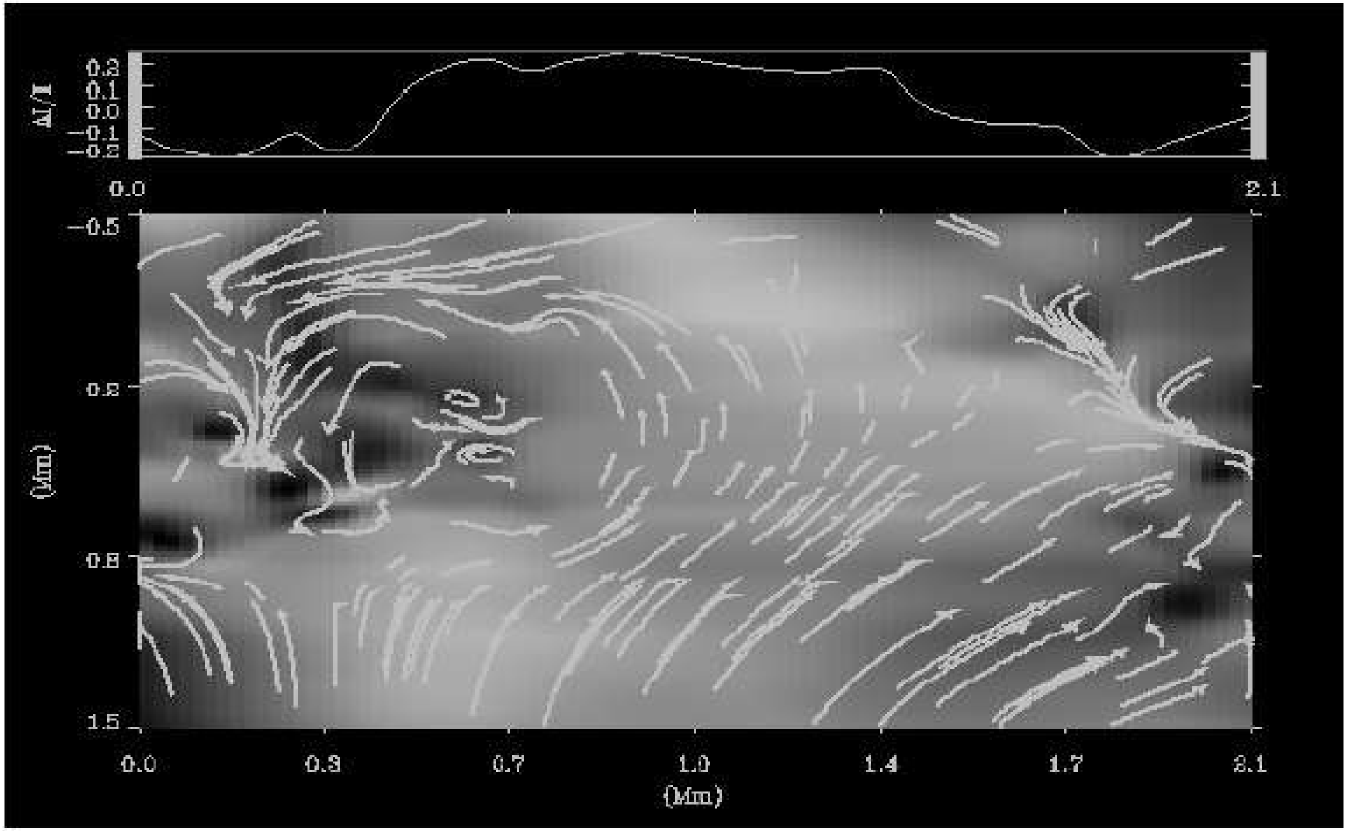



FIG. 4.—Pressure fluctuations about the mean hydrostatic equilibrium and the velocity field in an *xz* slice through a granule. The pressure is high above the centers of granules, which decelerates the warm upflowing fluid and diverts it horizontally. High pressure also occurs in the intergranular lanes where the horizontal motions are halted and gravity pulls the now cool, dense fluid down into the intergranular lanes. Horizontal rolls of high vorticity occur at the edges of the intergranular lanes. The emergent intensity profile across the slice is shown at the top.

From Stein & Nordlund (1998)

## Surface Convection

Some key features:

Slow, broad upward motions, and faster, thinner downward motions
Non-thermal velocity fields
Overshooting from zone where convection is efficient according to stability criteria (see Chap. 6)
Energy balance in upper layers not only controlled by radiative heating/cooling, but also by cooling from adiabatic expansion

See Stein & Nordlund (1998); Collet et al. (2006), etc.

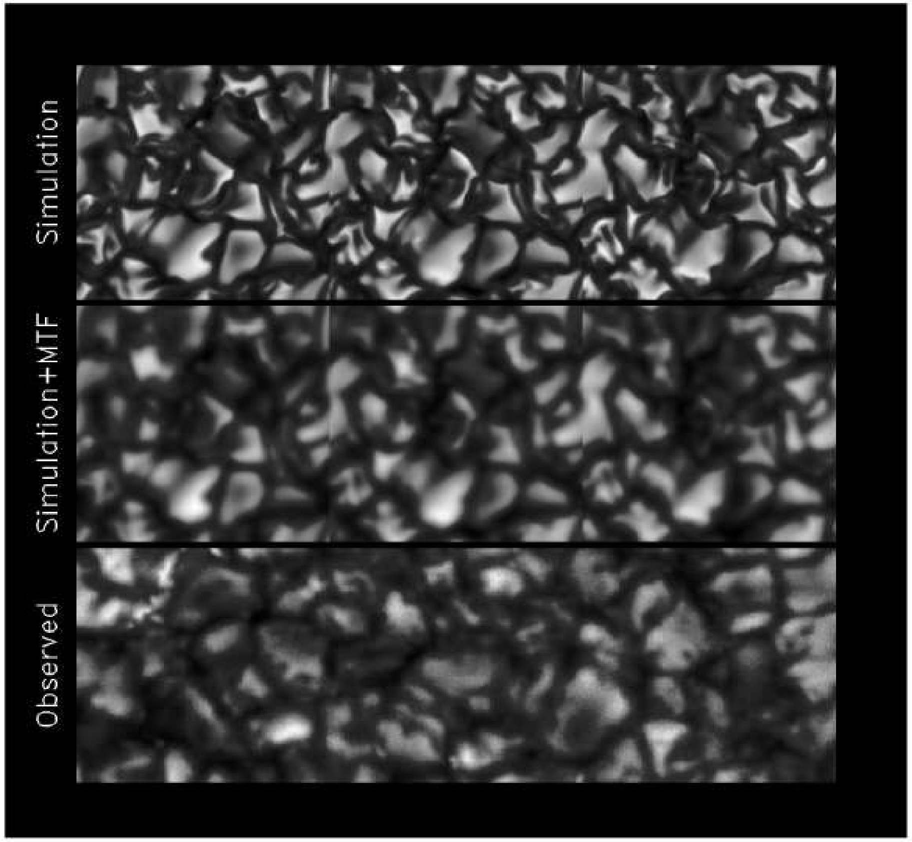



FIG. 19.—Comparison of granulation as seen in the emergent intensity from the simulations and as observed by the Swedish Vacuum Solar Telescope on La Palma. The top row shows three simulation images at 1 minute intervals, which together make a composite image 18 × 6 Mm in extent. The middle row shows this image smoothed by an Airy plus exponential point-spread function. The bottom row shows an 18 × 6 Mm white-light image from La Palma. Note the similar appearance of the smoothed simulation image and the observed granulation. The common edge brightening in the simulation is reduced when smoothed. Images by (Title 1996, private communication) taken in the CH G-band have much more contrast than white light and clearly reveal the edge brightening of granules.

Question: This does not look much like the traditional 1-D models we've discussed during the previous lecture! – Do you think we should throw them in the garbage?

## Surface Convection

blue: mean temperature from 3D hydro-model (scatter = dashed)
red: from 1D semi-empirical model (Holweger & Müller, see Chap. 5)
green: from 1D theoretical model atmospheres (MARCS)

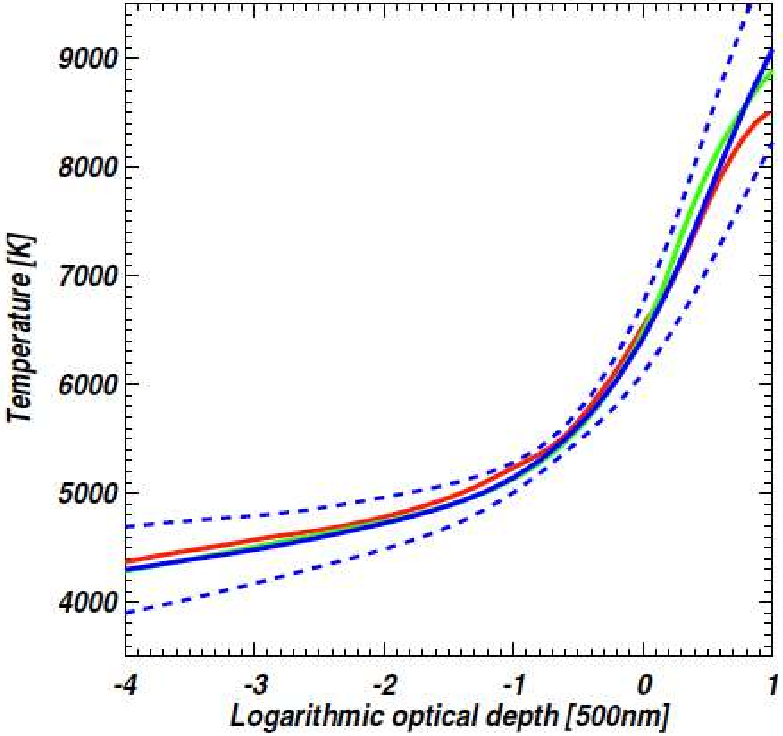



Figure 1: The mean temperature structure of the 3D hydrodynamical model of Trampedach et al. (2009) is shown as a function of optical depth at 500 nm (blue solid line). The blue dashed lines correspond to the spatial and temporal rms variations of the 3D model, while the red and green curves denote the 1D semi-empirical Holweger & Müller (1974) and the 1D theoretical MARCS (Gustafsson et al. 2008) model atmospheres, respectively.

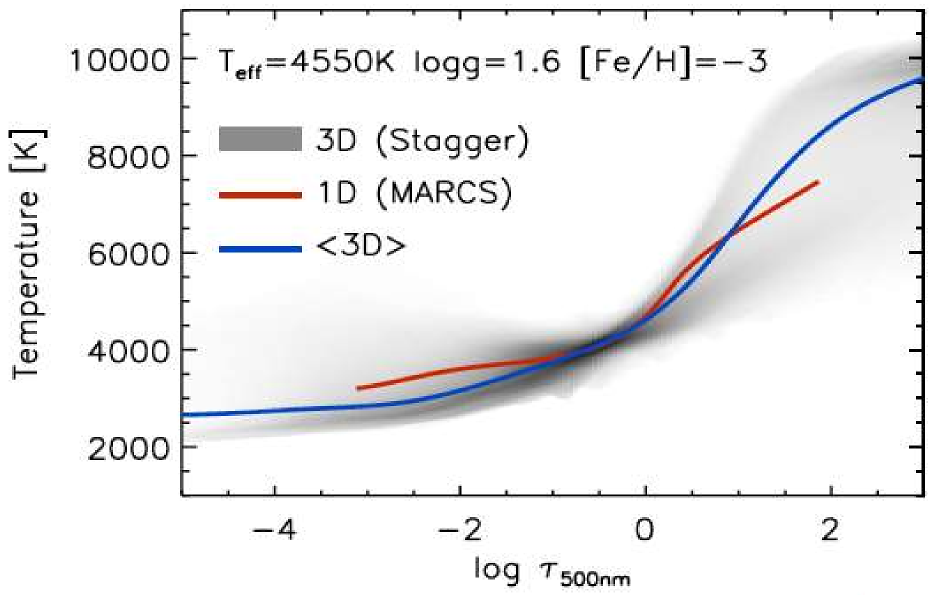
In many (though not all) cases, AVERAGE properties still quite OK:

Convection in energy balance approximated by "mixing-length theory"
Non-thermal velocity fields due to convective motions included by means of so-called "micro-" and "macro-turbulence"

BUT quantitatively we always need to ask:
To what extent can average properties be modeled by traditional 1-D codes?

Unfortunately, a general answer very difficult to give, need to be considered case by case

# Surface Convection



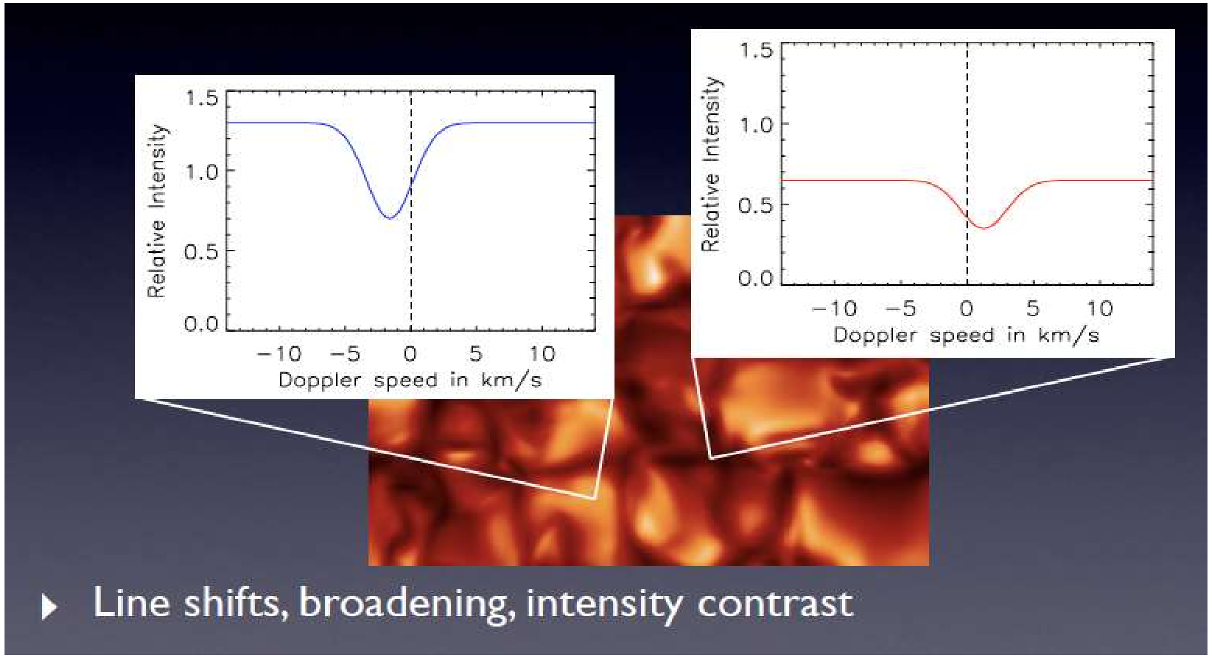

Metal-poor red giant, simulation by Remo Collet,
figure from talk by M. Bergemann

For example:

In metal-poor cool stars spectral lines are scarce (Question: Why?), and energy balance in upper photosphere controlled to a higher degree by adiabatic expansion of convectively overshot material.

In classical 1-D models though, these layers are convectively stable, and energy balance controlled only by radiation (radiative equilibrium, see part1).

## Surface Convection




From talk by Hayek

3-D radiation-hydro models successful in reproducing many solar features (see overview in Asplund et al. 2009), e.g:
Center-to-limb intensity variation
Line profiles and their shifts and variations (without micro/macroturbulence)
Observed granulation patterns

## Surface Convection

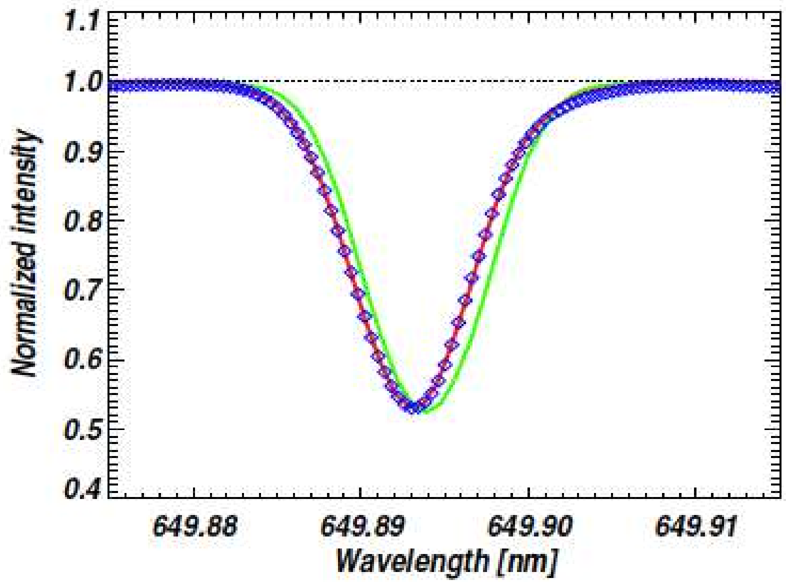



Figure 3: The predicted spectral line profile of a typical Fe I line from the 3D hydrodynamical solar model (red solid line) compared with the observations (blue rhombs). The agreement is clearly very satisfactory, which is the result of the Doppler shifts arising from the self-consistently computed convective motions that broaden, shift and skew the theoretical profile. For comparison purposes also the predicted profile from a 1D model atmosphere (here Holweger & Müller 1974) is shown; the 1D profile has been computed with a microturbulence of 1 km $s^{-1}$ and a tuned macroturbulence to obtain the right overall linewidth. Note that even with these two free parameters the 1D profile can neither predict the shift nor the asymmetry of the line.

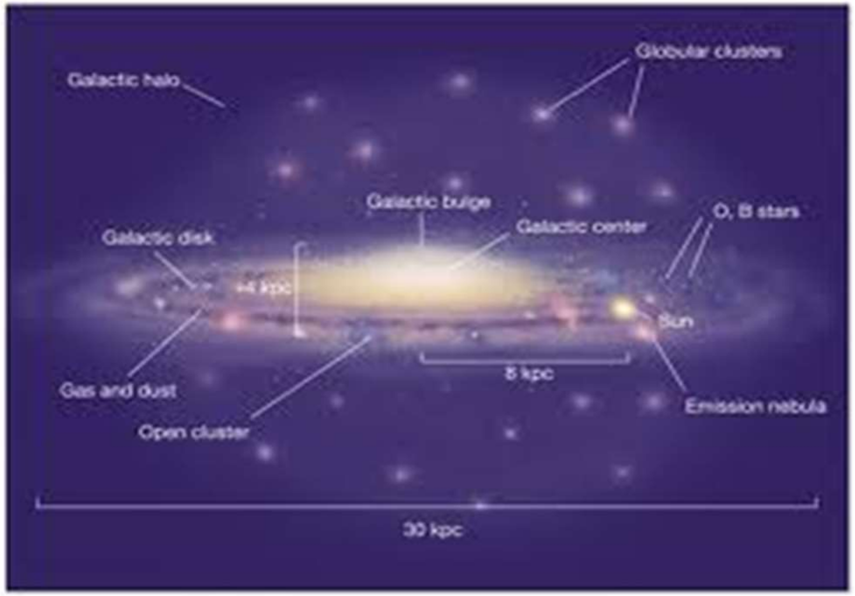
affects chemical abundance (determined by means of line profile fitting to observations)

One MAJOR result:
Effects on line formation has led to a downward revision of the CNO solar abundances and the solar metallicity, and thus to a revision of the *standard cosmic chemical abundance scale*

Fig. from Asplund et al. (2009) – "The Chemical Composition of the Sun"

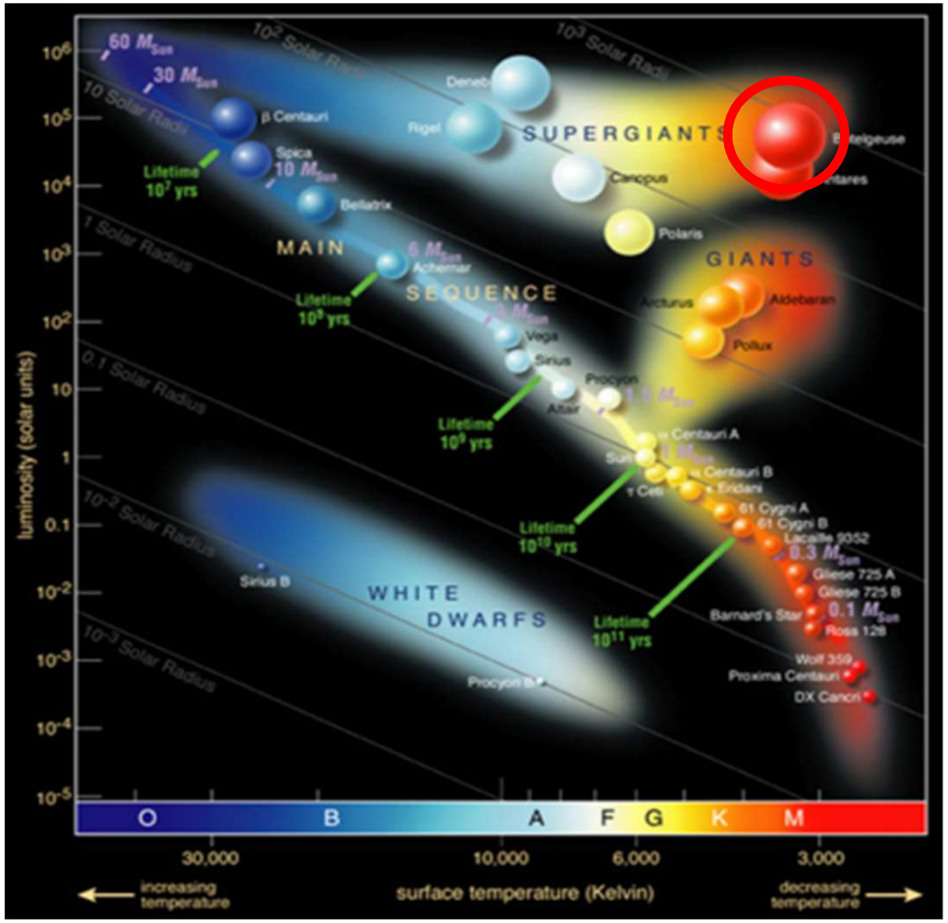
## Surface Convection

Also potentially critical for Galactic archeology…




…which traces the chemical evolution of the Universe by analyzing VERY old, metal-poor Globular Cluster stars –- relics from the early epochs (e.g. Anna Frebel and collaborators)

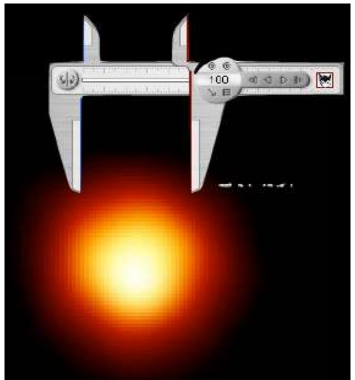
## Surface Convection




- Giant Convection Cells in the low-gravity, extended atmospheres of Red Supergiants
- Question: Why extended?

$H = a^2 / g$ (with $a = v_s$ the isothermal speed of sound)

Out to Jupiter…


Betelgeuse (HST)
Gilliland & Dupree 1996

$$a^2_{RSG} / a^2_{sun} \approx T_{RSG} / T_{sun} = 0.5...0.6$$

$$g_{RSG} / g_{sun} \approx 10^{-4}\,!$$

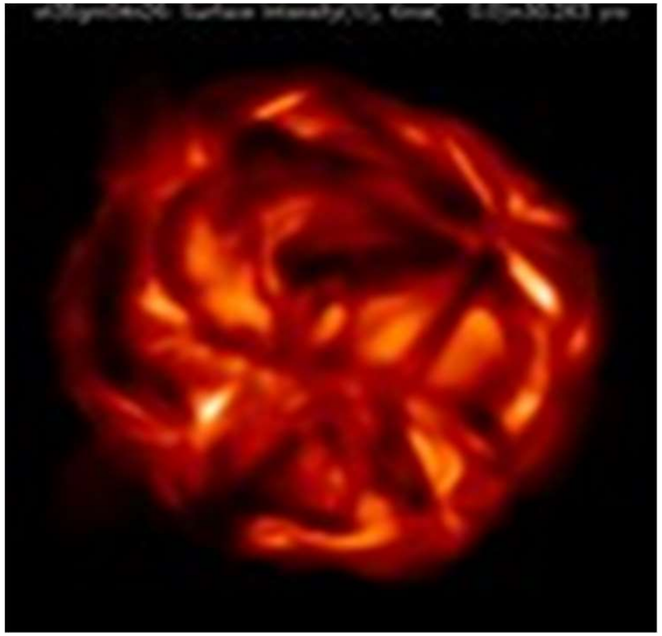
(see Chap. 6)

# Surface Convection

Supergiants (or models including a stellar wind):
Atmospheric extent > stellar radius:

Box-in-a-star → Star-in-a-box

(1D: Plane-parallel → Spherical symmetry, see Chap. 3)

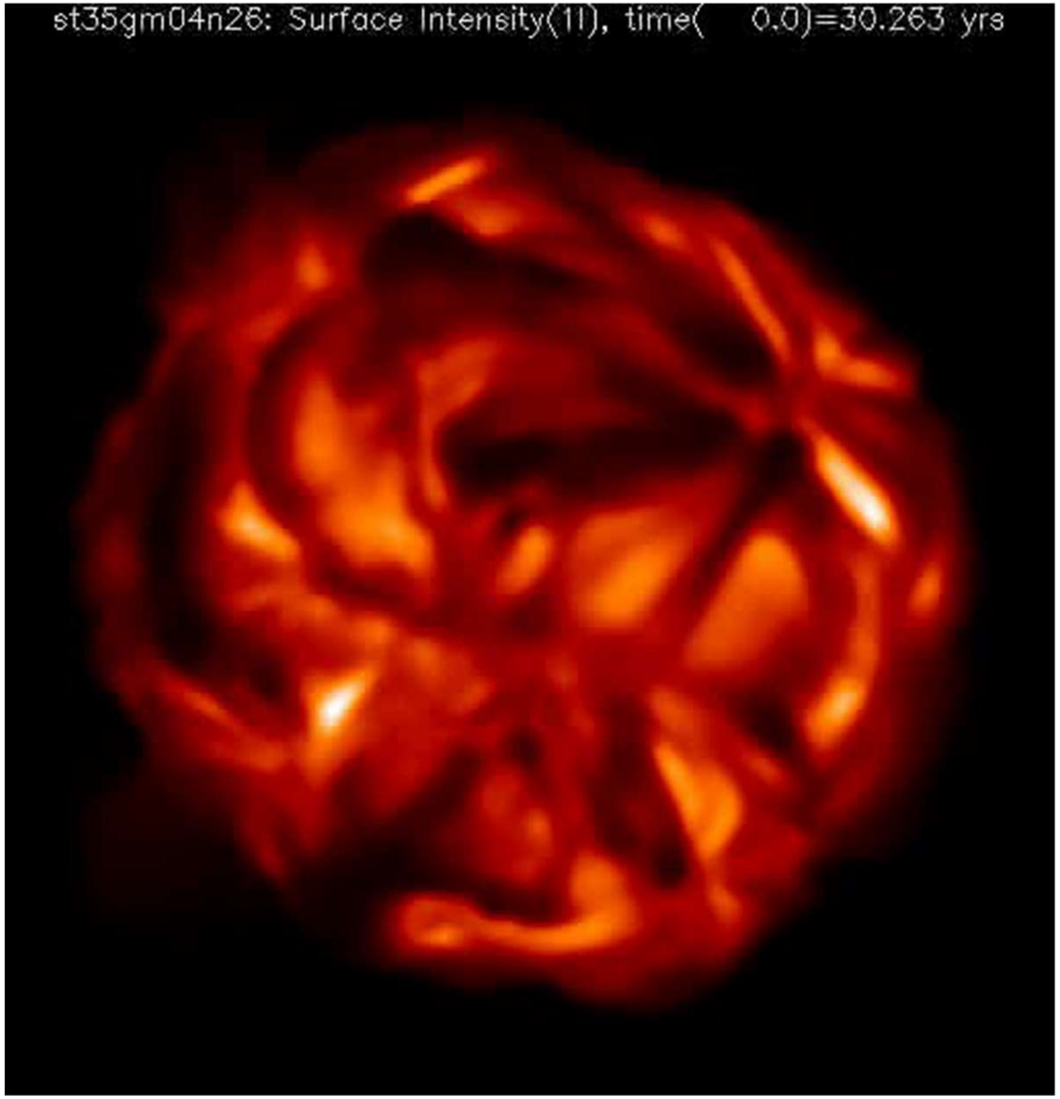


Star to model: Betelgeuse
Mass: 5 solar masses
Radius: 600 Rsun
Luminosity: 41400 Lsun
Grid: Cartesian cubical grid with $171^3$ points
Edge length of box 1674 solar radii

Model by Berndt Freytag, note the HUGE convective cells visible in the emergent intensity map!!

## Surface Convection

Star to model: Betelgeuse
Mass: 5 solar masses
Radius: 600 Rsun
Luminosity: 41400 Lsun
Grid: Cartesian cubical grid with $171^3$ points
Edge length of box 1674 solar radii
Movie time span: 7.5 years



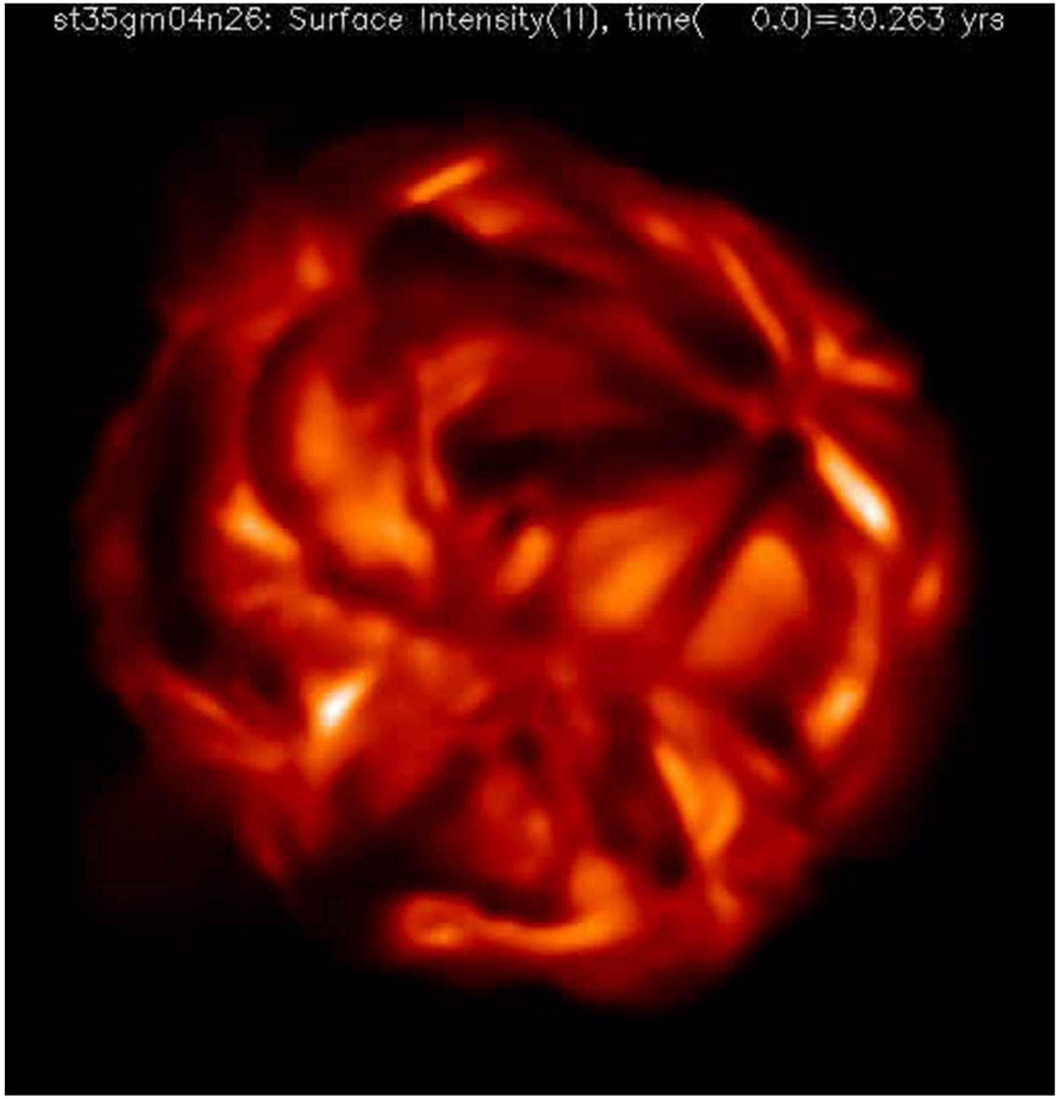

http://www.astro.uu.se/~bf/movie/dst35gm04n26/movie.html

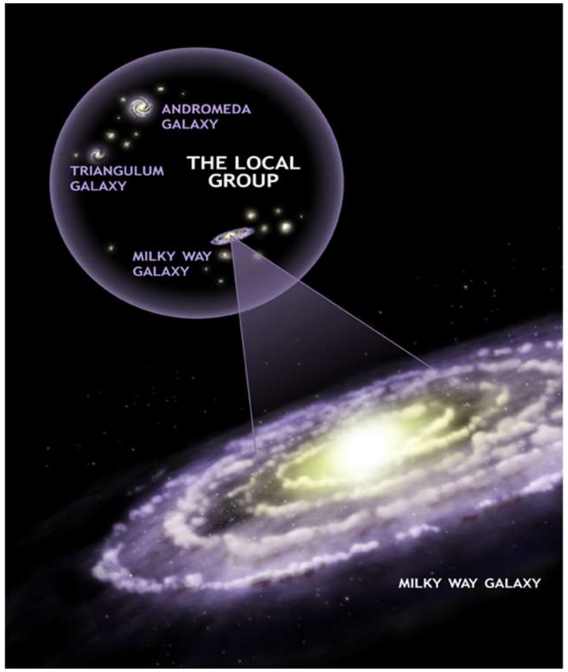
## Surface Convection

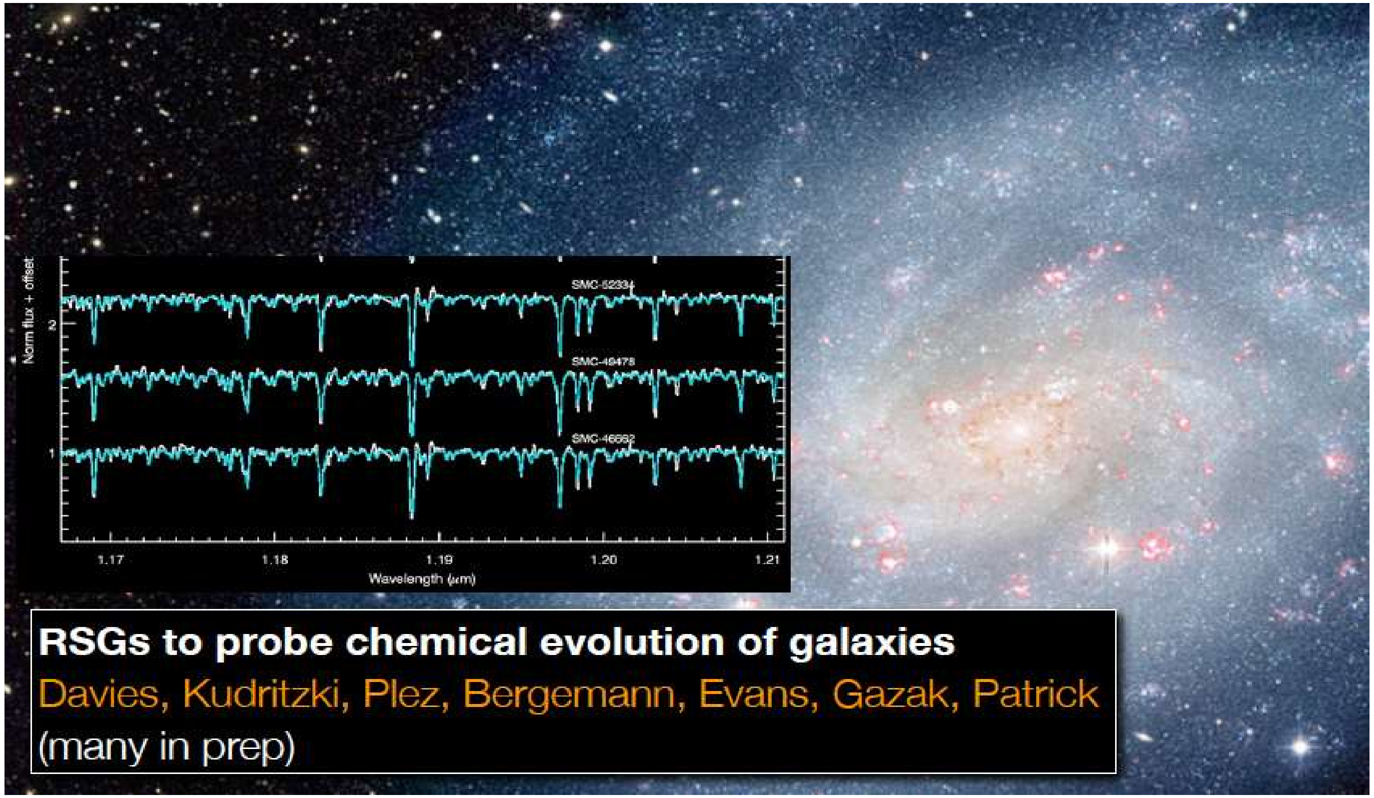
Extremely challenging,
models still in their infancies.
LOTS of exciting physics to explore, like

Pulsations
Convection
Numerical radiation-hydrodynamics
Role of magnetic fields
Stellar wind mechanisms

Also, to what extent can main effects be captured by 1-D models?
For quantitative applications like….





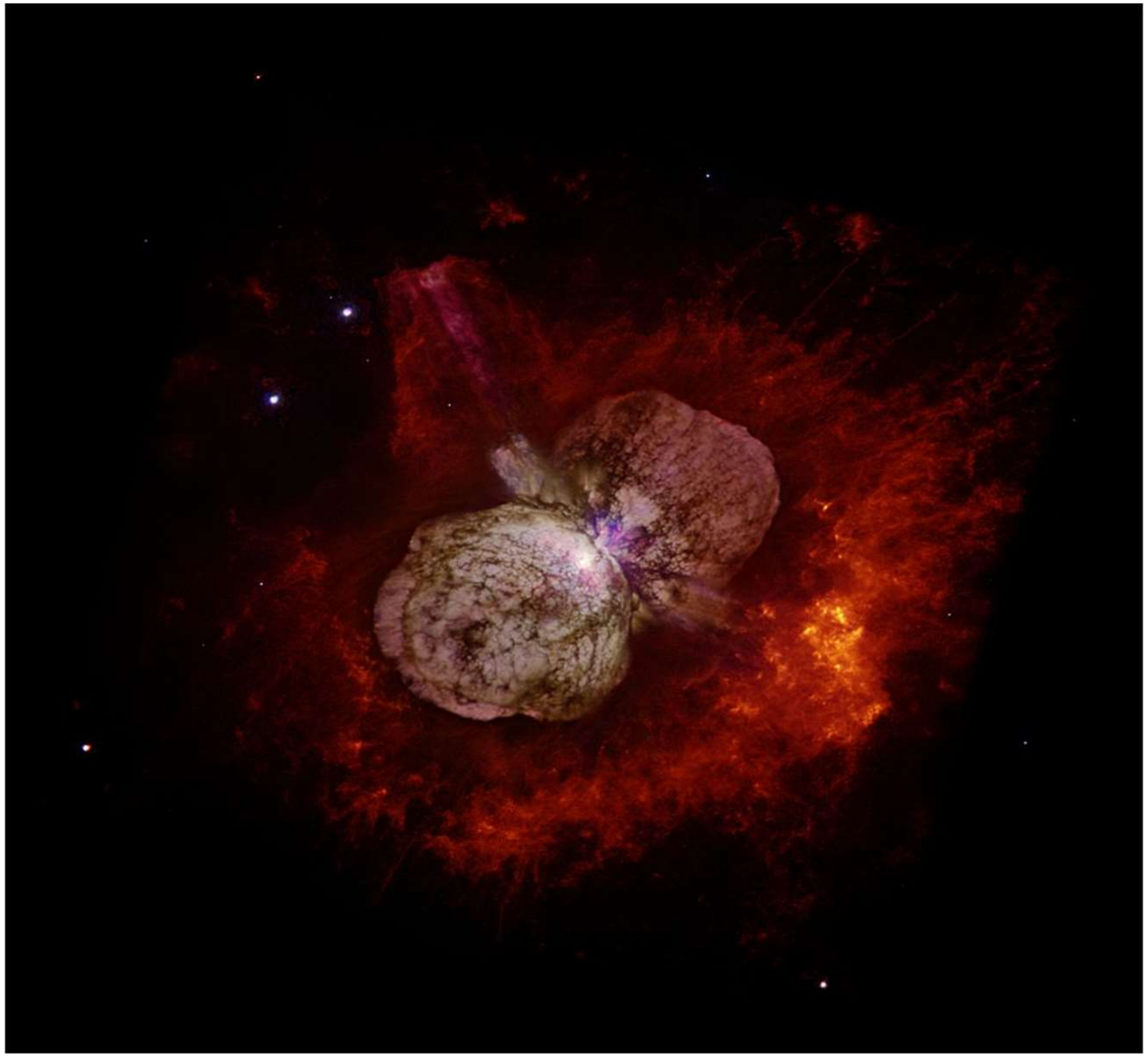

Question: Why are RSGs ideal for observational extragalactic stellar astrophysics, particularly in the near future?

# important codes and their features ....

| Codes | FASTWIND<br>CMFGEN<br>PoWR | WM-basic | TLUSTY<br>Detail/Surface | Phoenix | MARCS<br>Atlas | CO$^5$BOLD*<br>STAGGER |
|---|---|---|---|---|---|---|
| geometry | 1-D<br>spherical | 1-D<br>spherical | 1-D<br>plane-parallel | 1-D/3-D<br>spherical/<br>plane-parallel | 1-D<br>plane-parallel<br>(MARCS also spherical) | 3-D<br>Cartesian |
| LTE/NLTE | NLTE | NLTE | NLTE | NLTE/LTE | LTE | LTE simplified |
| dynamics | quasi-static photosphere + prescribed supersonic outflow | time-independent hydrodynamics | hydrostatic | hydrostatic or allowing for supersonic outflows | hydrostatic | hydrodynamic |
| stellar wind | yes | yes | no | yes | no | no |
| major application | hot stars with winds | hot stars with dense winds, ion. fluxes, SNRs | hot stars with negligible winds | cool stars, brown dwarfs, SNRs | cool stars | cool stars |
| comments | CMFGEN also for SNRs; FASTWIND using approx. line-blocking | line-transfer in Sobolev approx. (see part 2) | Detail/Surface with LTE-blanketing | convection via mixing-length theory | convection via mixing-length theory | very long execution times, but model grids start to emerge |

* COnservative COde for the COmputation of COmpressible COnvection

And then there are, e.g.,

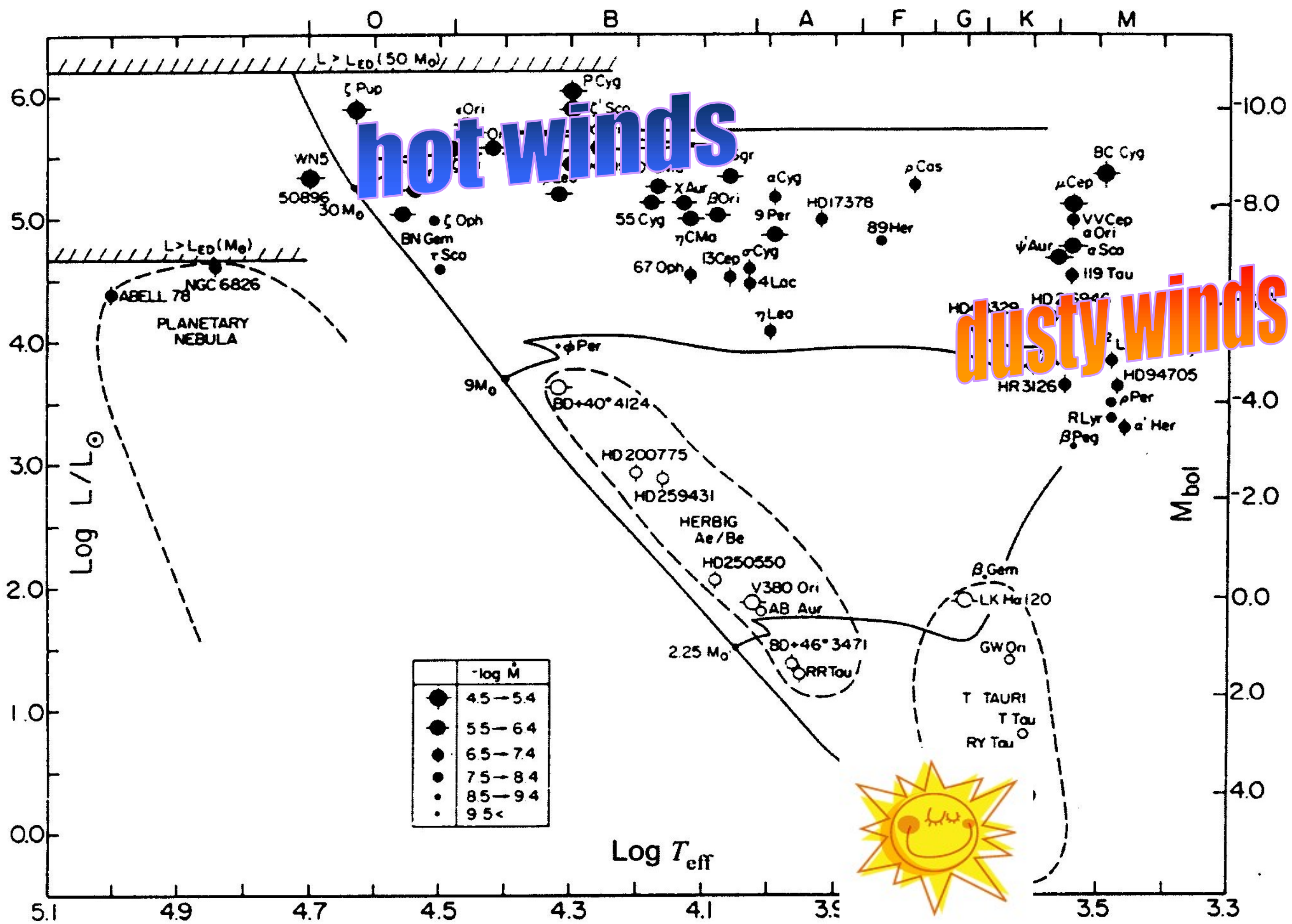


- Luminous Blue Variables (LBVs) like Eta Carina,
- Wolf-Rayet Stars (WRs)
- Planetary Nebulae (and their Central Stars)
- Be-stars with disks
- Brown Dwarfs
- Pre main-sequence T-Tauri and Herbig stars

…and many other interesting objects

Stellar astronomy alive and kicking! Very rich in both

Physics
Observational applications

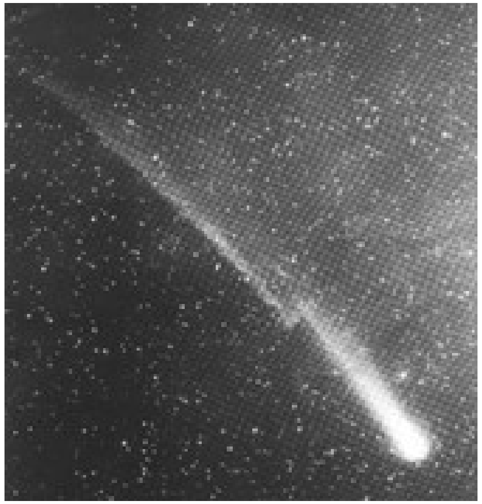


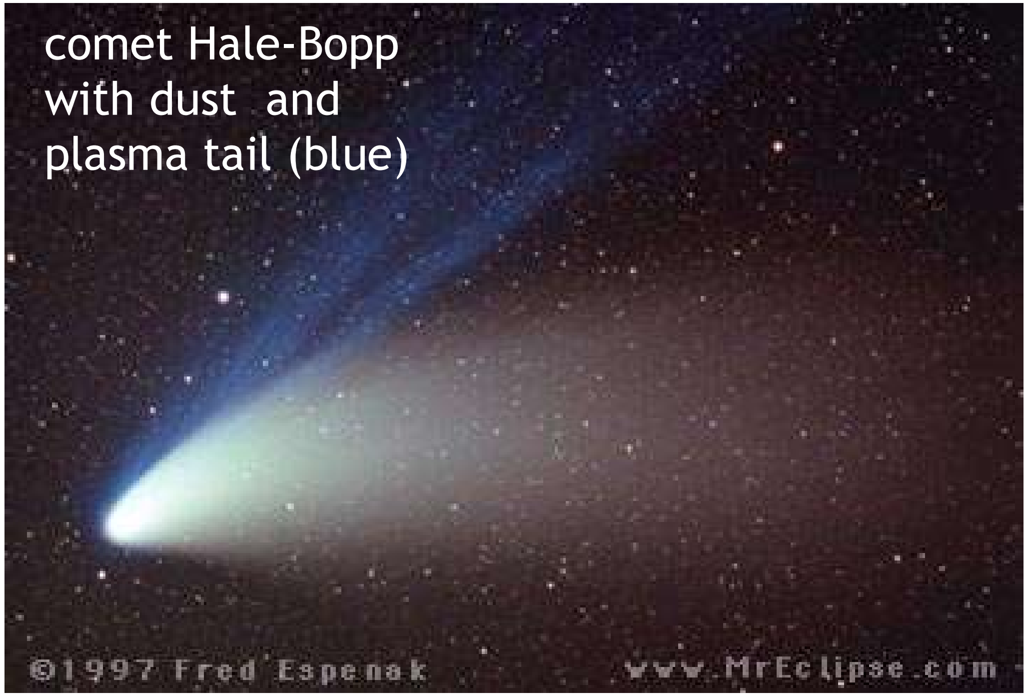

**ubiquitous phenomenon**

- solar type stars (incl. the sun)
- red supergiants/AGB-stars ("normal" + Mira Variables)
- hot stars (OBA supergiants, Luminous Blue Variables, OB-dwarfs, Central Stars of PN, sdO, sdB, Wolf-Rayet stars)
- T-Tauri stars
- and many more

# The solar wind – a suspicion

comet Halley, with „kink“ in tail



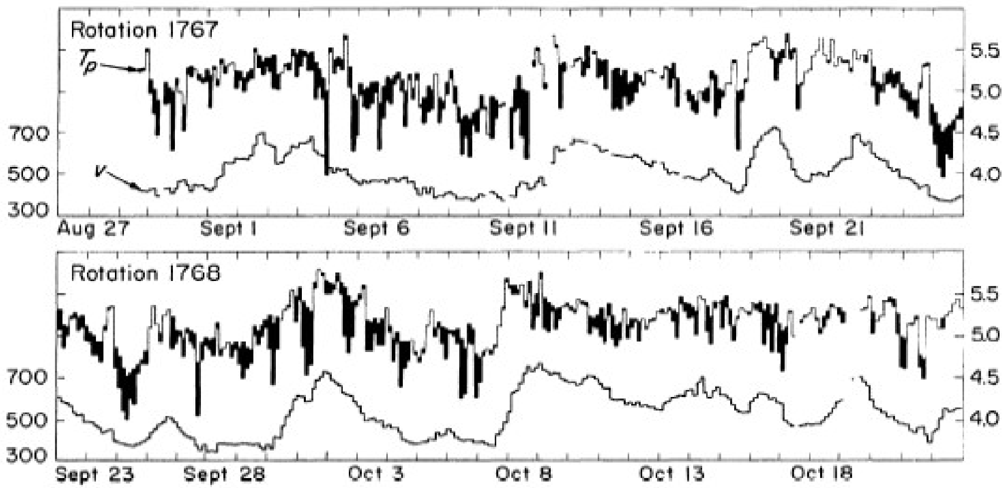



comet Hale-Bopp with dust and plasma tail (blue)

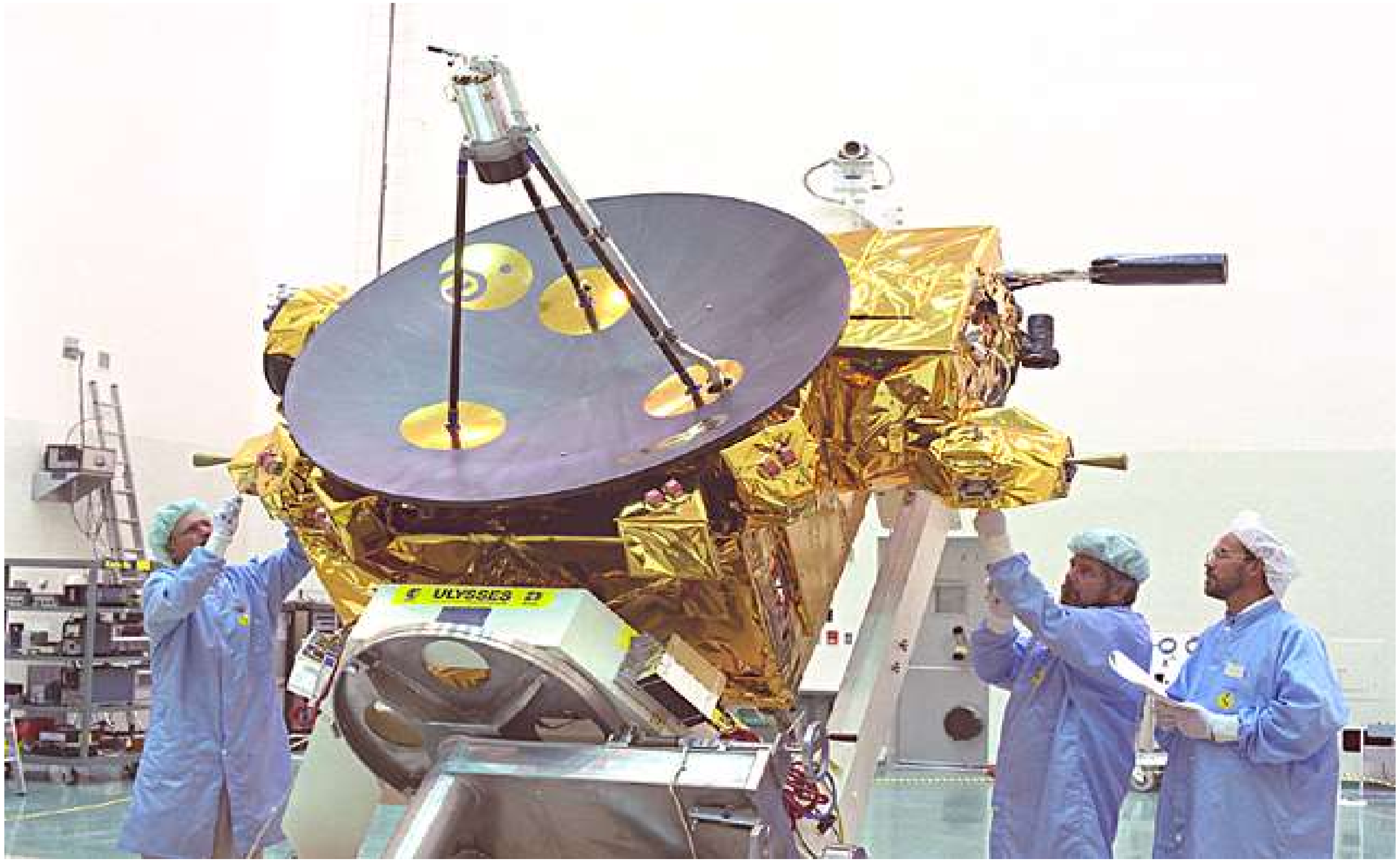
- comet tails *directed away from the sun*
- **Kepler:** influence of solar radiation pressure (-> radiation driven winds)
- *Ionic tail*: emits own radiation, sometimes different direction
- **Hoffmeister** (1943, subsequently Biermann): *solar particle radiation* different direction, since **v** (particle) comparable to **v** (comet)

# The solar wind – the discovery

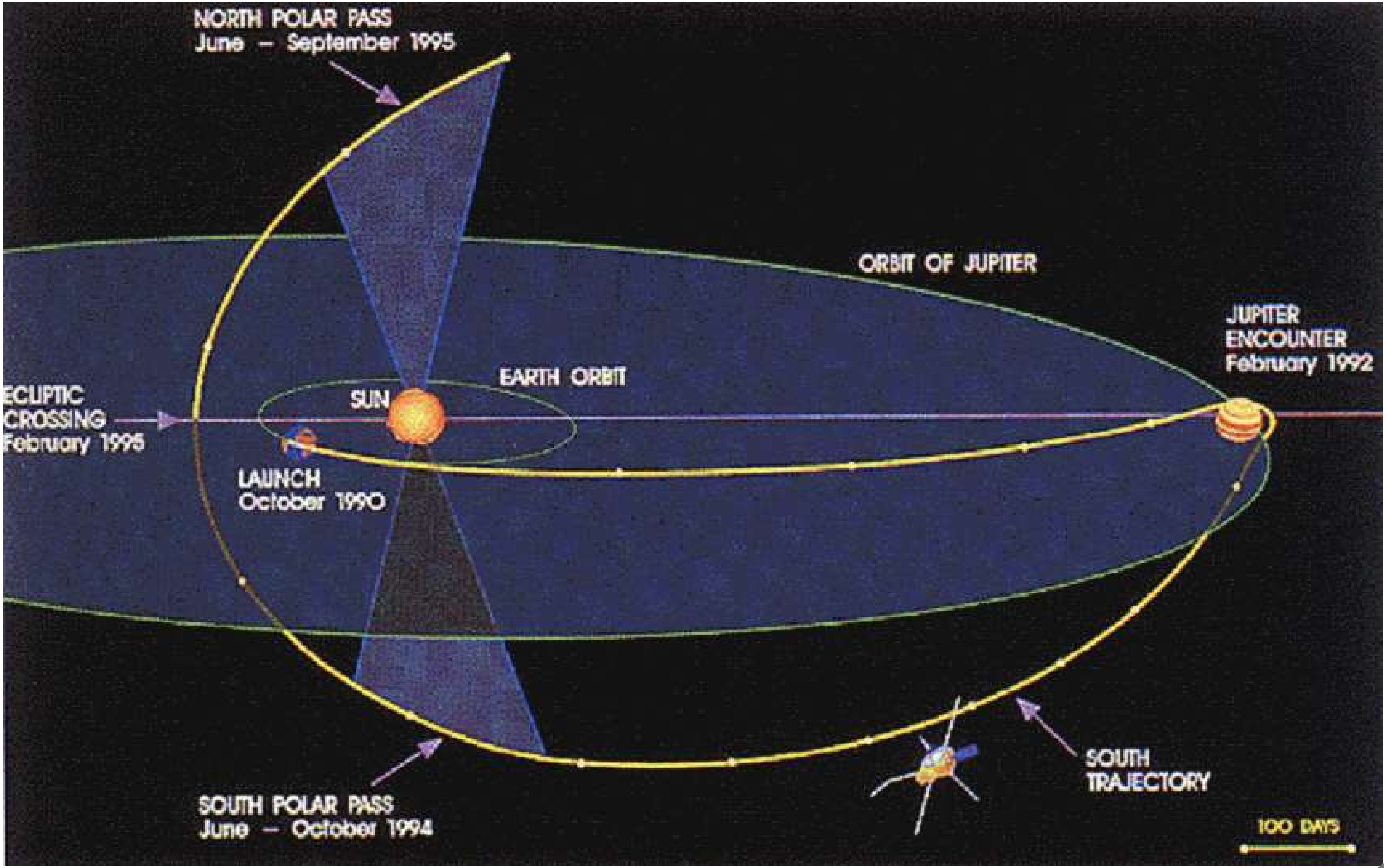
- **Eugene Parker (1958):** theoretical(!) investigation of coronal equilibrium: high temperature leads to (solar) wind (more detailed later on)
- **confirmed by**
  - Soviet measurements (Lunik2/3) with “ion-traps” (1959)
  - Explorer 10 (1961)
  - Mariner II (1962): measurement of fast and slow flows (27 day cycle -> co-rotating, related “coronal holes” and sun spots)

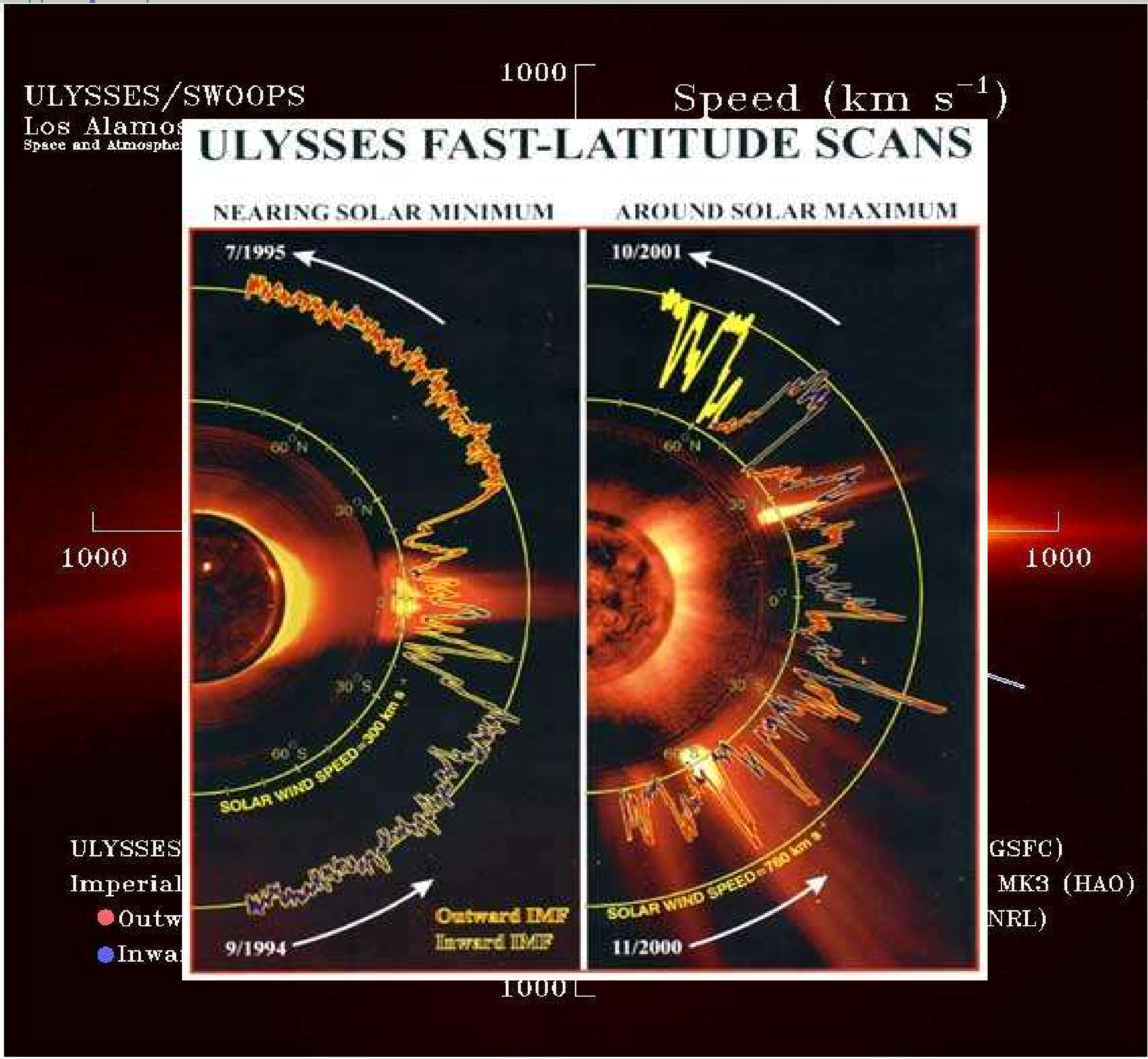


# The solar wind – Ulysses ...



# … surveying the polar regions

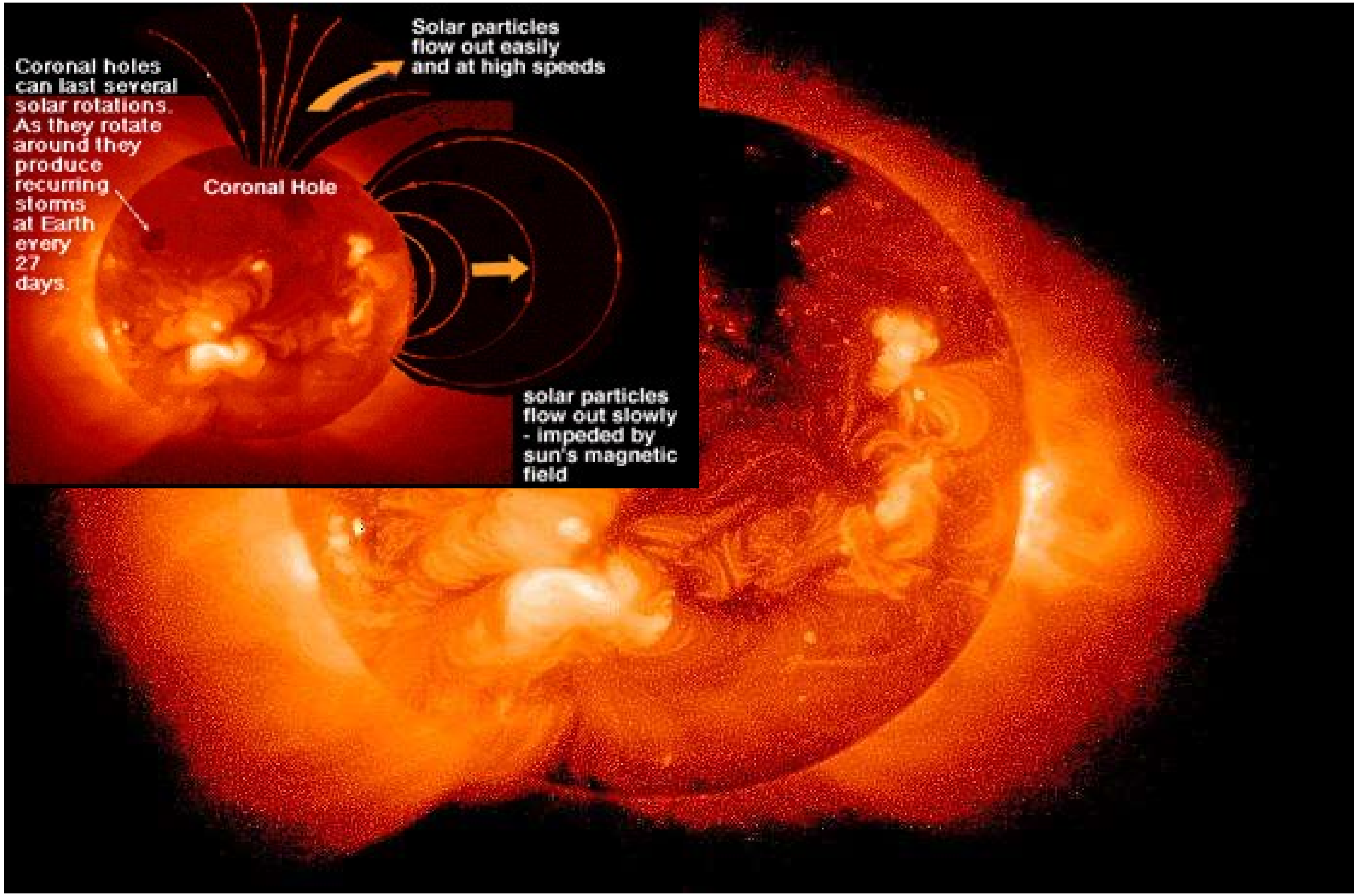





*polar* wind:
fast and thin

*equatorial* wind:
slow and dense

# The solar wind – coronal holes




*fast wind:*
over coronal holes (dark corona, “open” field lines, e.g., in polar regions)

coronal X-ray emission

⇒

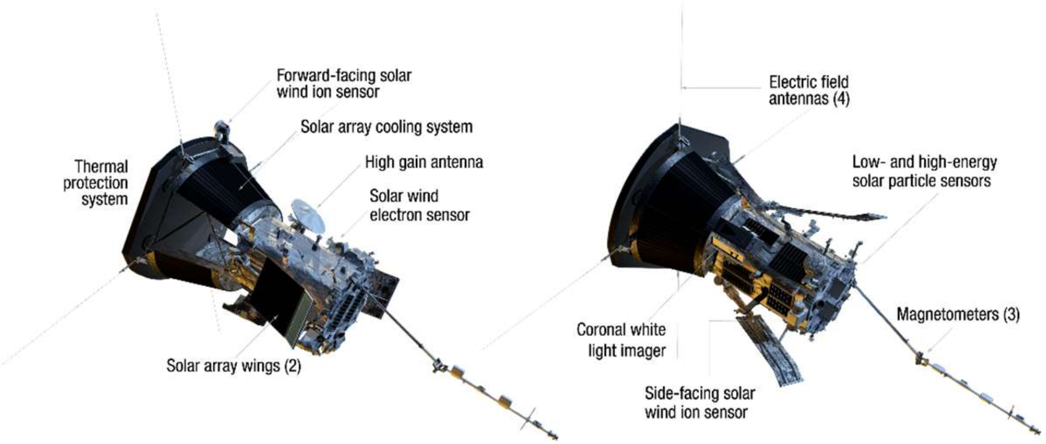
**very high temperatures**

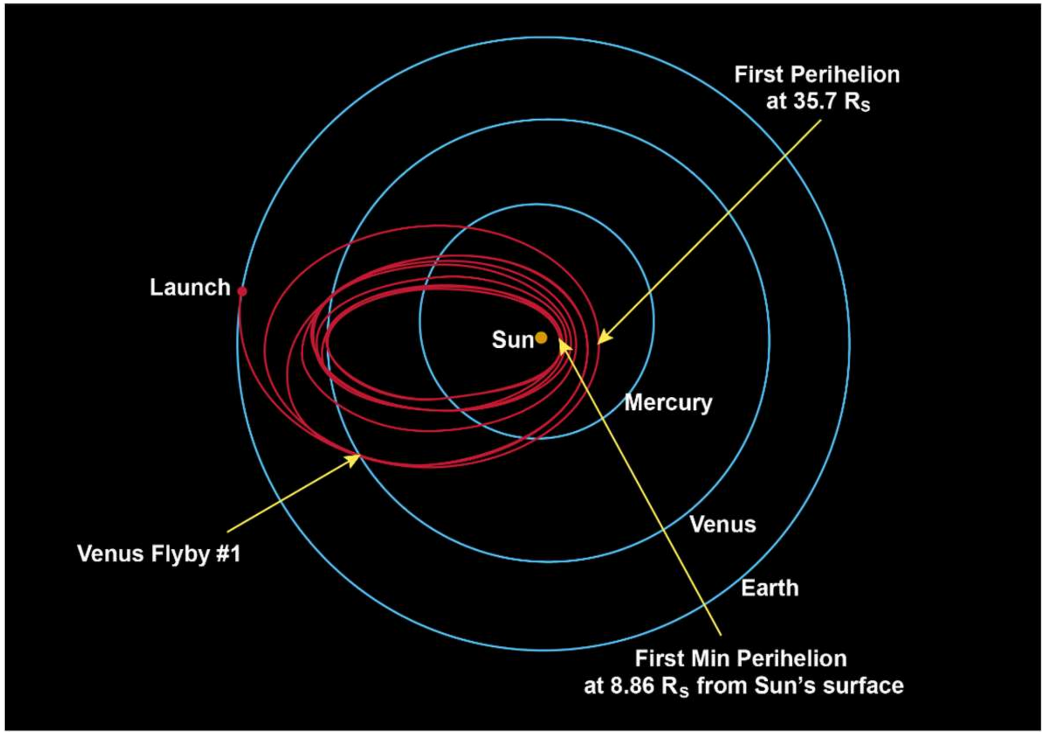
(Yohkoh Mission)

# Parker Solar Probe

## Primary objectives for the mission

- trace the energy flow, understand heating of the solar corona, study the outer corona.
- determine the structure and dynamics of the plasma and magnetic fields
- explore solar wind driving, and mechanisms that accelerate and transport energetic particles.

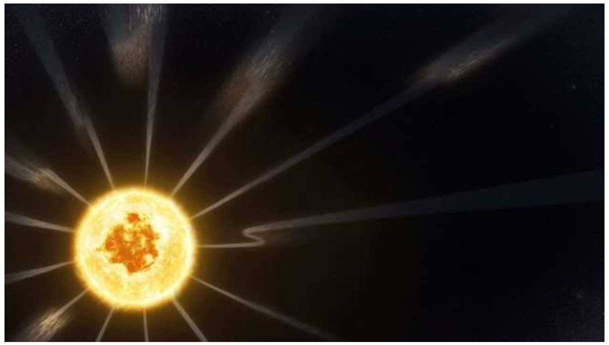



planned: 24 orbits, first perihelion on Nov. 5, 2018;
seven Venus-flybys over 7 years, to decrease perihelion distance from 36 to 8.9 $R_{sun}$ (6 Millionen km, with T~1100 K)

## First results (Nov. 2019)

- wind rotates, but up to 10 times faster than expected
- high speed plasma waves, up to c/6, can revert direction of B-field → "switchbacks": coherent (wind) structures
- coronal mass ejections much more irregular than expected
- dust cleared by solar wind



Credit: NASA's Goddard Space Flight Center/Conceptual Image Lab/Adriana Manrique Gutierrez

all material from: parkersolarprobe.jhuapl.edu

# The sun and its wind: mean properties

## The sun

radius = 695,990 km = 109 terrestrial radii
mass = 1.989 $10^{30}$ kg = 333,000 terrestrial masses
luminosity = 3.85 $10^{33}$ erg/s = 3.85 $10^{20}$ MW $\approx$ $10^{18}$ nuclear power plants
effective temperature = 5770 °K
central temperature = 15,600,000 °K
life time approx. 10 $10^9$ years
age = 4.57 $10^9$ years
distance sun earth approx. 150 $10^6$ km $\approx$ 400 times earth-moon

## The solar wind

temperature when leaving the corona: approx.1 $10^6$ K
average speed approx. 400-500 km/s (travel time sun-earth approx. 4 days)
particle density close to earth: approx. 6 $cm^{-3}$
temperature close to earth: $\lesssim 10^5$ K
*mass-loss rate*: approx $10^{12}$ g/s (1 Megaton/s) $\approx$ $10^{-14}$ solar masses/year
$\approx$ one Great-Salt-Lake-mass/day $\approx$ one Baltic-sea-mass/year
$\Rightarrow$ no consequence for solar evolution, since only 0.01% of total mass lost over total life time

**Need mechanism which accelerates material beyond escape velocity:**

- **pressure** driven winds
- **radiation** driven winds

**Note:** red giant winds still not understood,
only scaling relations available ("Reimers-formula")

remember equation of motion (conservation of momentum + stationarity, cf. Chap. 6, page 90)

$$v\frac{dv}{dr} = -\frac{1}{\rho}\frac{dp}{dr} + g^{ext} \quad \text{(in spherical symmetry), and } p = \rho a^2 \quad \text{(equation of state, with isothermal sound-speed } a\text{)}$$

$\Rightarrow$ with mass-loss rate $\dot{M}$, radius $r$, density $\rho$ and velocity $v$

$$\dot{M} = 4\pi r^2 \rho v,$$

**equation of continuity:**
conservation of mass

$$\left(1-\frac{a^2}{v^2}\right)v\frac{dv}{dr} = -\frac{GM}{r^2} + g_{rad} + \frac{2a^2}{r} - \frac{da^2}{dr}$$

**equation of motion:**
from conservation of momentum

| | | | |
|---|---|---|---|
| vel. field | grav. accel. | radiative accel. | (part of) accel. by pressure gradient |
| **positive** for v > a<br>**negative** for v < a | inwards | outwards | outwards |

# Pressure driven winds

$$\left(1-\frac{a^2}{v^2}\right)v\frac{dv}{dr} = -\frac{GM}{r^2} + \cancel{g_{rad}} + \frac{2a^2}{r} - \cancel{\frac{da^2}{dr}}$$

vel. field — grav. accel. — radiative accel. — "pressure"

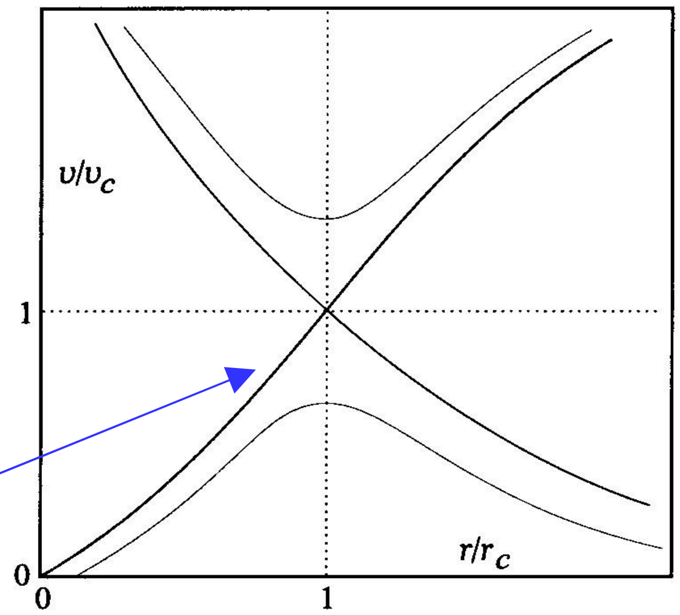
## The solar wind as a proto-type for pressure driven winds

- present in stars which have an (extremely) hot corona ($T \approx 10^6$ K)
- with $g_{rad} \approx 0$ and $T \approx$ const, the rhs of the equation of motion changes sign at

  $r_c = \dfrac{GM}{2a^2}$; with a (T=$1.5 \cdot 10^6$ K) $\approx 160$ km/s,

  we find for the sun $r_c \approx 3.9 R_{sun}$

  and obtain four possible solutions for $v/v_c$ ("c" = critical point)




- only one (the "transonic") solution compatible with observations
- pressure driven winds as described here rely on the presence of a hot corona (large value of a!)
- Mass-loss rate $\dot{M} \approx 10^{-14}$ $M_{sun}$ / yr, terminal velocity $v_\infty \approx 500$ km/s
- has to be heated (dissipation of acoustic and magneto-hydrodynamic waves)
- not completely understood so far

# Radiation driven winds

accelerated by radiation pressure:

$$\left(1-\frac{a^2}{v^2}\right)v\frac{dv}{dr} = -\frac{GM}{r^2} + g_{rad} + \underbrace{\frac{2a^2}{r}-\frac{da^2}{dr}}_{\text{important only in lowermost wind}}$$

pressure terms only of secondary order
(a ≈ 20 km/s for hot stars,
≈ 3 km/s for cool stars)

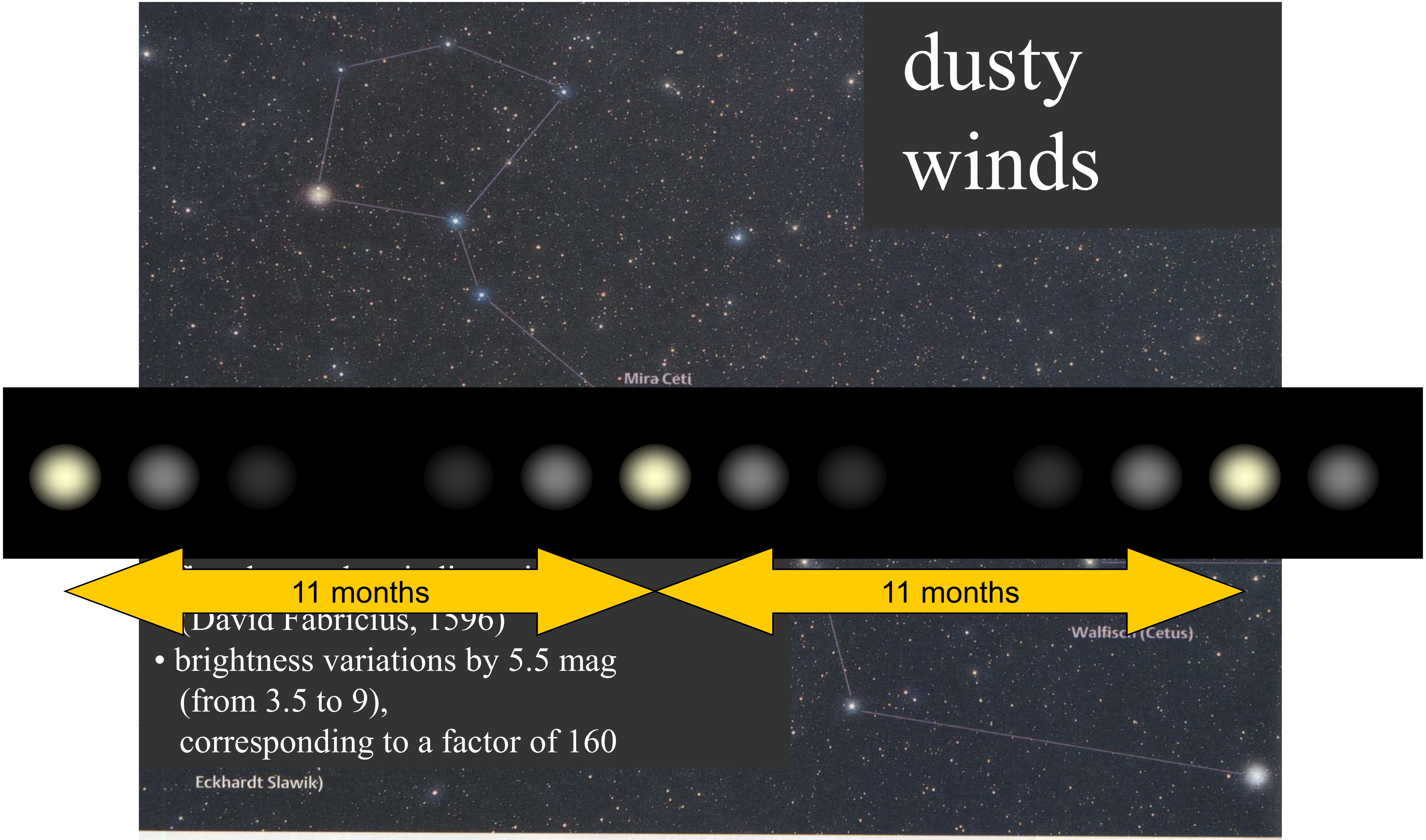
- cool stars (AGB): major contribution from **dust** absorption;
  coupling to "gas" by viscous drag force (gas – grain collisions)

  $\dot{M} \approx 10^{-6} M_{sun}/\text{yr}, \; v_\infty \approx 20$ km/s

- hot stars: major contribution from **metal line** absorption;
  coupling to bulk matter (H/He) by Coulomb collisions

  $\dot{M} \approx 10^{-6} \ldots 10^{-5}\, M_{sun}/\text{yr}, \; v_\infty \approx 2{,}000$ km/s

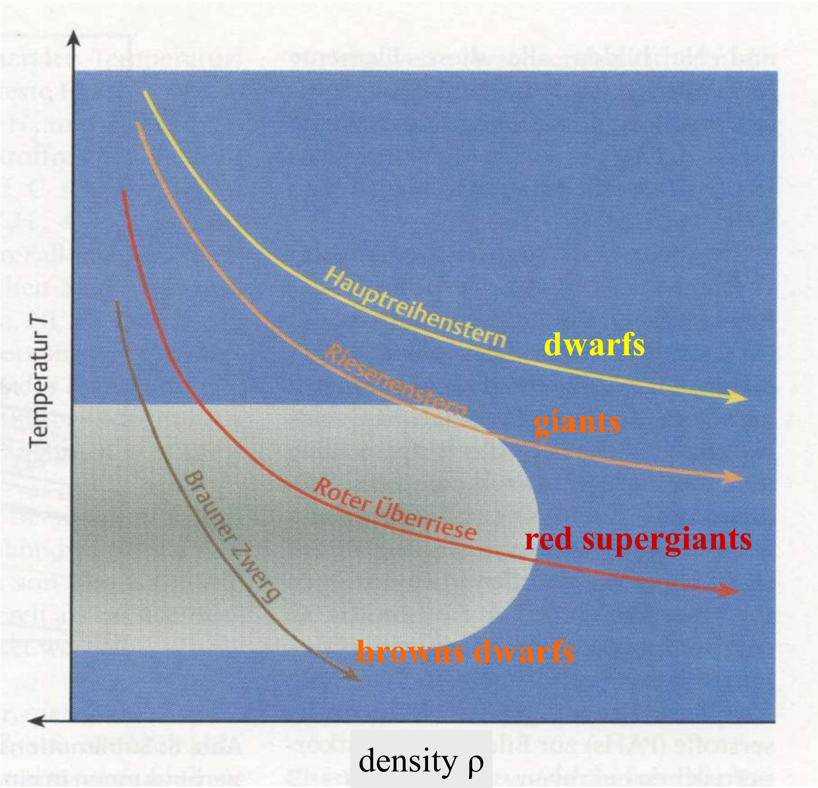
# dusty winds




(David Fabricius, 1596)

• brightness variations by 5.5 mag
(from 3.5 to 9),
corresponding to a factor of 160

# Cool supergiants: The dust-factories of our Universe




Material on this and following pages from Chr. Helling, *Sterne und Weltraum*, Feb/March 2002

**dust**: approx. 1% of ISM, 70% of this fraction formed in the winds of AGB-stars (cool, low-mass supergiants)

Red supergiants are located in dust-forming "window"

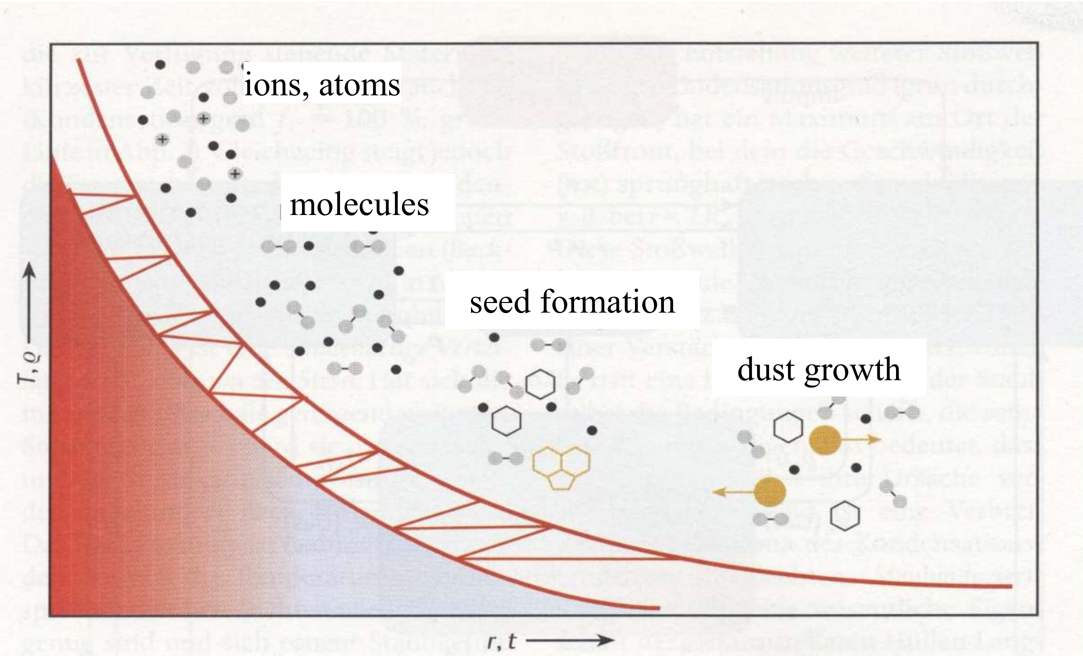
transition from gaseous phase to solid state possible only in **narrow range of temperature and density:**

gas density must be high enough and temperature low enough to allow for the chemical reactions:

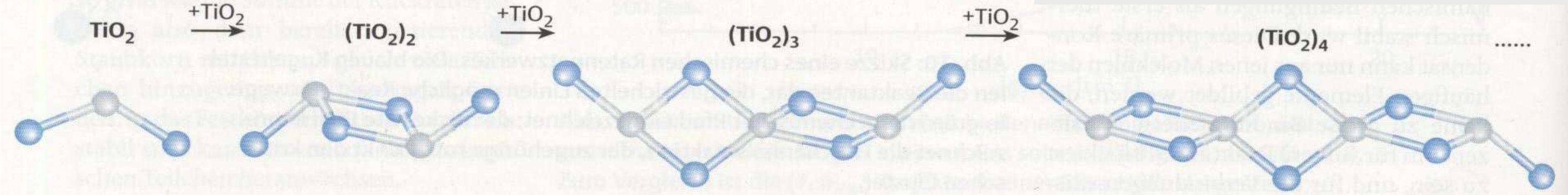
- sufficient number of dust forming molecules required
- the dust particles formed have to be thermally stable

# Growth of dust in matter outflow

- decrease of density and temperature
- more and more complex structures are forming
- dust: macroscopic, solid state body, approx. $10^{-7}$ m (1000 Angstrom), $10^{9}$ atoms

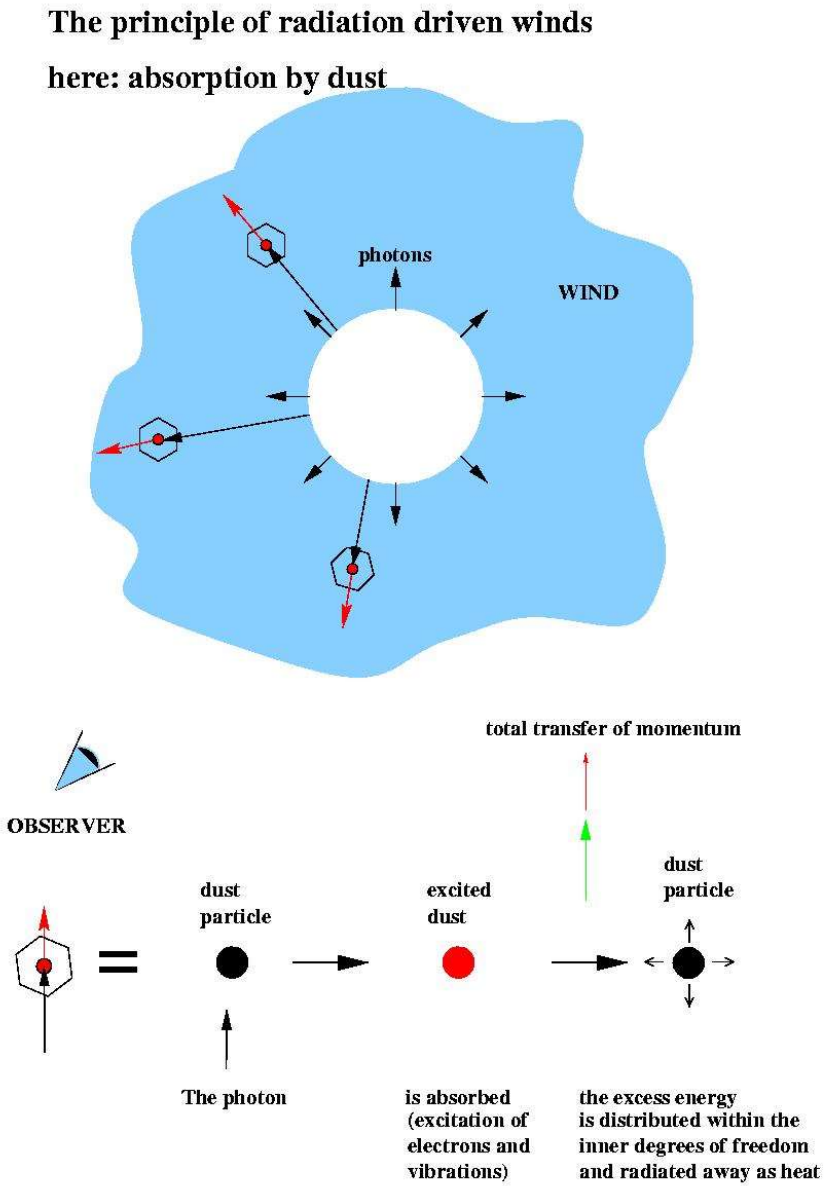



terrestrial, macroscopic rutile crystal ($TiO_2$, yellowish)




first steps of a linear reaction chain, forming the seed of $(TiO_2)_N$

# Dust-driven winds: the principle



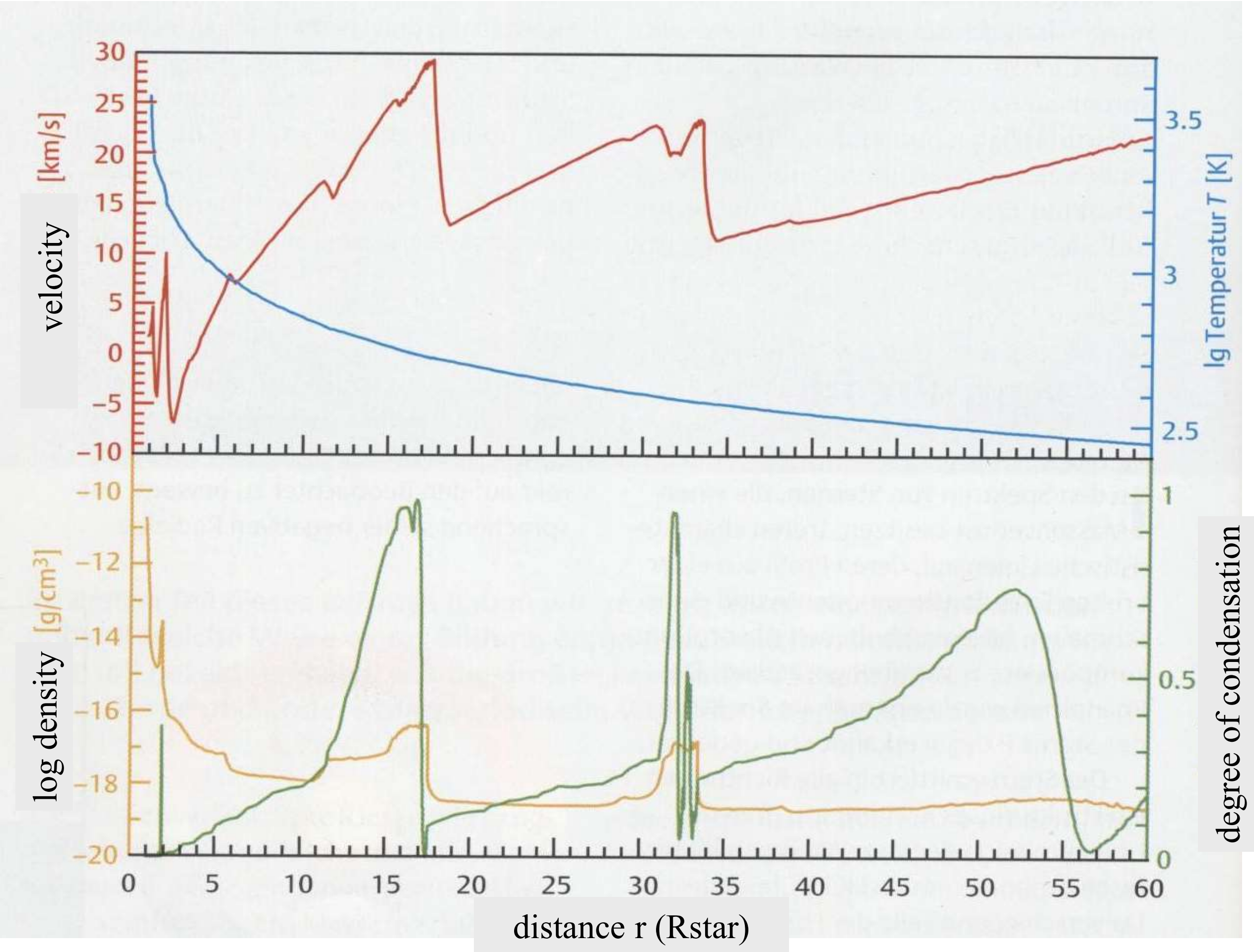

- star emits photons
- photons absorbed by dust
- momentum transfer accelerates dust
- gas accelerated by viscous drag force due to gas-dust collisions

acceleration
proportional to number of photons, i.e.,
proportional to *stellar luminosity* ***L***

$\Rightarrow$ mass-loss rate $\propto$ ***L***

dust driven winds at tip of AGB responsible for ejection of envelope
$\Rightarrow$ Planetary Nebulae

winds from massive red supergiants still not explained, but probably similar mechanism

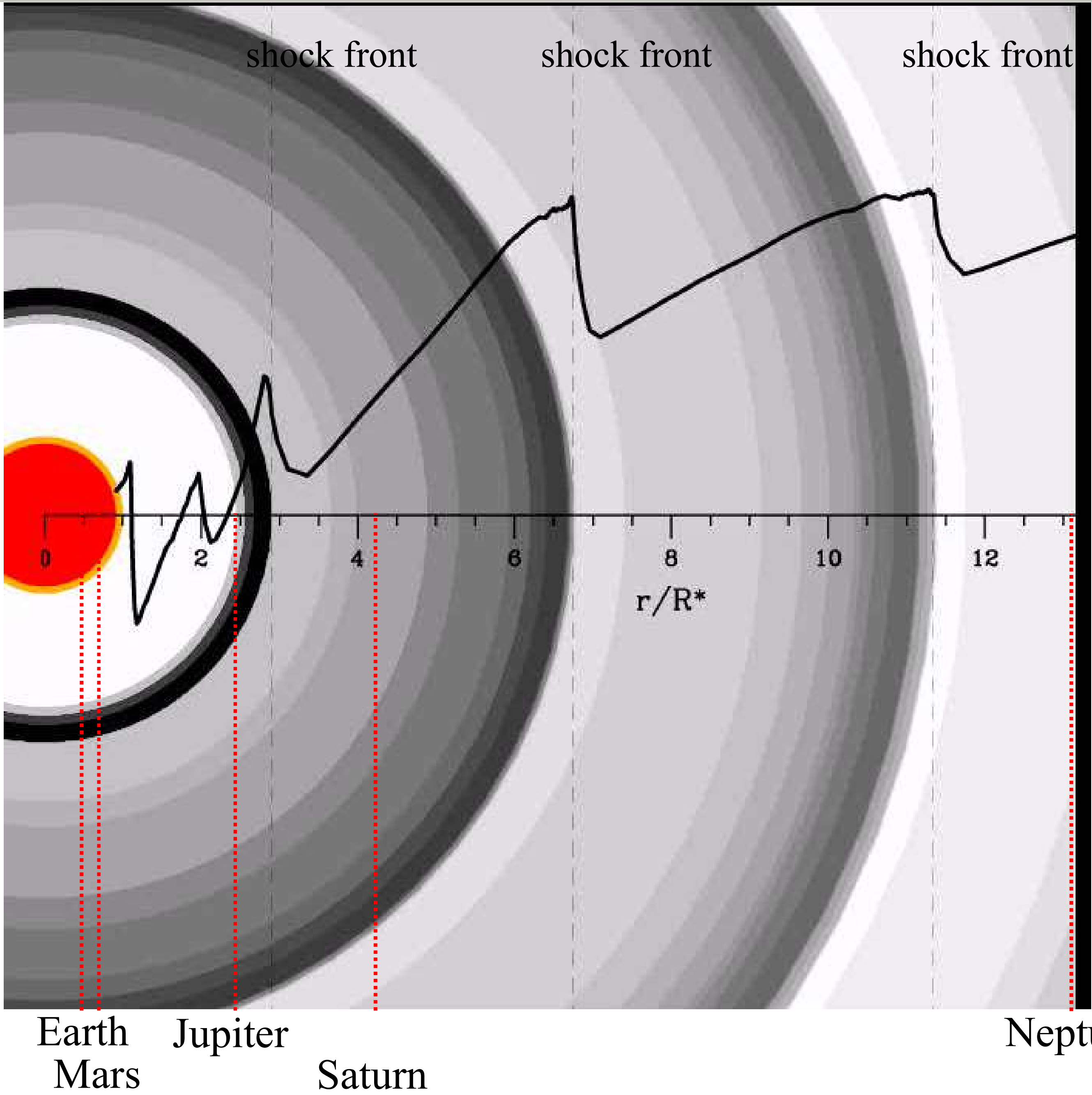



snapshot of a time-dependent hydro-simulation of a carbon-rich circumstellar envelope of an AGB-star. Model parameters similar to next slide.

- star (“surface”) pulsates,
- sound waves are created,
- steepen into shocks;
- matter is compressed,
- dust is formed
- and accelerated by radiation pressure

dust shells are blown away, following the pulsational cycle

$\Rightarrow$ periodic darkening of stellar disc
$\Rightarrow$ **brightness variations**




dark colors: dust shells

velocity

simulation of a
dust-driven wind
(*previous working group E. Sedlmayr, TU Berlin*)

$T = 2600$ K, $L = 10^4 L_{sun}$,
$M = 1 M_{sun}$, $\Delta v = 2$ km/s

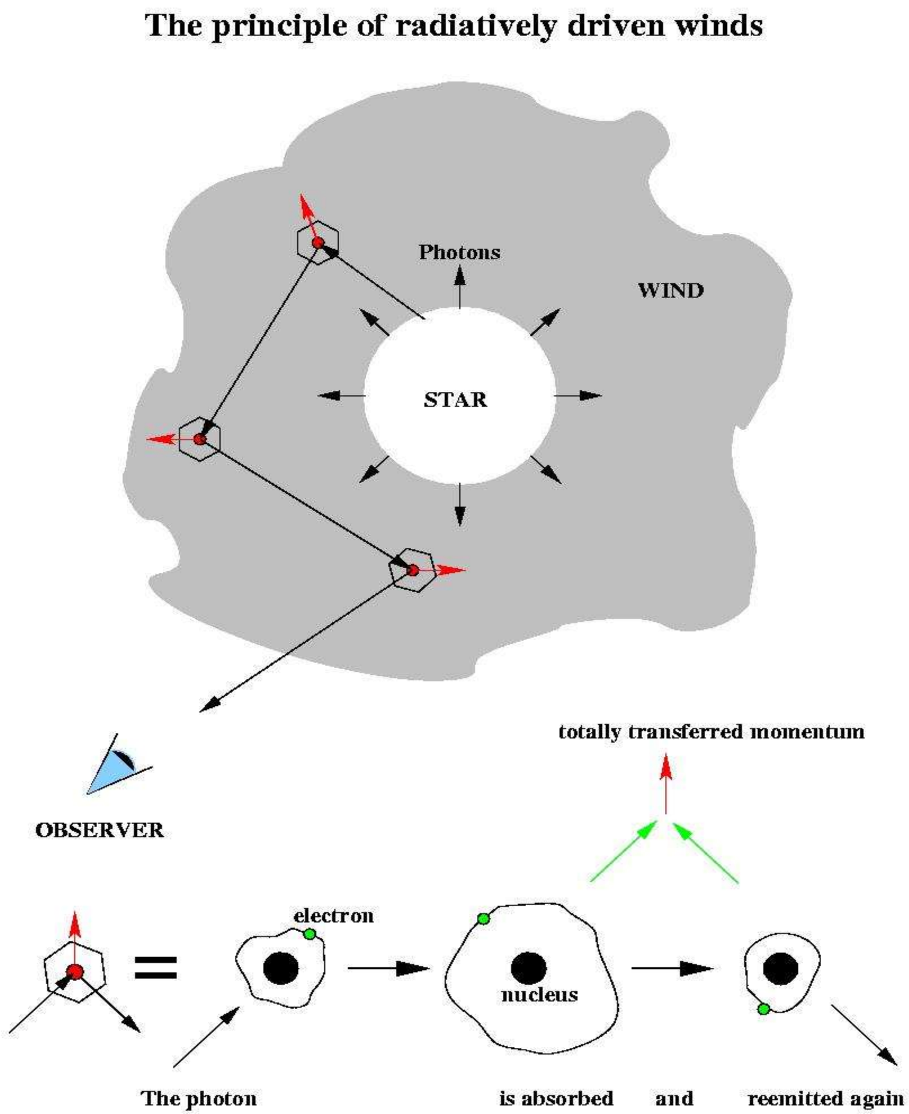
# Stars and their winds – typical parameters

| | The sun | Red AGB-stars | Blue supergiants |
|---|---|---|---|
| mass [$M_\odot$] | 1 | 1 ... 3 | 10...100 |
| luminosity [$L_\odot$] | 1 | $10^4$ | $10^5...10^6$ |
| stellar radius [$R_\odot$] | 1 | 400 | 10...200 |
| effective temperature [K] | 5570 | 2500 | $10^4...5\cdot10^4$ |
| | | | |
| wind temperature [K] | $10^6$ | 1000 | 8000...40000 |
| | | | |
| mass loss rate [$M_\odot$ /yr] | $10^{-14}$ | $10^{-6}...10^{-4}$ | $10^{-6}$ ... few $10^{-5}$ |
| terminal velocity [km/s] | 500 | 30 | 200...3000 |
| life time [yr] | $10^{10}$ | $10^5$ | $10^7$ |
| total mass loss [$M_\odot$] | $10^{-4}$ | $\gtrsim 0.5$ | up to 90% of total mass |

# Massive stars determine energy (kinetic and radiation) and momentum budget of surrounding ISM

Bubble Nebula
(NGC 7635)
in Cassiopeia

wind-blown
bubble around
BD+602522
(O6.5IIIf)




- accelerated by radiation pressure in **lines**
  $\dot{M} \approx 10^{-7} ... 10^{-5}\ M_{sun}$ / yr, $v_\infty \approx 200\ ...\ 3{,}000$ km/s
- momentum transfer from accelerated species (ions) to bulk matter (H/He) via Coulomb collisions

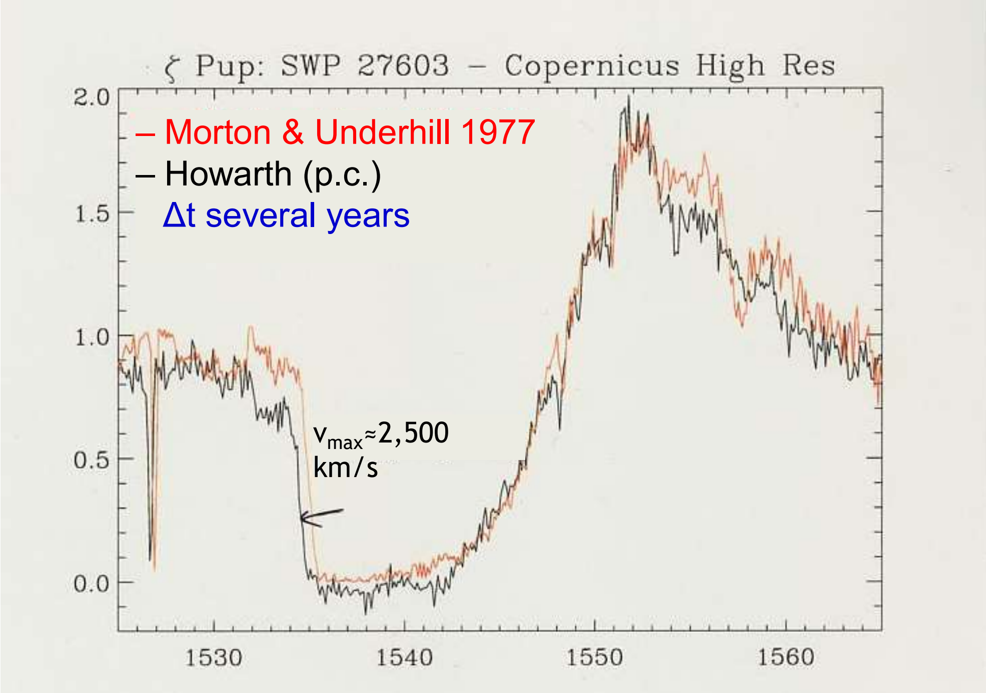
**Prerequesites for radiative driving**

- large number of photons => high luminosity
  $L \propto R_*^2\ T_{\rm eff}^4$ => supergiants or hot dwarfs
- line driving:
  large number of lines close to flux maximum
  (typically some $10^4 ... 10^5$ lines relevant)
  with high interaction probability
  (=> mass-loss dependent on metal abundances)

- line driven winds important for chemical evolution of (spiral) Galaxies, in particular for starbursts
- transfer of momentum (=> can induce *star formation*, hot stars mostly in *associations*), energy and nuclear processed material to surrounding environment
- dramatic impact on stellar evolution of massive stars (mass-loss rate vs. life time!)

**pioneering investigations** by
Lucy & Solomon, 1970, ApJ 159
Castor, Abbott & Klein, 1975, ApJ 195 (CAK)

**reviews** by Kudritzki & Puls, 2000, ARAA 38
Puls et al. 2008 A&Arv 16, issue 3




- Observational findings:
  massive star have outflows, at least quasi-stationary
- only small, in NO WAY dominant variability of global quantities $(\dot{M}, v_\infty)$
- $\dot{M}, v_\infty, v(r)$ have to be explained
- diagnostic tools have to be developed
- predictions have to be given

# Equation of motion in the standard model

Hydro-equations

$$\frac{\partial}{\partial t}\rho + \nabla\cdot(\rho\mathbf{v}) = 0 \qquad \text{continuity equation}$$

$$\frac{\partial}{\partial t}(\rho\mathbf{v}) + \nabla\cdot(\rho\mathbf{v}\mathbf{v}) = -\nabla p + \rho\mathbf{a}^{\text{ext}} \qquad \text{momentum equation}$$

$\Rightarrow$ (use continuity equation)

$$\frac{\partial}{\partial t}\mathbf{v} + (\mathbf{v}\cdot\nabla)\mathbf{v} = -\frac{1}{\rho}\nabla p + \mathbf{a}^{\text{ext}} \qquad \text{equation of motion}$$

$\Rightarrow$ (with $\frac{\partial}{\partial t} = 0$, 1-D spherically symmetric)

$$\boxed{4\pi r^2\rho(r)\mathrm{v}(r) = \text{const} = \dot{M}} \qquad \text{mass-loss rate}$$

$$\mathrm{v}\frac{d\mathrm{v}}{dr} = -\frac{1}{\rho(r)}\frac{dp}{dr} + a^{\text{ext}}(r)$$

$$p = NkT \quad (\text{equation of state}) = \frac{kT}{\mu m_{\mathrm{H}}}\rho = \mathrm{v}_{\mathrm{s}}^2\rho$$

$\mathrm{v}_{\mathrm{s}}$ isothermal sound speed, $\mu$ mean molecular weight

$$\Rightarrow \quad \boxed{\mathrm{v}\left(1-\frac{\mathrm{v}_{\mathrm{s}}^2}{\mathrm{v}^2}\right)\frac{d\mathrm{v}}{dr} = \frac{2\mathrm{v}_{\mathrm{s}}^2}{\mathrm{r}} - \frac{d\mathrm{v}_{\mathrm{s}}^2}{dr} + a^{\text{ext}}}$$

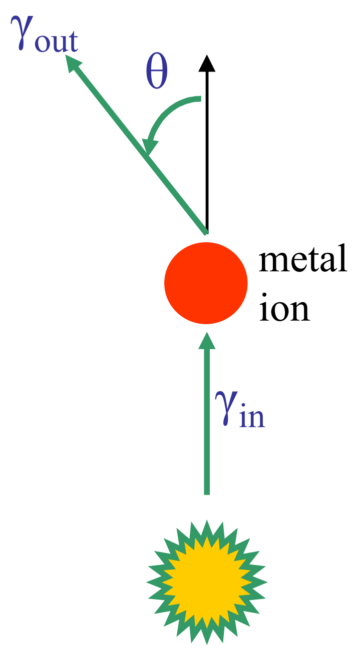
(assumption here: $\mathrm{v}_{\mathrm{s}}^2 \sim T$ known)

$$\boxed{a^{\text{ext}}(r) = -\frac{GM}{r^2}(1-\Gamma) + g_{\text{Rad}}^{\text{true cont}}(r) + g_{\text{Rad}}^{\text{line}}(r)}$$

$$\Gamma = \frac{g_{\text{Rad}}^{\text{Thomson}}(r)}{g_{\text{grav}}(r)} = \text{const is Eddington factor,}$$

corrects for radiative acceleration due to Thomson scattering

# Principle idea of line acceleration




$$\left.\begin{array}{l}\cos\theta_{in} \approx 1 \\ \text{isotropic reemission} \\ \langle\cos\theta_{out}\rangle = 0\end{array}\right\} \quad \langle\Delta P\rangle = \frac{h\nu_{in}}{c}$$

$$\Rightarrow g_{rad} = \frac{\langle\Delta P\rangle_{tot}}{\Delta t\,\Delta m} = \frac{\sum_{\text{all lines}}\langle\Delta P\rangle_i}{\Delta t\,\Delta m}$$

a) scattering of continuum light in resonance lines

$$\Delta P_{radial} = P_{in} - P_{out}$$
$$= \frac{h}{c}(\nu_{in}\underbrace{\cos\theta_{in}}_{\text{absorption}} - \nu_{out}\underbrace{\cos\theta_{out}}_{\text{reemission}})$$

b) momentum transfer from metal ions (fraction $10^{-3}$) to bulk plasma (H/He) via Coulomb collisions (see Springmann & Pauldrach 1992)

- velocity drift of ions w.r.t. H/He is compensated by frictional force as long as $v_D/v_{th} < 1$ (linear regime, "Stokes" law)

$$R_{ij}^{\text{fric}} \sim G(x_{ij}) \qquad x_{ij} = \sqrt{A_{ij}}\,\frac{|\mathrm{v}_i - \mathrm{v}_j|}{\mathrm{v}_{\text{th}}(\text{prot})} \qquad A_{ij} \text{ is reduced mass}$$

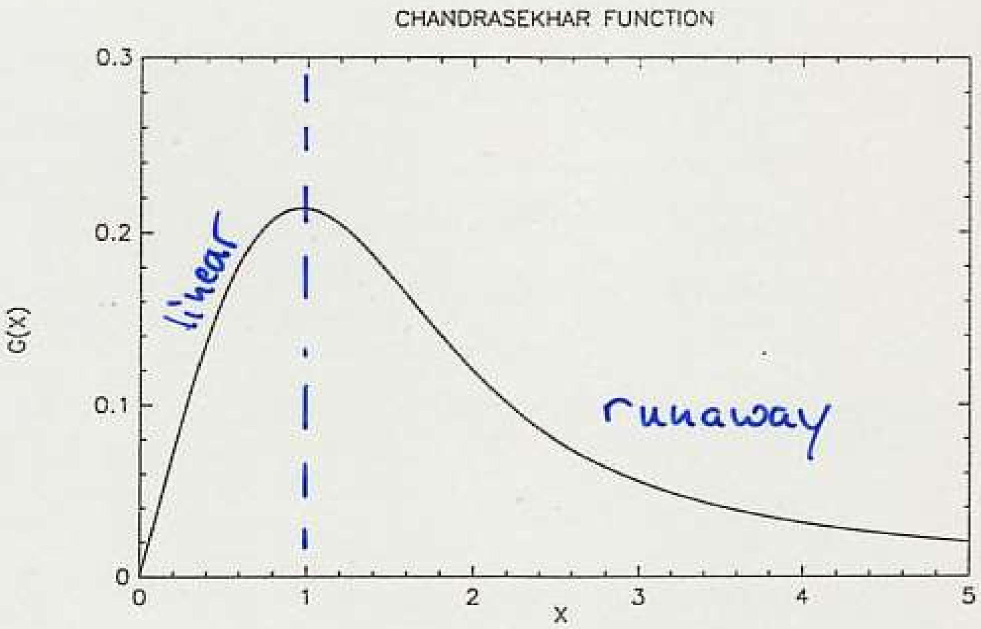



Fig. 1. The Chandrasekhar function $G(x)$ which gives the frictional force on test particles by field particles of unit density for an inverse square law of Coulomb interaction. The variable $x$ is essentially the ratio of the velocity of the test particles in the rest frame of the field particles to the thermal velocity of the field particles (see text). The limiting cases are $G(x) \sim x$ for $x \ll 1$ and $G(x) \sim x^{-2}$ for $x \gg 1$

from Springmann & Pauldrach (1992, A&A 262)
see also Owocki & Puls (2002, ApJ 568)

approximate description (supersonic regime) by linear diffusion equation

$$\mathrm{v}_{\text{ion}} \frac{d}{dr} \mathrm{v}_{\text{ion}} = g_{\text{Rad}}^{\text{ion}} - \frac{GM}{r^2} - \frac{w}{\tau_{ib}} \qquad w \text{ drift velocity}$$

$$\mathrm{v}_{\text{bulk}} \frac{d}{dr} \mathrm{v}_{\text{bulk}} = \quad - \frac{GM}{r^2} + \frac{w}{\tau_{bi}} \qquad \text{bulk} \approx \text{H/He},$$

$\tau$ relaxation time between collisions

in order to obtain one-component fluid,

$$\mathrm{v}_{\text{ion}} \frac{d\mathrm{v}_{\text{ion}}}{dr} = \mathrm{v}_{\text{bulk}} \frac{d\mathrm{v}_{\text{bulk}}}{dr}$$

$$\Rightarrow w = g_{\text{Rad}}^{\text{ion}} \left( \frac{1}{\tau_{ib}} + \frac{1}{\tau_{bi}} \right)^{-1} \approx g_{\text{Rad}}^{\text{tot}} \frac{\rho_{\text{tot}}}{\rho_{\text{ion}}} \cdot \tau \sim g_{\text{Rad}}^{\text{tot}} \frac{1}{Z} \frac{1}{\rho}$$

tot = bulk + ion, $Z$ is metallicity

for low $\rho \sim \frac{\dot{M}}{V}$ and/or low $Z$ $\rightarrow$ drift large $\rightarrow$ runaway

e.g., winds of A-dwarfs, Babel et al. 1995, A&A 301

# The photon-tiring limit

What is the maximum mass-loss rate that can be accelerated???

- mechanical luminosity in wind at infinity is

$$L_{\rm wind} = \dot{M}\left(\frac{v_\infty^2}{2} + \frac{GM}{R}\right) = \dot{M}\left(\frac{v_\infty^2}{2} + \frac{v_{\rm esc}^2}{2}\right) \quad \text{with } v_{\rm esc} = \sqrt{\frac{2GM}{R}}$$

- maximum mass loss, if $L_{\rm wind} = L_*$ $\quad \Rightarrow \; L(\infty) = 0,$ star becomes invisible

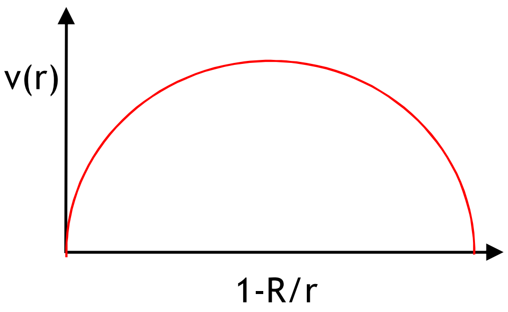
$$\dot{M}_{\max} = \frac{2L_*}{v_\infty^2 + v_{\rm esc}^2}$$

$$\Rightarrow \eta_{\max} = \frac{\dot{M}_{\max} v_\infty}{L/c} = \frac{2c}{v_\infty\left(1+\left(\frac{v_{\rm esc}}{v_\infty}\right)^2\right)}$$

typical values: $v_\infty \approx 2000...3000$ km/s $\approx 0.01c$, $v_{\rm esc}/v_\infty \approx 1/3$ $\quad \rightarrow \eta_{\max} \approx 200$

$\dot{M}_{\rm tir}$ (Owocki & Gayley 1997) is maximum mass-loss rate when wind just escapes the gravitational potential, with $v_\infty \rightarrow 0$

$$\boxed{\dot{M}_{\rm tir} = \frac{2L_*}{v_{\rm esc}^2} = 0.032\frac{M_\odot}{\rm yr}\frac{L_*}{10^6 L_\odot}\frac{R}{R_\odot}\frac{M_\odot}{M} = 0.0012\frac{M_\odot}{\rm yr}\Gamma_e\frac{R}{R_\odot}}$$



# Calculation of the line force

crucial point of the problem

$$g_{\mathrm{Rad}}^{\mathrm{line}} = \frac{4\pi}{c\rho}\frac{1}{2}\int_0^\infty d\nu \int_{-1}^{1} \mu d\mu \left[\chi_\nu^{\mathrm{line}}(r,\mu) I_\nu(r,\mu) - \cancel{\eta_\nu^{\mathrm{line}}(r,\mu)}\right]$$

absorbed emitted

→ (in single-line approximation: no interaction of different lines)

$$g_{\mathrm{Rad}}^{\mathrm{line}} = \frac{2\pi}{c\rho}\sum_{\mathrm{lines\ i}} \int_{\mathrm{line}} d\nu \int_{-1}^{1} \mu d\mu \; \chi_\nu^{\mathrm{i}}(r,\mu) I_\nu^{i}(r,\mu)$$

- two quantities to be known
  - force/line in response to $\chi_\nu$
  - distribution of lines with $\chi_\nu$ and $\nu$

# The force per line

- super-simplified
- simplified: “Sobolev approximation” :
  assume that opacities and source functions
  are constant inside τ-integral,
  i.e., over Doppler-shifted profile function
  → analytic solution possible, purely local
- “exact”:
  - comoving frame, special cases
  - observer’s frame, instability

# Super-simplified theory

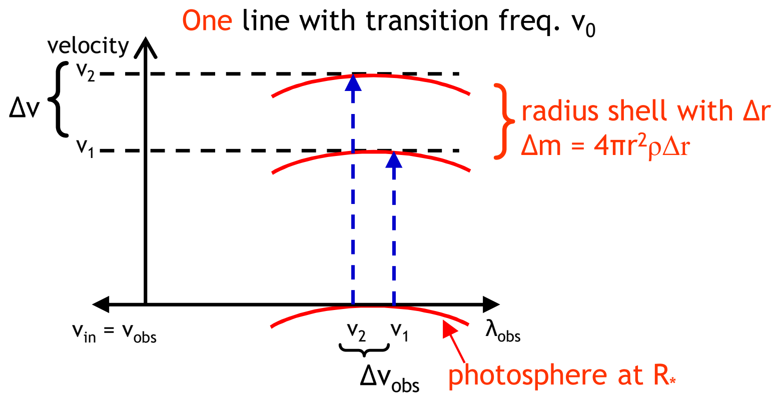
interaction with line at $\nu_0$, when comoving frame frequency of photon starting at $R_*$ with $\nu_{\mathrm{obs}}$ is equal to $\nu_0$

(finite profile width neglected, interaction probability = 1)

$$\nu_{\mathrm{CMF}} = \nu_{\mathrm{obs}} - \frac{\nu_0 \mathrm{v}(r)}{c} =: \nu_0$$ (Doppler shift, radial photons, $\mu$=1, assumed)

$$\left.\begin{aligned} \nu_0 &= \nu_1^{\mathrm{obs}} - \frac{\nu_0}{c}\mathrm{v}_1(r) \\ \nu_0 &= \nu_2^{\mathrm{obs}} - \frac{\nu_0}{c}\mathrm{v}_2(r) \end{aligned}\right\}$$ scattering at larger v requires 'bluer' photons

$$\Rightarrow \Delta\nu_{\mathrm{obs}} = \frac{\nu_0}{c}\Delta\mathrm{v}$$




Number of photons in interval $\left[\nu_1^{\mathrm{obs}}, \nu_2^{\mathrm{obs}} = \nu_1^{\mathrm{obs}} + \Delta\nu_{\mathrm{obs}}\right]$ per unit time

$$\frac{N_\nu \Delta\nu}{\Delta t} = \frac{L_\nu \Delta\nu}{h\nu_{\mathrm{obs}}} \quad \Rightarrow \quad \left(g_{\mathrm{Rad}} = \frac{\Delta P}{\Delta t \Delta m}\right)$$

$$g_{\mathrm{Rad}} = \frac{h\nu_{\mathrm{obs}}}{\mathrm{c}} \cdot \frac{L_\nu \Delta\nu}{h\nu_{\mathrm{obs}}} \cdot \frac{1}{\Delta m} = \qquad \left(\Delta\nu = \frac{\nu_0}{c}\Delta\mathrm{v}\right)$$

$$= \frac{L_\nu \nu_0}{c^2} \frac{\Delta\mathrm{v}}{\Delta r} \frac{1}{4\pi r^2 \rho} \propto \boxed{\frac{d\mathrm{v}}{dr}} \frac{1}{r^2 \rho}$$

# Why $g_{rad} \propto dv/dr$ ?

shell of matter with spatial extent $\Delta r$,

and velocity $v_0 + \left(\frac{dv}{dr}\right)_1 \Delta r$

absorption of photons at $\nu_0 \pm \delta\nu$

**in frame of matter**

photons must start at higher (stellar)

frequencies, are "seen" at $\nu_0 \pm \delta\nu$

in frame of matter because of Doppler-effect.

Let $\Delta\nu$ be frequency band contributing to

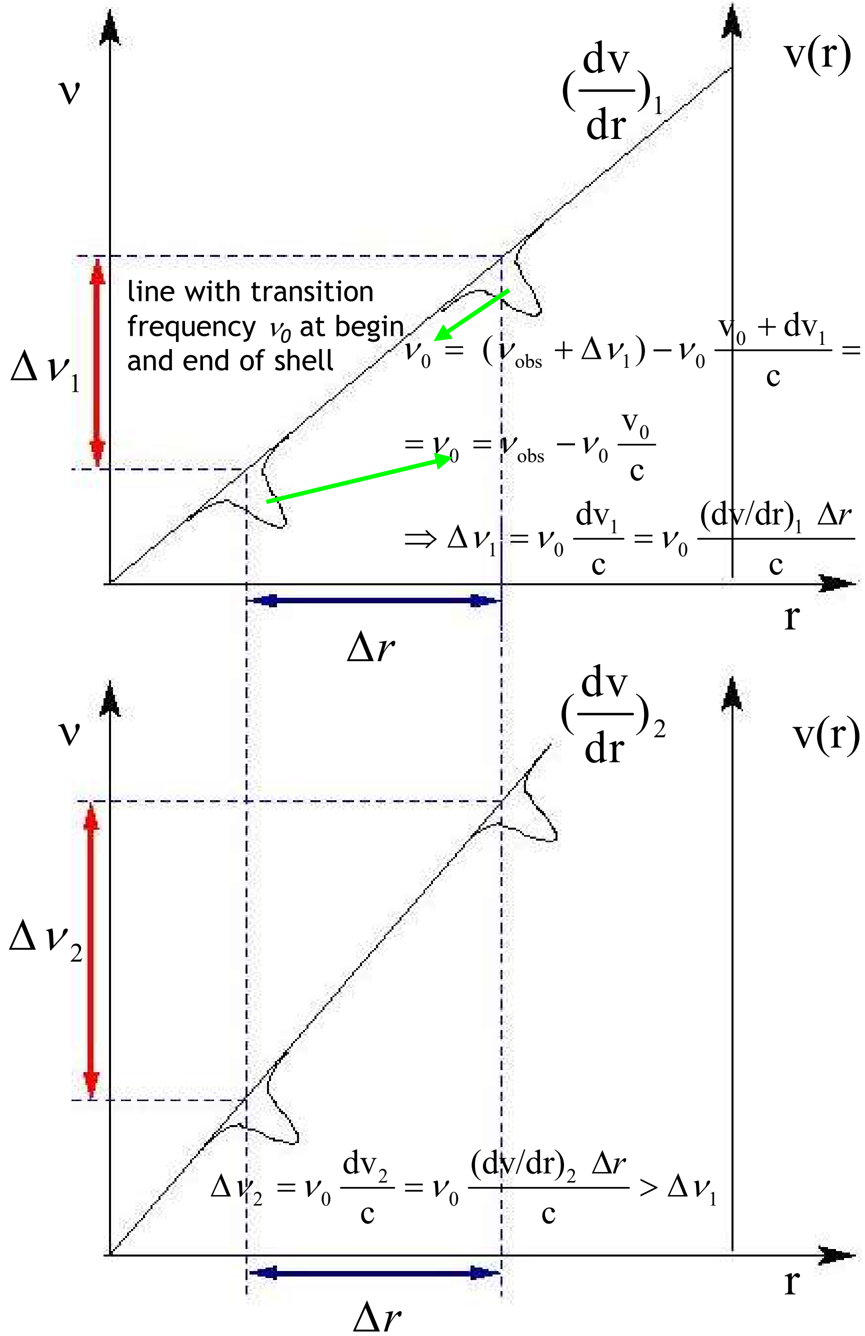
acceleration of matter in $\Delta r$

The larger $\frac{dv}{dr}$,

- the larger $\Delta\nu$
- the more photons can be absorbed
- the larger the acceleration

$$\boxed{g_{rad} \propto \frac{dv}{dr}}$$

(assuming that each photon is absorbed, i.e., acceleration from optically thick lines)



# Accounting for finite interaction probability

$$\boxed{g_{\text{rad}}(\text{one line at } \nu_0) = \frac{L_\nu \nu_0}{c^2}\frac{\Delta \text{v}}{\Delta r}\frac{1}{4\pi r^2}}$$

Assumption was: each photon is scattered

Then: $g_{\text{rad}}$ independent of cross-sections, occuption numbers etc.
only dependent on hydro-structure and flux distribution

What happens if interaction probability < 1?

interaction probability $= 1 - \mathrm{e}^{-\tau}$, with optical depth $\tau$

$\tau \gg 1$ prob = 1

$\tau \ll 1$ prob = $\tau$

Now: division in two classes

optically thick lines, $\tau \geq 1$ $\xrightarrow{\approx}$ prob = 1 (saturation, independent of $\tau$)

optically thin lines $\tau < 1$ $\xrightarrow{\approx}$ prob = $\tau$

$$\Rightarrow \boxed{g_{\text{rad}}(\text{optically thin line}) = \tau \cdot g_{\text{rad}}(\text{optically thick line})}$$

# Line acceleration from a line ensemble

$$g_{\mathrm{Rad}}^{\mathrm{tot}}(r) = \sum_{\mathrm{thick}} g_{\mathrm{Rad}}^{\mathrm{i}}(r) + \sum_{\mathrm{thin}} g_{\mathrm{Rad}}^{\mathrm{j}}(r) =$$

$$= \frac{1}{4\pi r^2 c^2}\left( \sum_{\mathrm{thick}} L_\nu \nu_{\mathrm{i}} \underbrace{\frac{d\mathrm{v}}{dr}\frac{1}{\rho}}_{k_1} + \sum_{\mathrm{thin}} L_\nu \nu_{\mathrm{i}} \underbrace{\frac{d\mathrm{v}}{dr}\frac{\tau_i}{\rho}}_{k_{\mathrm{i}}} \right)$$

$$\tau_{\mathrm{i}} = \frac{\bar{\chi}_{\mathrm{Li}}\lambda_{\mathrm{i}}}{d\mathrm{v}/dr} =: \frac{k_{\mathrm{i}}\rho(r)}{d\mathrm{v}/dr} \qquad \left(\text{precisely: } k_{\mathrm{i}} = \frac{\bar{\chi}_{\mathrm{Li}}\lambda_{\mathrm{i}}}{\rho s_{\mathrm{e}} \mathrm{v}_{\mathrm{th}}}\right)$$

$\uparrow$ optical depth of line in "Sobolev theory"

$k_{\mathrm{i}}$ is line strength $\sim \dfrac{\sigma_{\mathrm{i}} n_{\mathrm{i}}(r)\lambda_{\mathrm{i}}}{\rho(r)}$ $\quad \sigma_{\mathrm{i}}$ cross section,

$n_{\mathrm{i}}$ lower occup. number of line transition

$k_{\mathrm{i}}$ roughly constant in wind!!!

Which line strength corresponds to 'border' $\tau_{\mathrm{i}} = 1$?

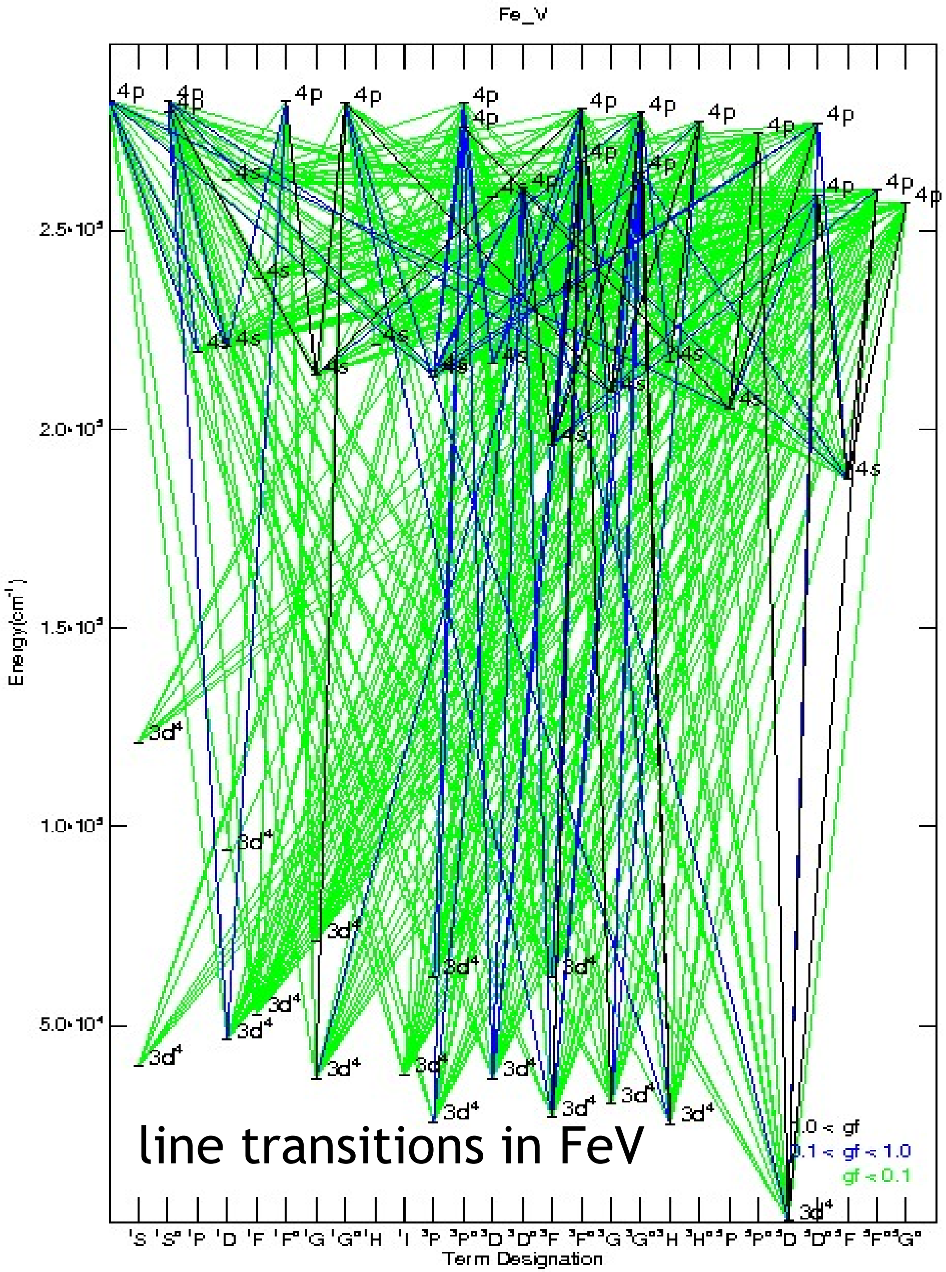
$$1 = \frac{k_1 \rho}{d\mathrm{v}/dr} \quad \Rightarrow \quad k_1 = \frac{d\mathrm{v}/dr}{\rho}$$

$$\Rightarrow \quad g_{\mathrm{Rad}}^{\mathrm{tot}}(r) = \frac{1}{4\pi r^2 c^2}\left( k_1 \sum_{k_{\mathrm{i}} > k_1} L_\nu \nu_{\mathrm{i}} + \sum_{k_{\mathrm{i}} < k_1} L_\nu \nu_{\mathrm{i}} k_{\mathrm{i}} \right)$$

optically thick — optically thin

depends on hydrostruct. — depends on line-strength

# Millions of lines ....




... are present
... and needed!

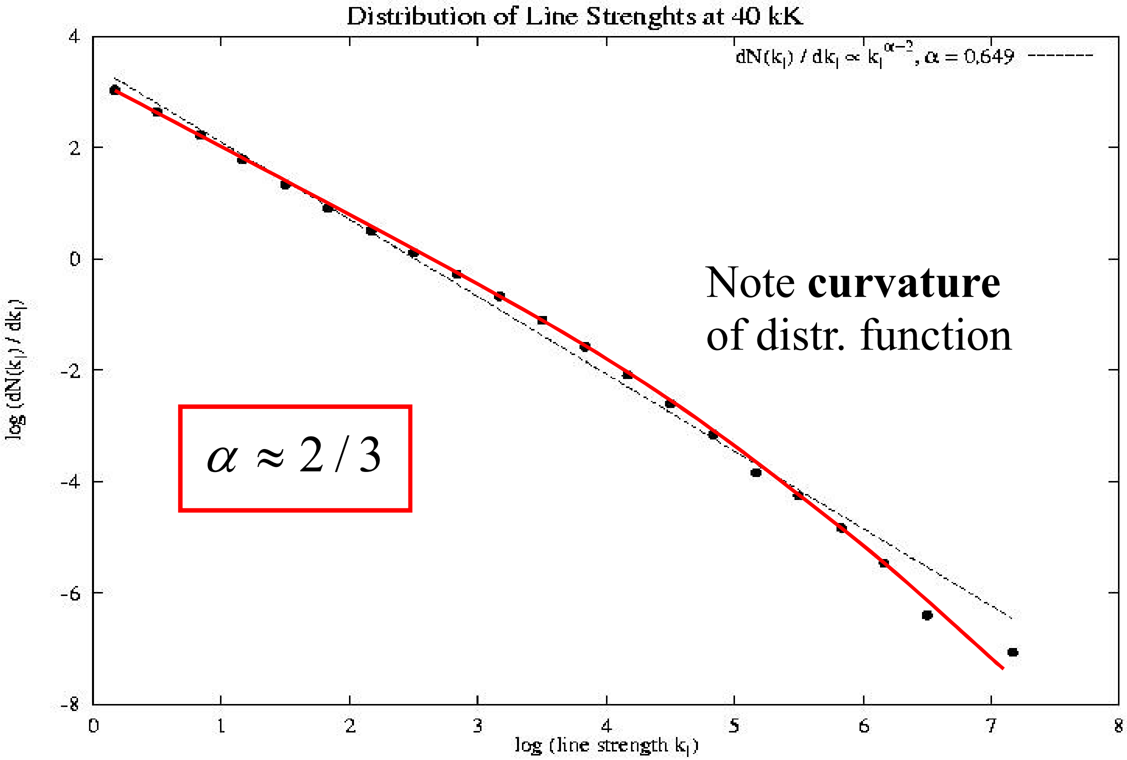
$$g_{\rm Rad}^{\rm tot} = \sum_{\text{all lines}} g_{\rm Rad}^{\rm i},$$

$$g_{\rm Rad}^{\rm thin} = L_\nu^{\rm i} \nu_{\rm i}\, k_{\rm i}, \qquad k_{\rm i} \propto \frac{\bar{\chi}_{\rm i} \lambda_{\rm i}}{\rho} \quad \text{(line-strength)}$$

$$g_{\rm Rad}^{\rm thick} = L_\nu^{\rm i} \nu_{\rm i} \frac{d{\rm v}/dr}{\rho} \propto L_\nu^{\rm i} \nu_{\rm i} k_1$$

# The line distribution function

- pioneering work by **Castor, Abbott & Klein** (CAK, 1975):
  - from glance at CIII atom in LTE, they suggested that ALL line-strengths follow a power-law distribution
- first realistic line-strength distribution function by Kudritzki et al. (1988)
- NOW: couple of MI (Mega lines), 150 ionization stages (H – Zn), NLTE




$$\frac{dN(k)}{dk} = k^{\alpha-2}, \quad \alpha \approx 0.6 \ldots 0.7$$

+ 2nd empirical finding: valid in *each* frequential subinterval

$$dN(k,\nu) = -N_0 f(\nu) d\nu \, k^{\alpha-2} dk$$

Logarithmic plot of line-strength distribution function for an O-type wind at 40,000 K and corresponding power-law fit (see Puls et al. 2000, A&AS 141)

# Force/line + line-strength distribution

$$\Rightarrow g_{\mathrm{Rad}}^{\mathrm{tot}}(r) = \frac{1}{4\pi r^2 c^2}\left( k_1 \sum_{k_i > k_1} L_\nu \nu_i + \sum_{k_i < k_1} L_\nu \nu_i k_i \right) \rightarrow$$

$$\rightarrow \frac{1}{4\pi r^2 c^2}\left( \int_0^\infty k_1 \int_{k_{\max}}^{k_1} L(\nu)\nu \, dN(k,\nu) + \int_0^\infty \int_{k_1}^0 L(\nu)\nu k \, dN(k,\nu) \right) =$$

$$= \frac{N_0 \int L(\nu)\nu f(\nu) d\nu}{4\pi r^2 c^2} \left( \underbrace{\underbrace{k_1 \int_{k_1}^{k_{\max}} k^{\alpha-2} dk}_{k_1 \frac{1}{1-a} k_1^{\alpha-1}} + \underbrace{\int_0^{k_1} k \cdot k^{\alpha-2} dk}_{\frac{1}{\alpha} k_1^{\alpha}}}_{\frac{1}{\alpha(1-\alpha)} k_1^{\alpha}} \right)$$

$\Rightarrow$ final result

$$\boxed{g_{\mathrm{Rad}}^{\mathrm{tot}}(r) = \frac{\mathrm{const}}{4\pi r^2} k_1^{\alpha}}$$

very ‘strange’ acceleration, non-linear in dv/dr

$$k_1 = \frac{d\mathrm{v}/dr}{\rho} = \frac{4\pi}{\dot{M}} r^2 \mathrm{v} \frac{d\mathrm{v}}{dr}; \quad \mathrm{const} = \frac{N_0 \int L(\nu)\nu f(\nu) d\nu}{c^2 \alpha(1-\alpha)}$$

# 9.2 Theoretical predictions for line-driven winds

first hydro-solution developed by CAK 1975, ApJ 195, improved for non-radial photons and ionization effects by Pauldrach, Puls & Kudritzki 1986, A&A 164 and Friend & Abbott 1986, ApJ 311

- non-linear differential equation
- has 'singular point' in analogy to solar wind
- $v_{crit} >> v_s$ (100… 200 km/s)
- solution: iteration of singular point location/velocity, integration inwards and outwards

had equation of motion

$$v\left(1-\frac{v_s^2}{v^2}\right)\frac{dv}{dr}=\frac{2v_s^2}{r}-\frac{dv_s^2}{dr}+a^{ext}(r)$$

$$a^{ext}(r)=-\frac{GM}{r^2}(1-\Gamma)+\cancel{g_{Rad}^{true\ cont}}(r)+g_{Rad}^{line}(r)$$

for 'normal' winds

$$g_{Rad}^{line}(r)=f\cdot\frac{L}{r^2}k_1^{\alpha}$$

$$k_1=\frac{r^2 v dv/dr}{\dot{M}/(4\pi)} \qquad f=f(r,v,\frac{dv}{dr},\dot{M}) \text{ if all subtleties included}$$

All together

$$v\left(1-\frac{v_s^2}{v^2}\right)\frac{dv}{dr}=-\frac{GM}{r^2}(1-\Gamma)+\frac{2v_s^2}{r}-\frac{dv_s^2}{dr}+\frac{f\cdot L}{r^2}\left(\frac{\dot{M}}{4\pi}\right)^{-\alpha}\left(r^2 v\frac{dv}{dr}\right)^{\alpha}$$

# Approximate solution

(see also Kudritzki et al., 1989, A&A 219)

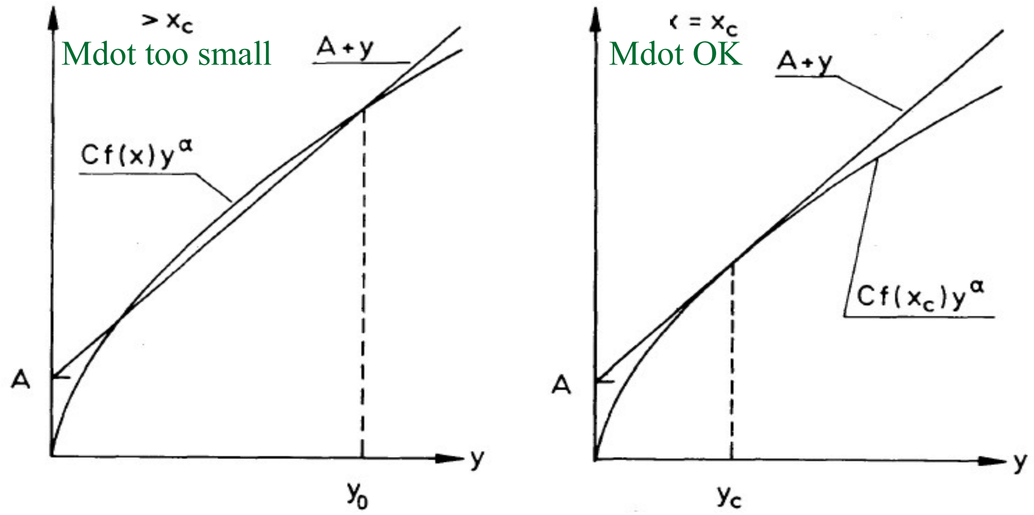
- supersonic → pressure terms vanish
- radially streaming photons → f $(4\pi)^\alpha$ → const

$$v\frac{dv}{dr} = -\frac{GM}{r^2}(1-\Gamma) + \frac{\text{const}\cdot L}{r^2}\dot{M}^{-\alpha}\left(r^2 v\frac{dv}{dr}\right)^\alpha$$

$$\Rightarrow\; y + A = \text{const}\cdot L\cdot \dot{M}^{-\alpha} y^\alpha \quad \Rightarrow \quad y \text{ is constant}$$

with $A = GM(1-\Gamma)$, $\quad y = r^2 v\dfrac{dv}{dr}$

graphical solution (Cassinelli et al. 1979, ARAA 17, Kudritzki et al. 1989)




for unique solution, derivatives have to be EQUAL!

$$y + A = \text{const}\cdot L\cdot \dot{M}^{-\alpha} y^\alpha \qquad \text{equation of motion}$$

and equality of derivatives

$$1 = \text{const}\cdot L\cdot \dot{M}^{-\alpha}\,\alpha\, y^{\alpha-1} \quad \text{at critical point } y_c$$

$$\dot{M}^{-\alpha} = \frac{1}{\text{const}\cdot L\cdot\alpha} y_c^{1-\alpha}$$

in equation of motion at critical point

$$y_c + A = \frac{1}{\alpha} y_c, \qquad \text{i.e., } y_c\left(1-\frac{1}{\alpha}\right) = -GM(1-\Gamma)$$

$$y_c = \frac{\alpha}{1-\alpha} GM(1-\Gamma) \overset{!}{=} y$$

finally …

# Scaling relations for line-driven winds (without rotation)

- $\dot{M} \propto N_{\rm eff}^{\frac{1}{\alpha'}} L^{\frac{1}{\alpha'}} \left(M(1-\Gamma)\right)^{1-\frac{1}{\alpha'}}$ scaling law for $\dot{M}$
- $r^2 \mathrm{v} \dfrac{d\mathrm{v}}{dr} = \dfrac{\alpha}{1-\alpha} GM(1-\Gamma)$

  $\rightarrow$ Integration between $\infty$ and $R_*$
- $\mathrm{v}(r) = \mathrm{v}_\infty \left(1 - \dfrac{R_*}{r}\right)^\beta, \quad \beta = \begin{cases} 0.5 \text{ for approx. solution, "CAK-velocity law"} \\ 0.8 \text{ (O-stars) ... 2 (BA-SG), see next slide} \end{cases}$
- $\mathrm{v}_\infty = \left(\dfrac{\alpha}{1-\alpha}\right)^{\frac{1}{2}} \left(\dfrac{2GM(1-\Gamma)}{R_*}\right)^{\frac{1}{2}}$ scaling law for $v_\infty$
- $\rightarrow \mathrm{v}_\infty \approx 2.25 \dfrac{\alpha}{1-\alpha} \mathrm{v}_{\rm esc}$, if all subtleties included

$\Gamma$ Eddington factor, accounting for acceleration by Thomson-scattering, diminishes effective gravity

$N_{\rm eff}$ number of lines effectively driving the wind, dependent on metallicity and spectral type

$\alpha$ exponent of line-strength distribution function, $0 < \alpha < 1$
large value: more optically thick lines

$\alpha' = \alpha - \delta$, with $\delta$ ionization parameter,
typical value for O-stars: $\alpha' \approx 0.6$

# The wind-momentum luminosity relation (WLR)

- use scaling relations for Mdot and $v_\infty$, calculate
  modified wind-momentum rate

$$\dot{M}\,\mathrm{v}_\infty \propto N_{\mathrm{eff}}^{1/\alpha'}\, L^{1/\alpha'} \big(M(1-\Gamma)\big)^{1-1/\alpha'} \frac{\big(M(1-\Gamma)\big)^{1/2}}{R_*^{1/2}}$$

$$\dot{M}\,\mathrm{v}_\infty\, R_*^{1/2} \propto N_{\mathrm{eff}}^{1/\alpha'}\, L^{1/\alpha'} \cancel{\big(M(1-\Gamma)\big)^{3/2-1/\alpha'}}$$

# The wind-momentum luminosity relation (WLR)

- use scaling relations for Mdot and $v_\infty$, calculate
  modified wind-momentum rate

$$\dot{M}\, v_\infty R_*^{1/2} \propto N_{\text{eff}}^{1/\alpha'} L^{1/\alpha'} \qquad \text{since } (\alpha' \approx \frac{2}{3})$$

independent of $M$ and $\Gamma$!!!!!

$$\boxed{\log(\dot{M}\, v_\infty R_*^{1/2}) \approx \frac{1}{\alpha'} \log L + const(z,\ \text{sp.type})}$$

- stellar winds contain info about stellar radius!!!

(Kudritzki, Lennon & Puls 1995)

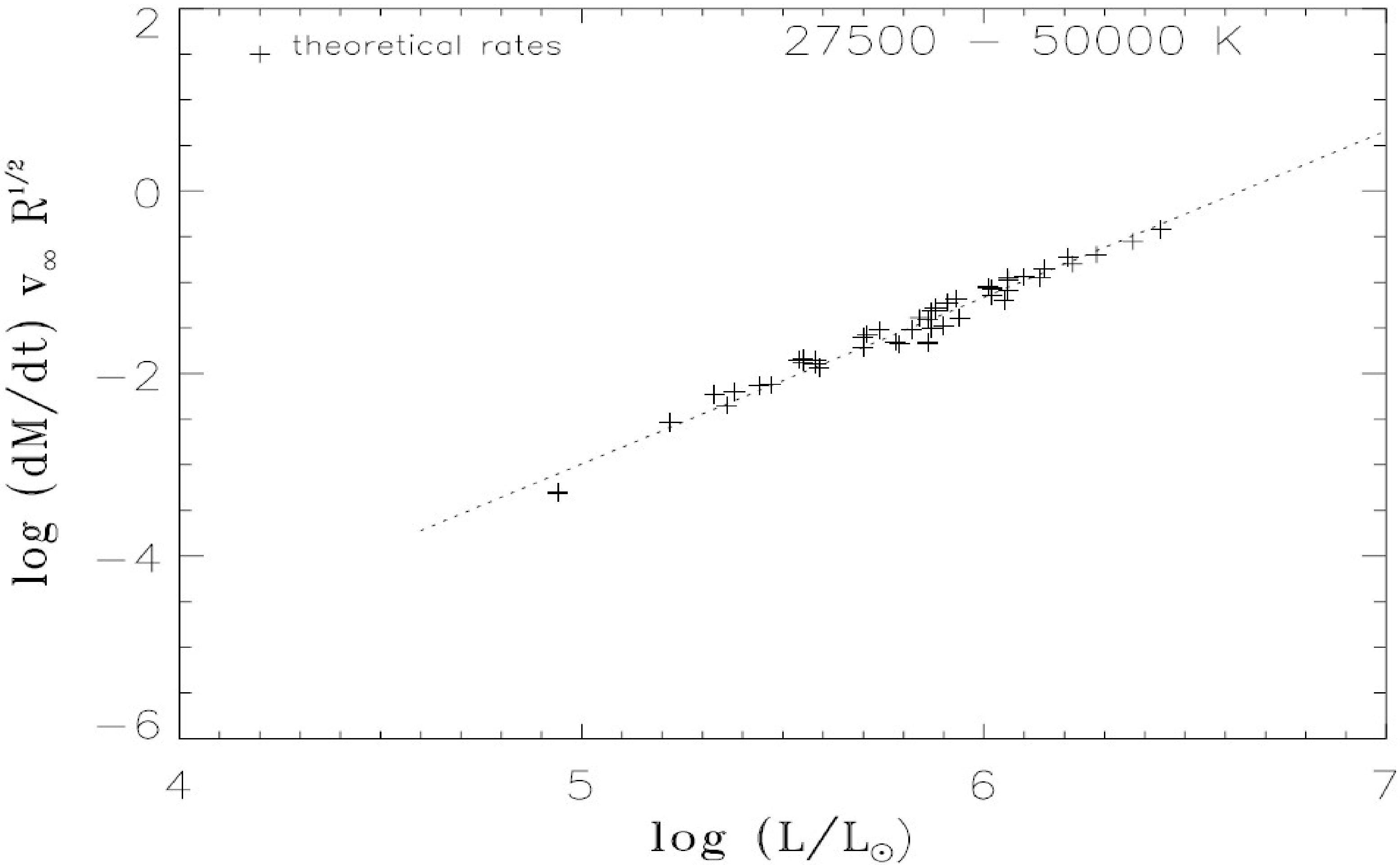
- (at least) two applications

  (1) construct observed WLR, calibrate as a function of spectral type and metallicity ($N_{\text{eff}}$ and α' depend on both parameter)

  – **independent tool to measure extragalactic distances** from *wind-properties*, Teff and metallicity

  (2) compare with theoretical WLR to test validity of radiation driven wind theory

# Validity of WLR concept




Theoretical wind-momentum rates as a function of luminosity, as calculated by Vink et al. (2000). Though multi-line effects (line overlaps) are included, the WLR concept (derived from simplified arguments) holds!

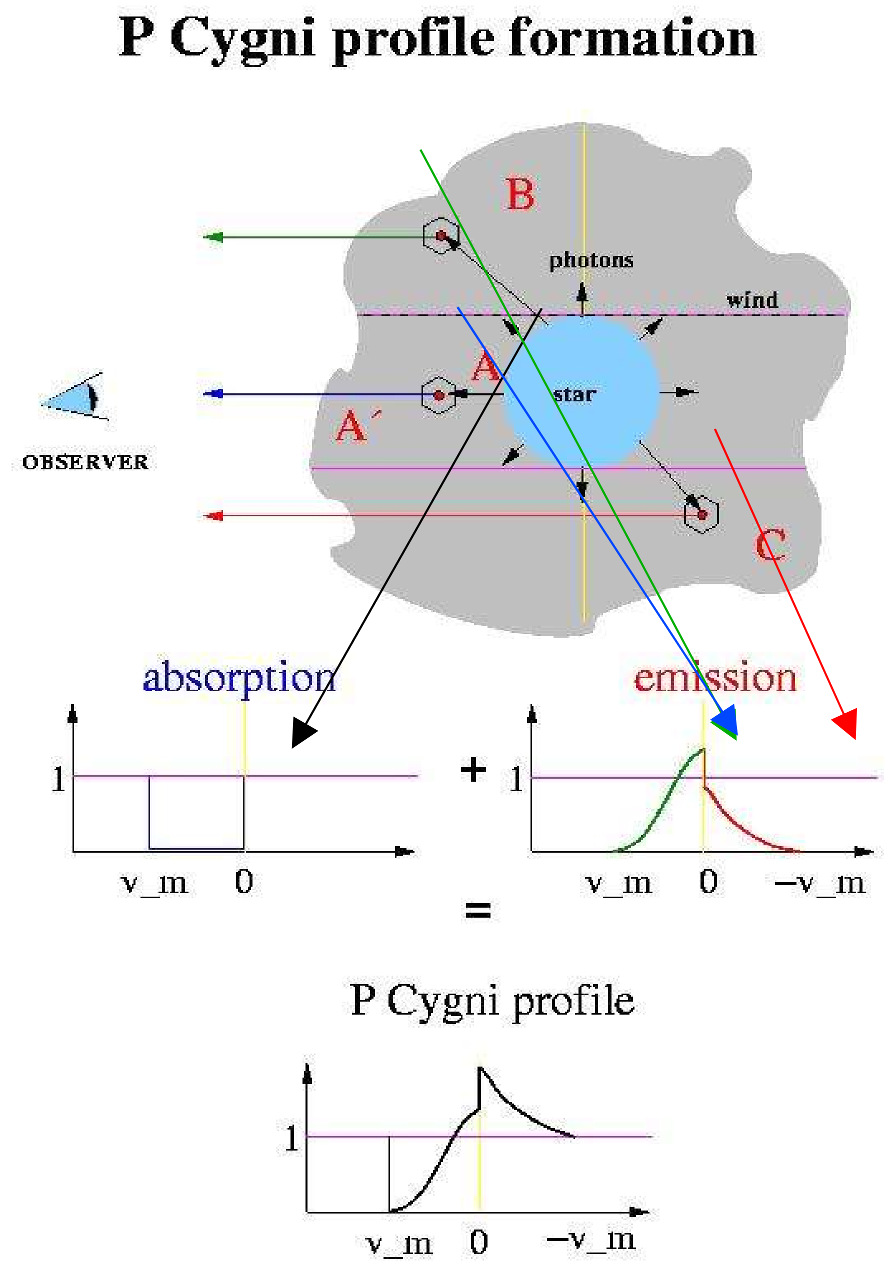
# Determination of wind-parameters: $v_\infty$

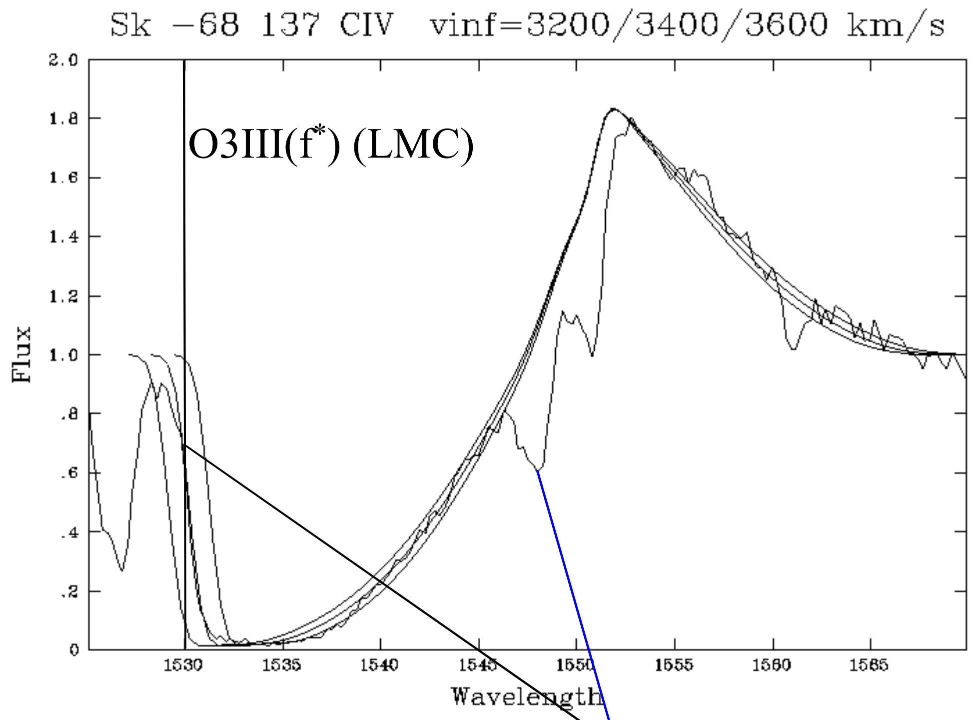
$$\nu_{\text{obs}} = \nu_0 \left(1 + \frac{\mu v(r)}{c}\right); \quad \nu_0 \text{ line frequency in CMF}$$

$$\mu v(r) > 0: \quad \nu_{\text{obs}} > \nu_0 \quad \text{blue side}$$

$$\mu v(r) < 0: \quad \nu_{\text{obs}} < \nu_0 \quad \text{red side}$$

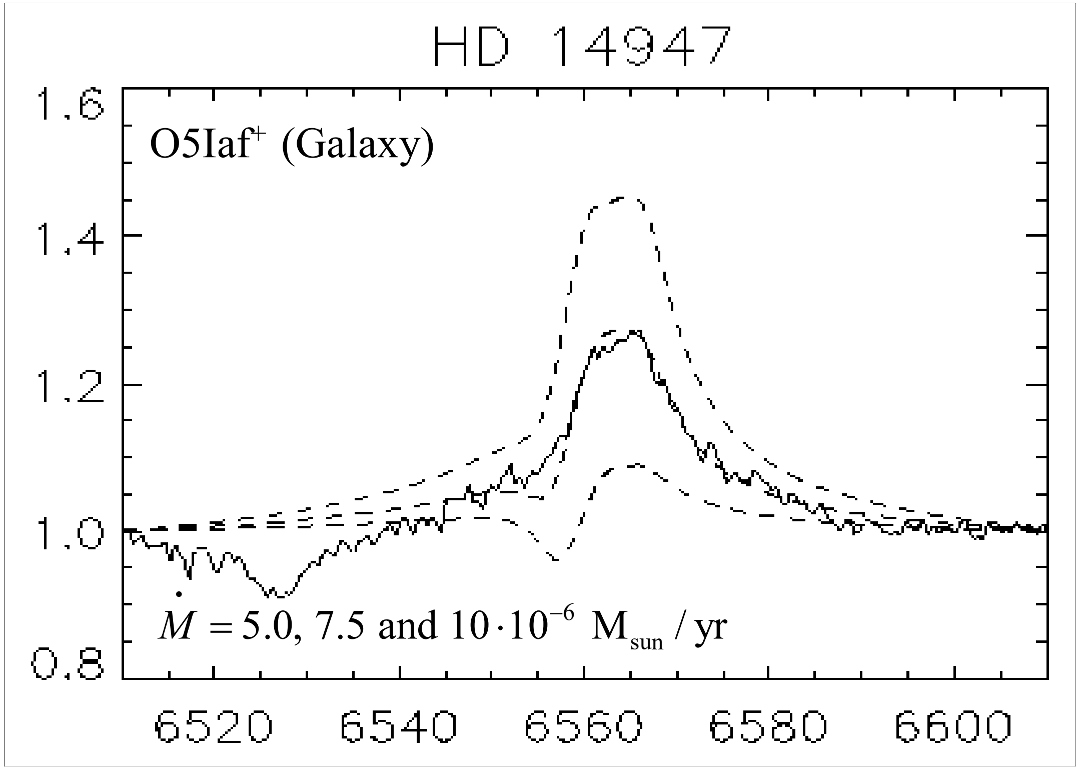
$$\frac{v_{\text{m}}}{\text{c}} = \frac{\nu_{\max} - \nu_0}{\nu_0} = 1 - \frac{\lambda_{\min}}{\lambda_0}$$

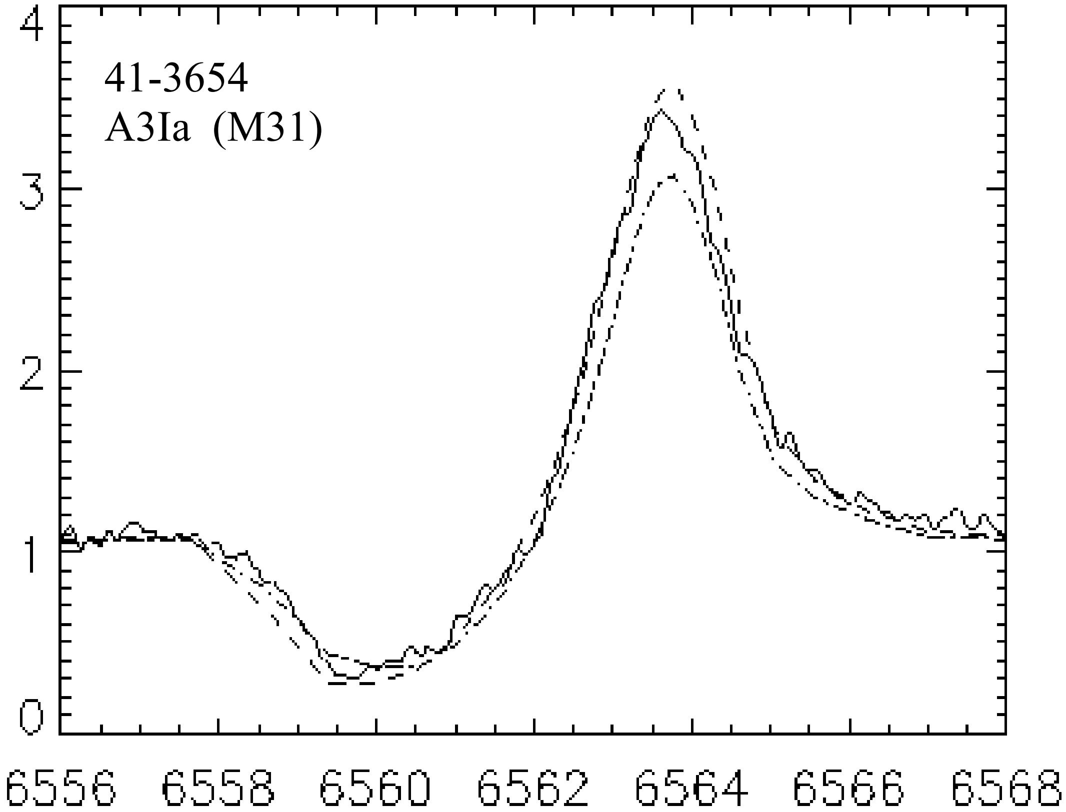



$$v_{\text{m}} \approx 2.998 \cdot 10^5 \left(1 - \frac{1530}{1548}\right) \approx 3{,}480 \text{ km/s}$$

# Determination of mass-loss rate from $H_\alpha$



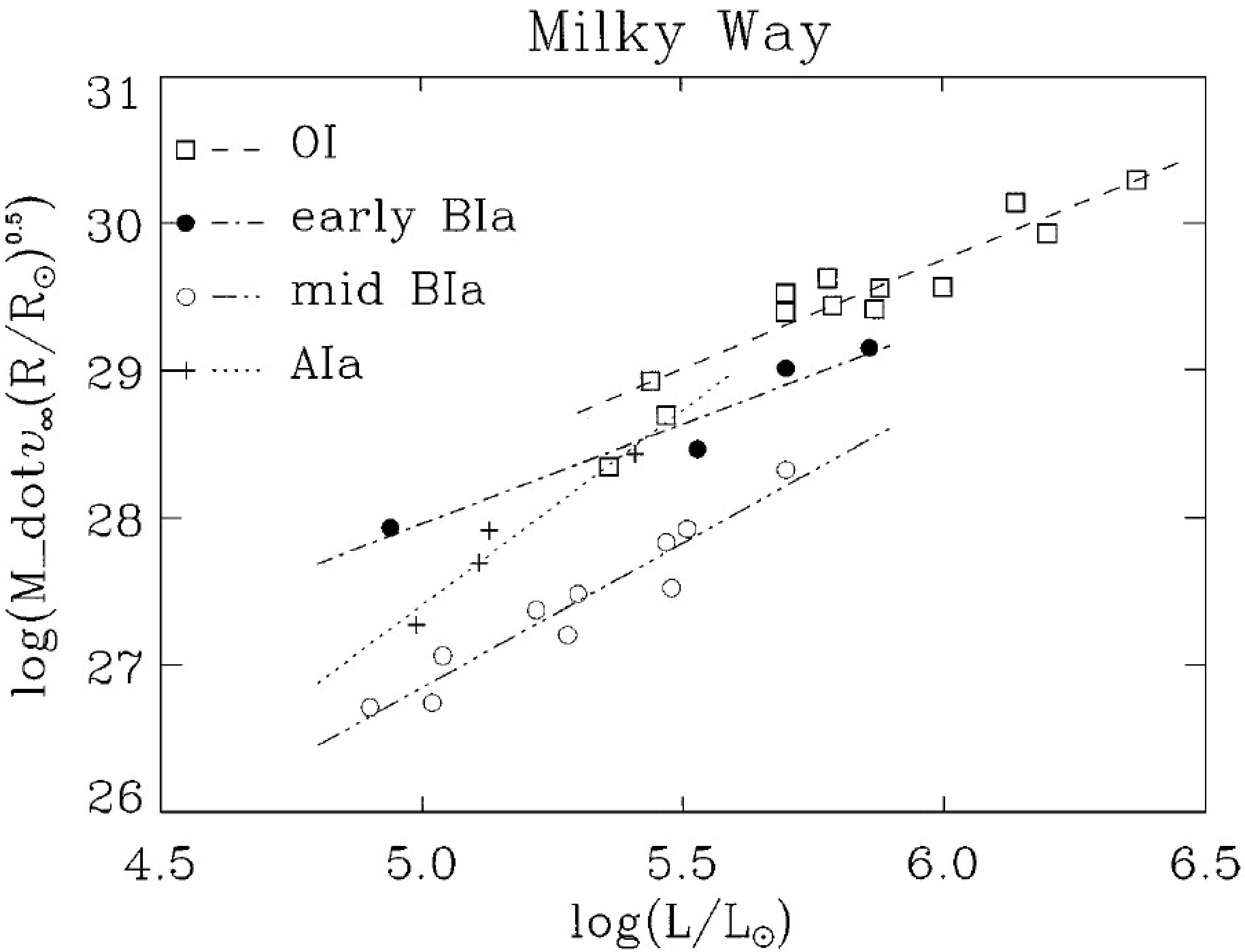

**Note:** Wind parameters can be cast into one quantity

$$Q = \frac{\dot{M}}{(R_* v_\infty)^{1.5}} \quad \text{or} \quad Q' = \frac{\dot{M}}{R_*^{1.5}}$$

For same values of Q(') (albeit different combinations of Mdot, $v_\infty$ and $R_*$), profiles look almost identical!

$H_\alpha$ taken with the Keck HIRES spectrograph, compared with two model calculations adopting $\beta = 3$, $v_\infty = 200$ km/s and *Mdot* = 1.7 and 2.1 $\times 10^{-6}$ $M_{sun}$/yr.

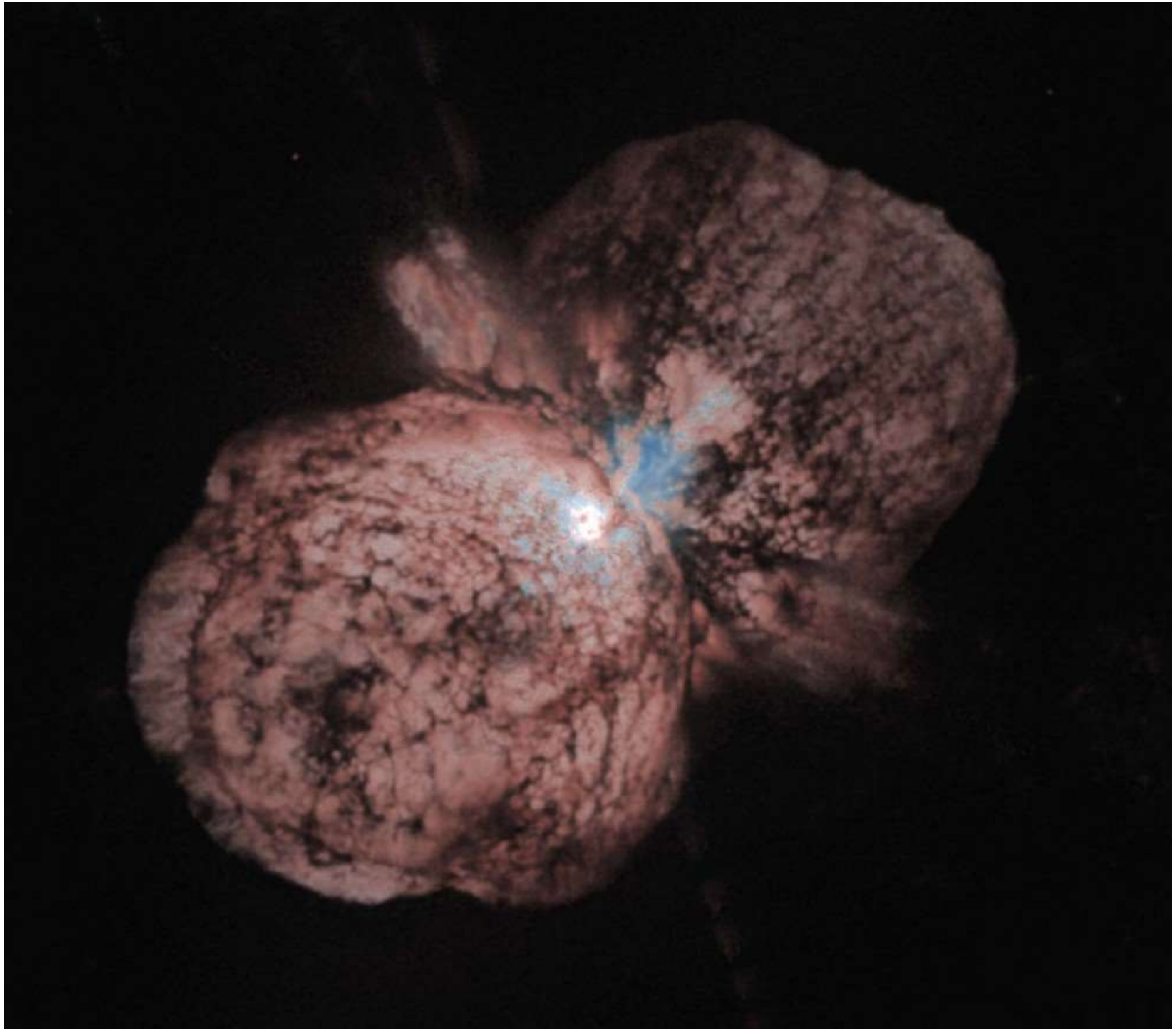
# Observed WLR




Modified wind momenta of Galactic O-, early B-, mid B- and A-supergiants as a function of luminosity, together with specific WLR obtained from linear regression. (From Kudritzki & Puls, 2000, ARAA 38).

# η Car: Aspherical ejecta

image by HST



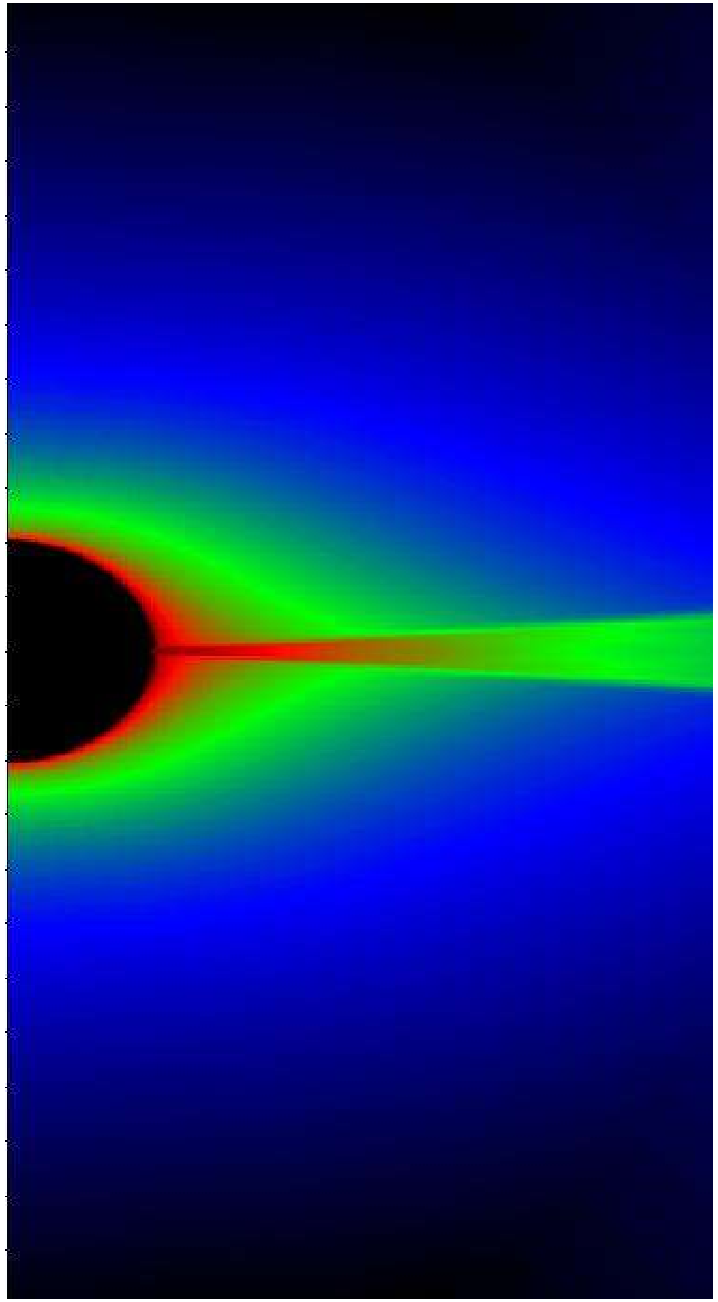
# Influence of rotation

hot, massive stars = young stars

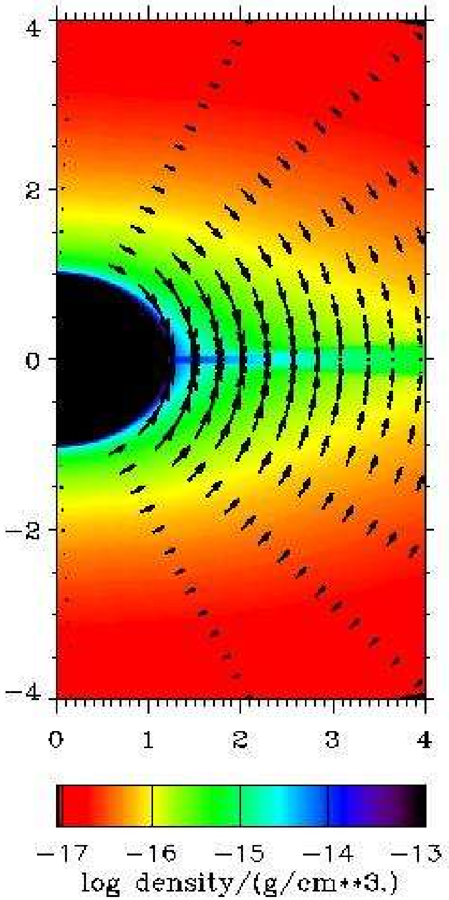
rapidly rotating (up to several 100 km/s)

twofold effect

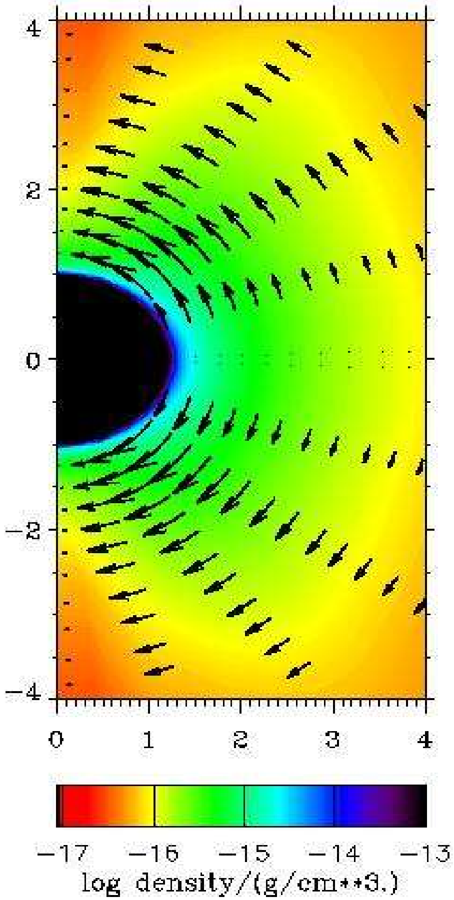
- star becomes “oblate”
- wind has to react on additional centrifugal acceleration, large in equatorial, small in polar regions



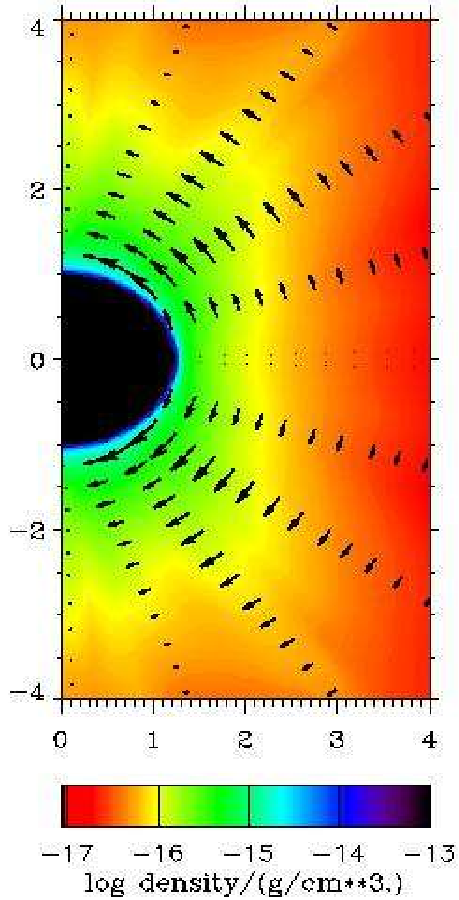
# Prolate or oblate wind structure?




purely radial radiative acceleration:
wind-compressed disk

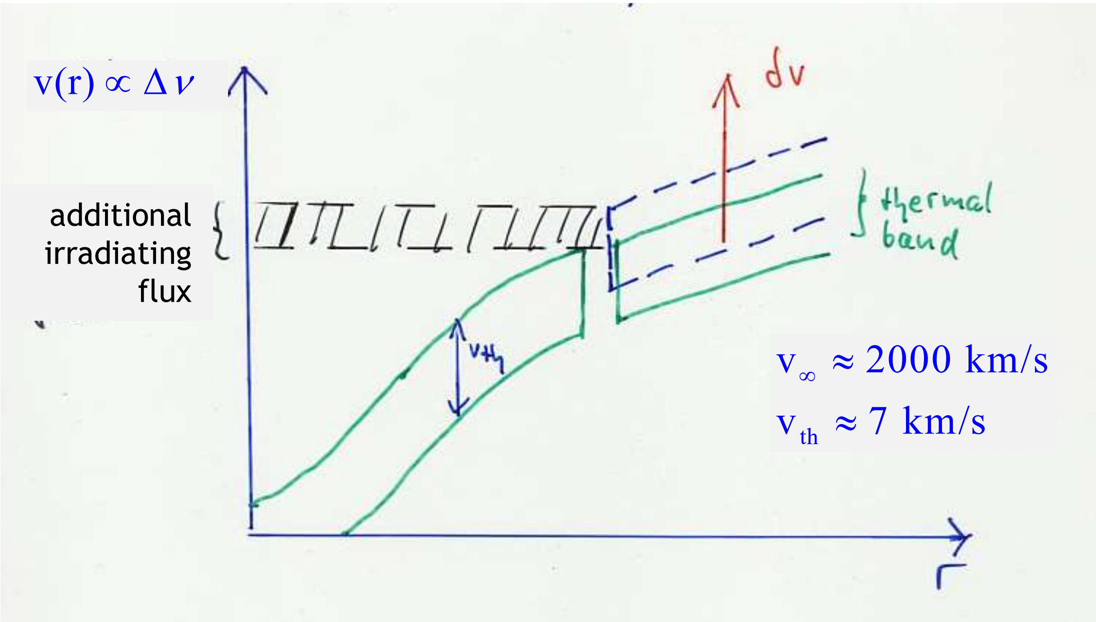



inclusion of non-radial component of line-acceleration (rotation breaks symmetry)




non-radial line-acceleration plus „gravity darkening“:
prolate geometry

# The line-driven instability




perturbation $\delta v \uparrow$

→ profile shifted to higher freq.

→ line 'sees' more stellar flux

→ line force grows $\delta g \uparrow$

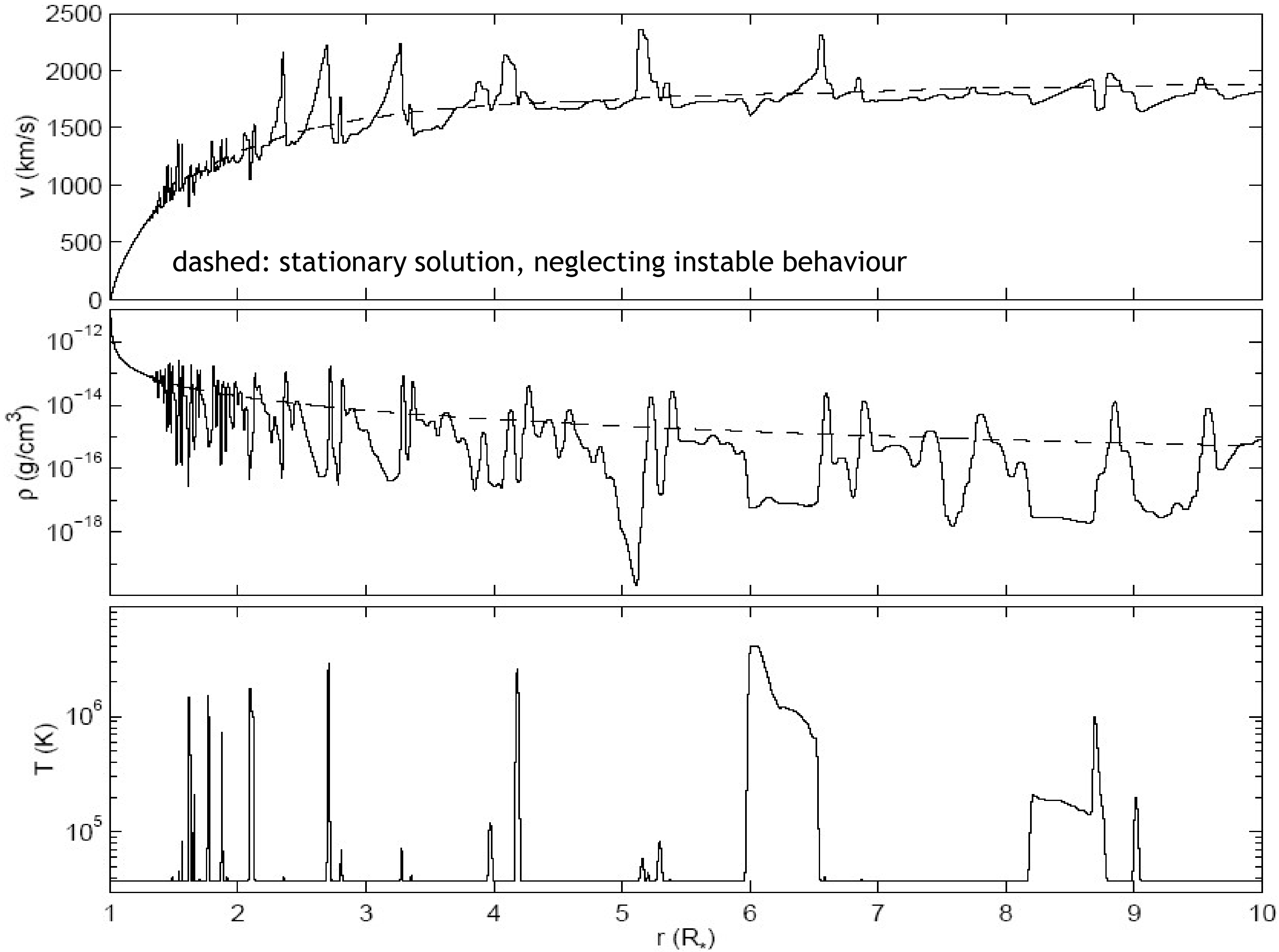
→ additional acceleration $\delta v \uparrow$

exponential growth of perturbation

$\delta g_{\mathrm{Rad}} \propto \delta v$

[for details, see MacGregor et al.1979 and Carlberg 1980]

# Time dependent hydro-simulations of line-driven winds: Snapshot of density, velocity and temperature structure



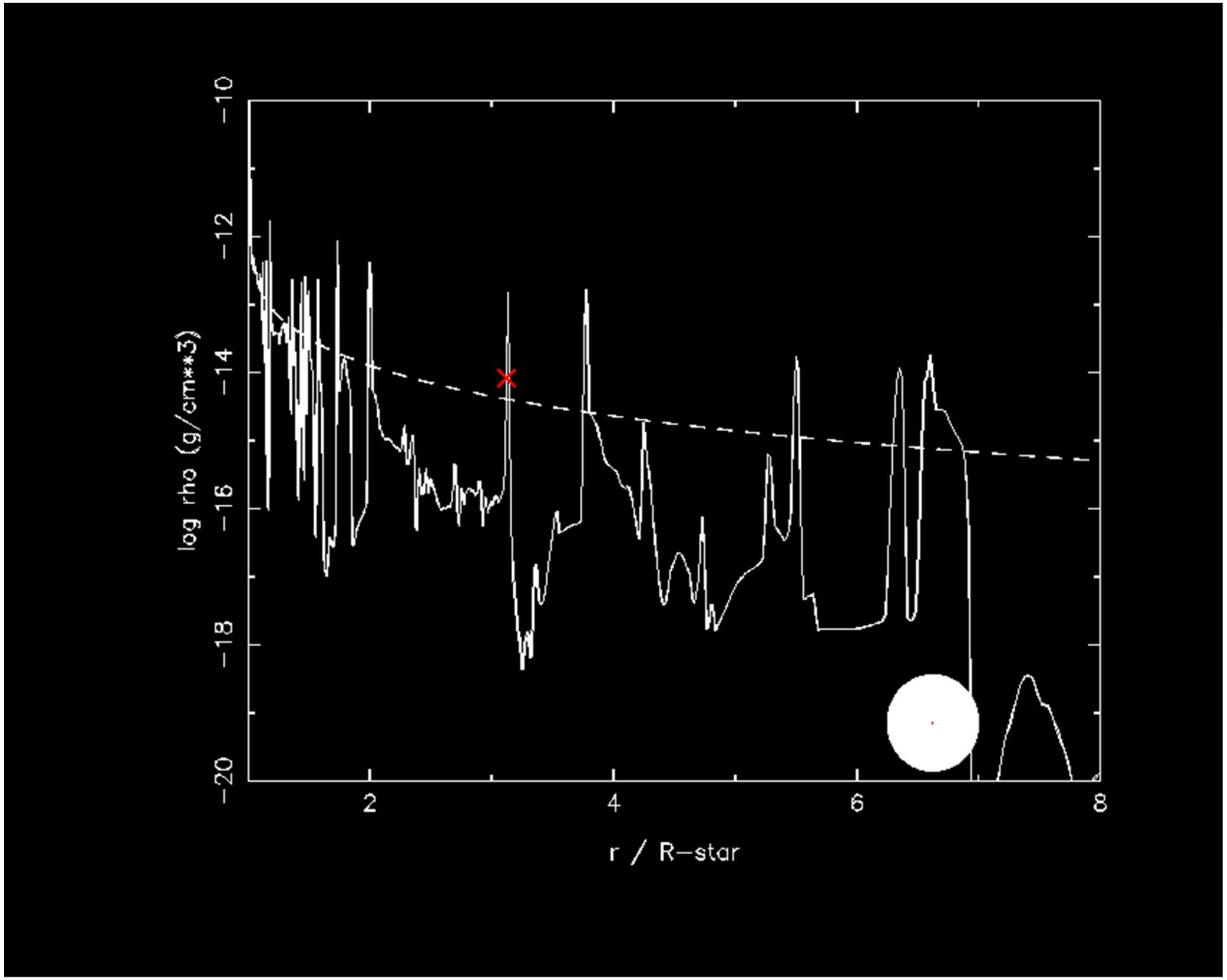

average hydro-structure not too different from stationary approx.: Most line profiles fairly similar, but effect ("clumping") needs to be accounted for in analysis

(very) hot gas
→ X-ray emission
(observed!)

From Runacres & Owocki, 2002, A&A 381

# Density evolution in an unstable wind

x
X-ray
"flash"



**Determine atmospheric parameters from observed spectrum**

**Required**
$T_{eff}$, log g, R, $Y_{He}$, Mdot, $v_\infty$, β (+ metal abundances)
(R stellar radius at $\tau_R = 2/3$)

**also necessary**
$v_{rad}$ (radial velocity)
v sin i (projected rotational velocity)

**Given**
- *reduced* optical spectra (eventually +UV, +IR, +X-ray)
- λ/Δλ, resolution of observed spectrum
- Visual brightness V
- distance d (from cluster/association membership), partly rather insecure
- NLTE-code(s), "model grid"

1. Rectify spectrum, i.e. divide by continuum (experience required)
2. Shift observed spectrum to lab wavelengths (use narrow **stellar** lines as reference):

$$\lambda_{lab} \approx \lambda_{obs}\left(1 - \frac{v_{rad}}{c}\right), \quad v_{rad} \text{ assumed as positive if object moves away from observer}$$

- **Alternative set of parameters**

  **L, M, R** *or*
  **L, M, $T_{eff}$** *or*
  **$T_{eff}$, log g, R** …

- **interrelations**

$$L = 4\pi R_*^2 \sigma_B T_{eff}^4$$

$$g = \frac{GM}{R_*^2}$$

- **Useful scaling relations**
  If L, M, R in *solar units*, then

$$R_* = \frac{L^{0.5}}{T_{eff}^2} \cdot 3.327 \cdot 10^7$$

$$\log g = \log\left(\frac{M}{R_*^2} \cdot 2.74 \cdot 10^4\right)$$

$$v_{esc} = \sqrt{R_* g (1-\Gamma) \cdot 1.392 \cdot 10^{11}}$$

$$\Gamma = s_e T_{eff}^4 / g \cdot 1.8913 \cdot 10^{-15}$$

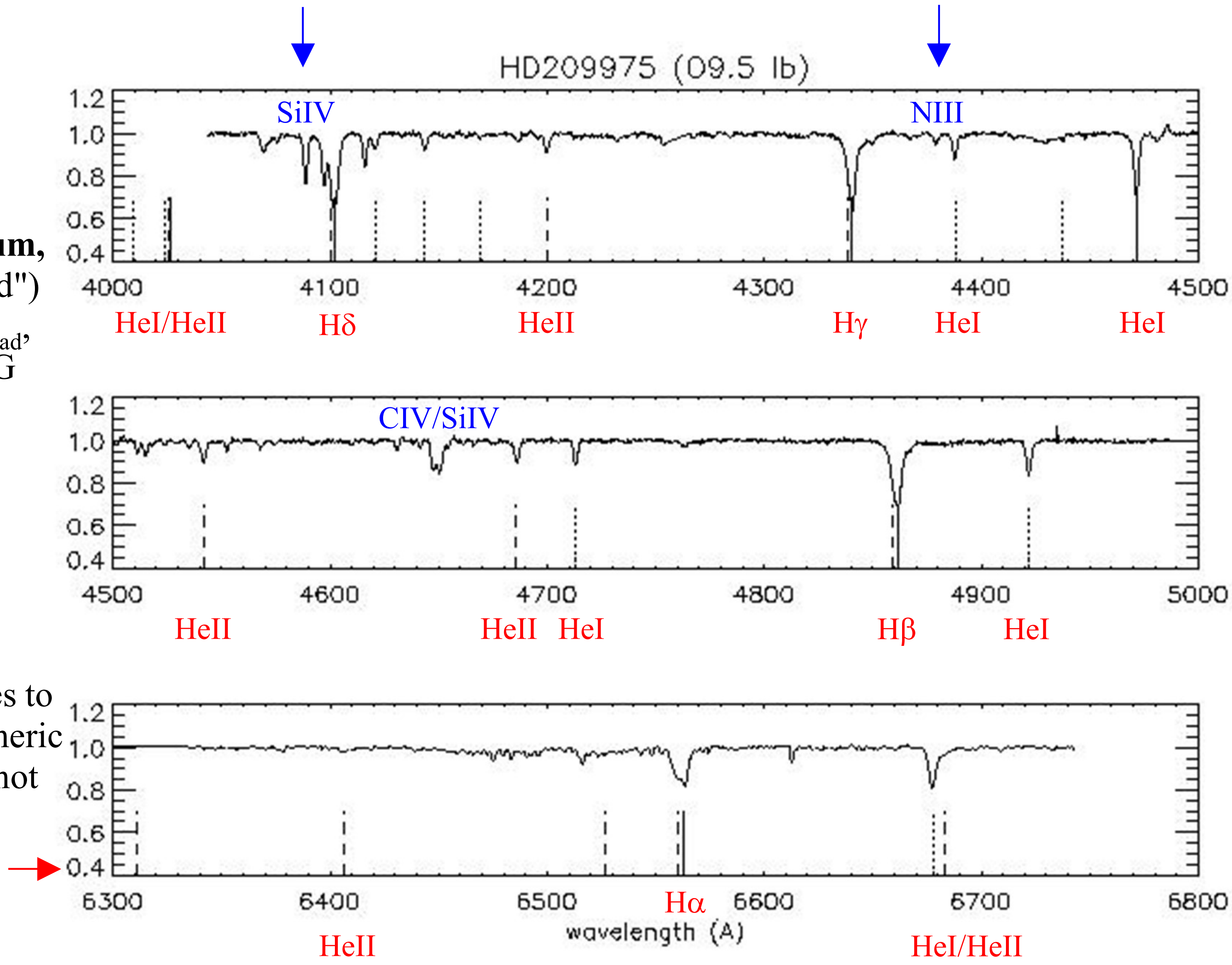
$$s_e = 0.4 \frac{1 + I_{He} Y_{He}}{1 + 4 Y_{He}}$$

with $I_{He}$ number of free electrons per Helium atom
(e.g.,=2, if completely ionized)

rectified
**optical spectrum,**
("blue" and "red")
corrected for $v_{rad}$,
of the late O-SG
*19 Cep*

___ Hydrogen
...... Helium I
- - - Helium II

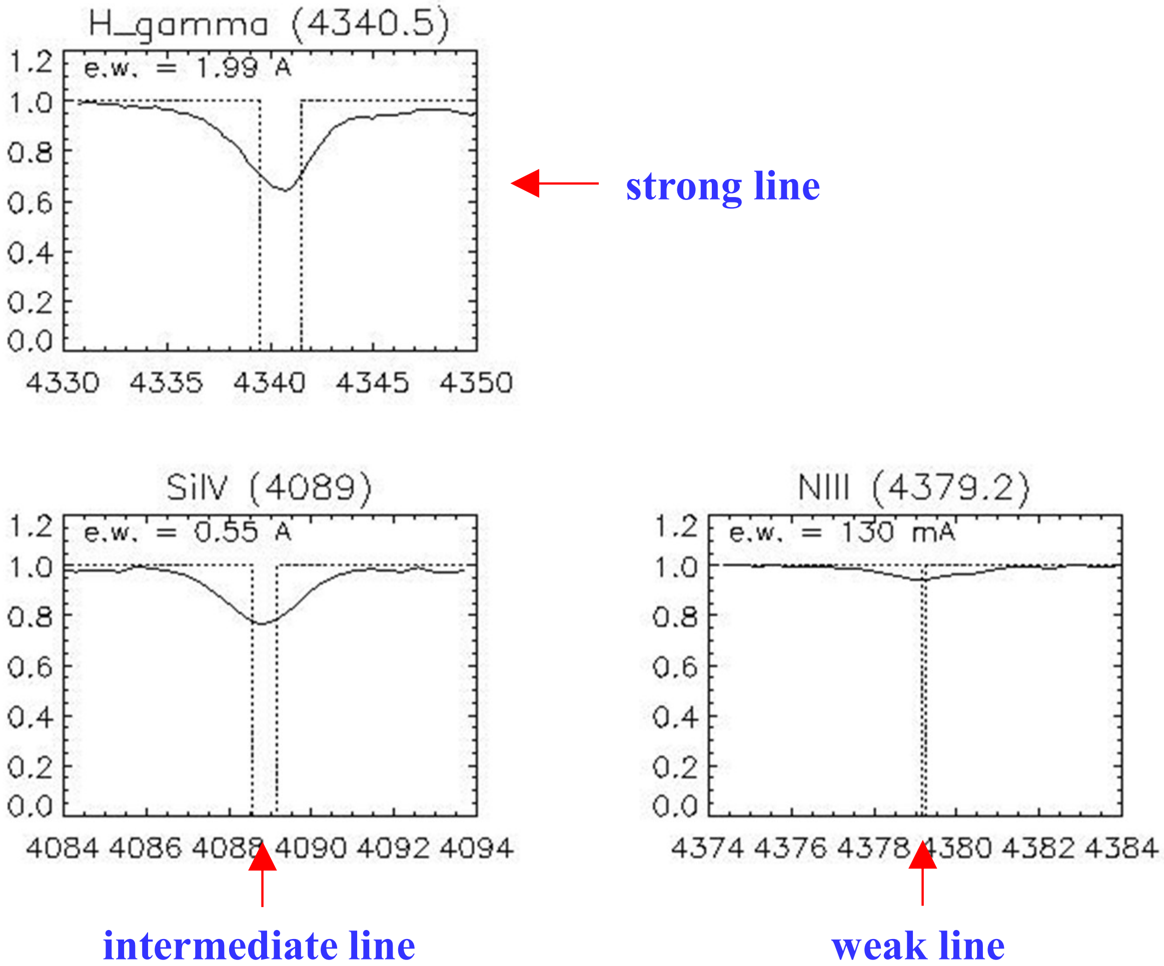
in "red":
"strategic" lines to derive atmospheric parameters in hot stars



equivalent width $W_\lambda = \int\limits_{\text{line}} \frac{H_{\text{cont}} - H_{\text{line}}(\lambda)}{H_{\text{cont}}} d\lambda = \int\limits_{\text{line}} \left(1 - R(\lambda)\right) d\lambda,$

area of profile under continuum, $\dim[W_\lambda]$ = Angstrom or milliAngstrom, mÅ

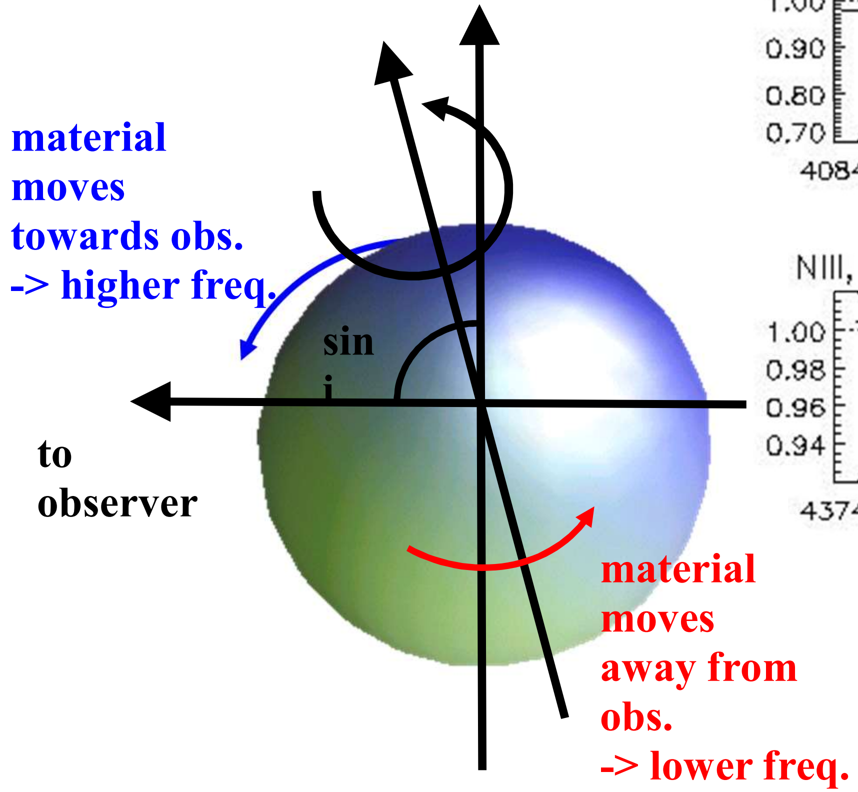
corresponds to width of saturated profile $(R(\lambda) = 0)$ with same area

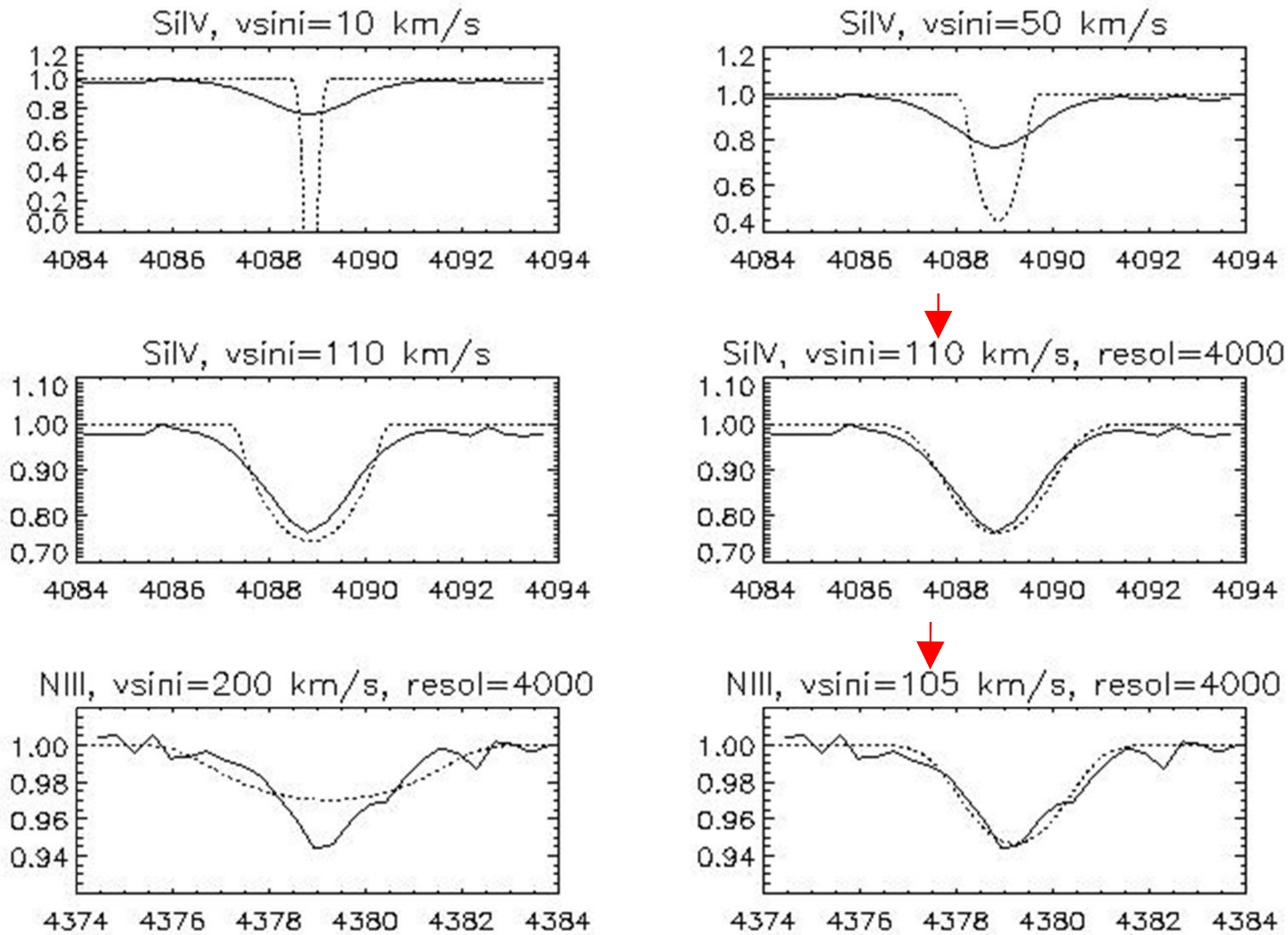


# Determine projected rotational speed v sin i

Use weak metal lines to derive v sin i: Convolve theoretical line with rotational profile.

Convolve finally with instrumental profile (~ Gauss) according to spectral resolution

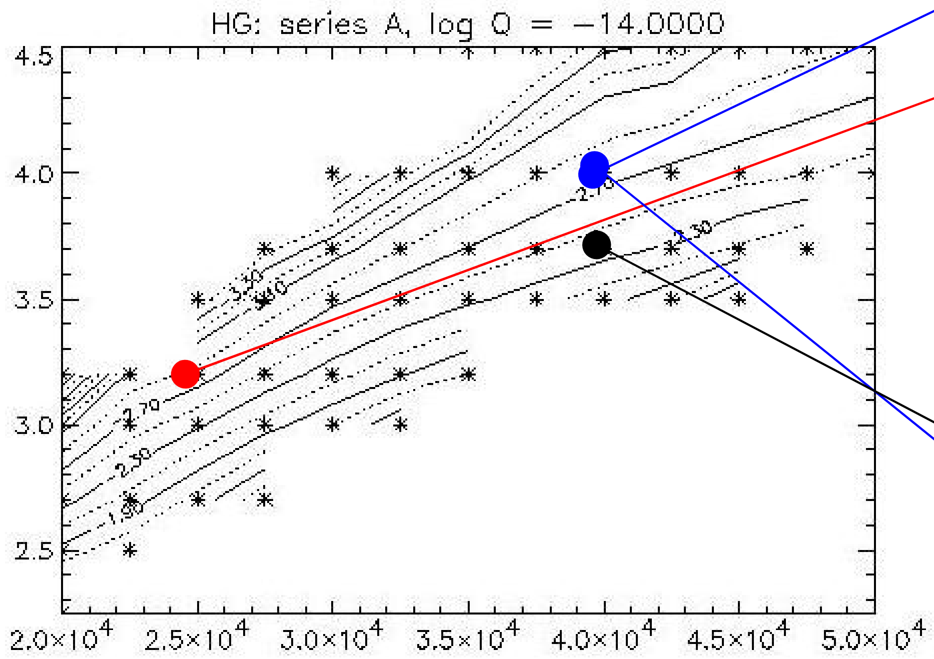


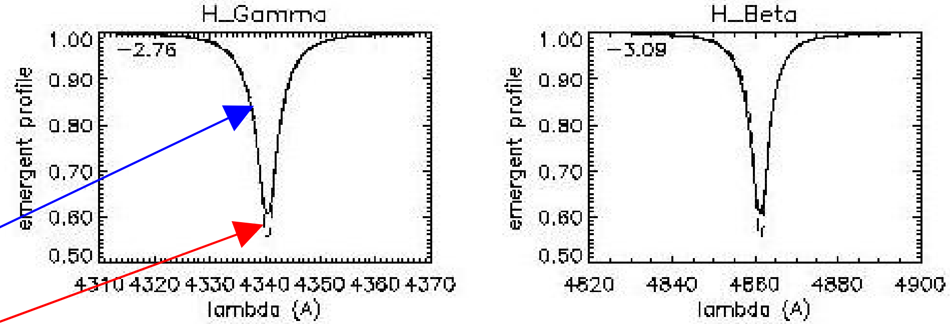

**Convolution with rotational and instrumental profile conserves equivalent width!!!**

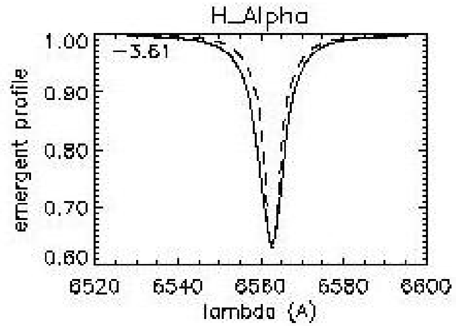
**Recent methods use a Fourier technique to infer vsini**

# Hγ - log g and $T_{eff}$

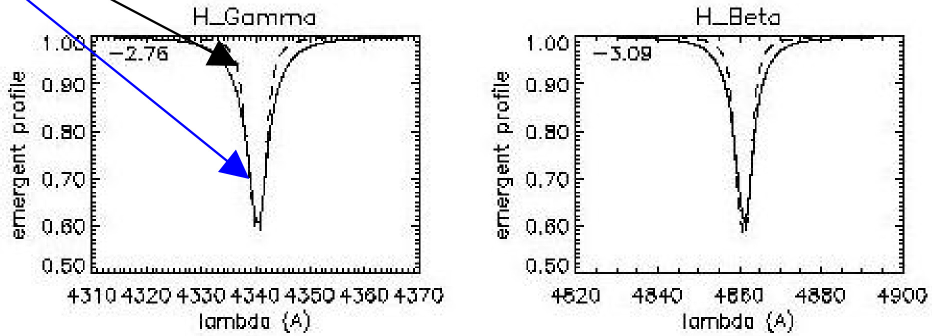


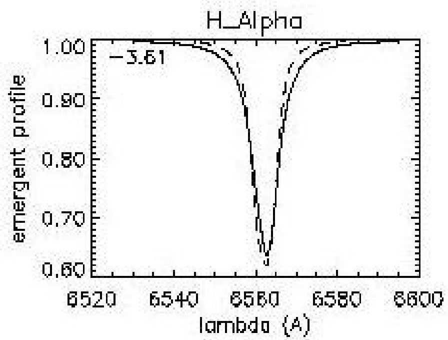

Iso-contours of equiv. widths for Hγ (from model grid), for solar Helium abundance and (very) thin winds




degeneracy of profiles:
(almost) identical lines for
$T_{eff}$ =40,000 and log g=4.0
and
$T_{eff}$ =25,000 and log g=3.2



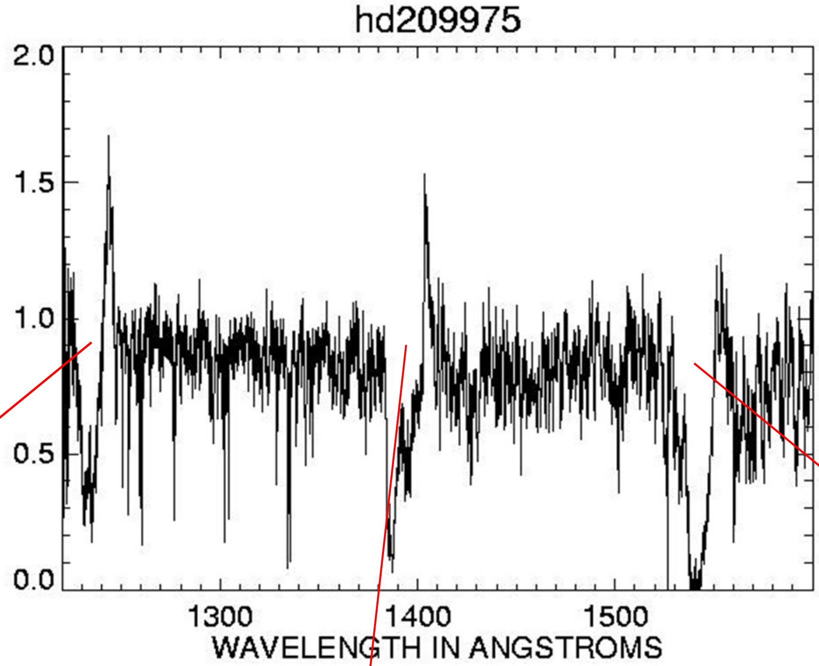

to derive $T_{eff}$, log g and $Y_{He}$,
at least 3 lines have to be fitted in parallel
(if no wind is present):

**Hγ** defines **log g** (for given $T_{eff}$ )
**HeII/HeI** define $\mathbf{T_{eff}}$ (for given log g)
absolute **strength of He lines** define $\mathbf{Y_{He}}$

usually, wind emission has to be accounted for (profiles shallower)

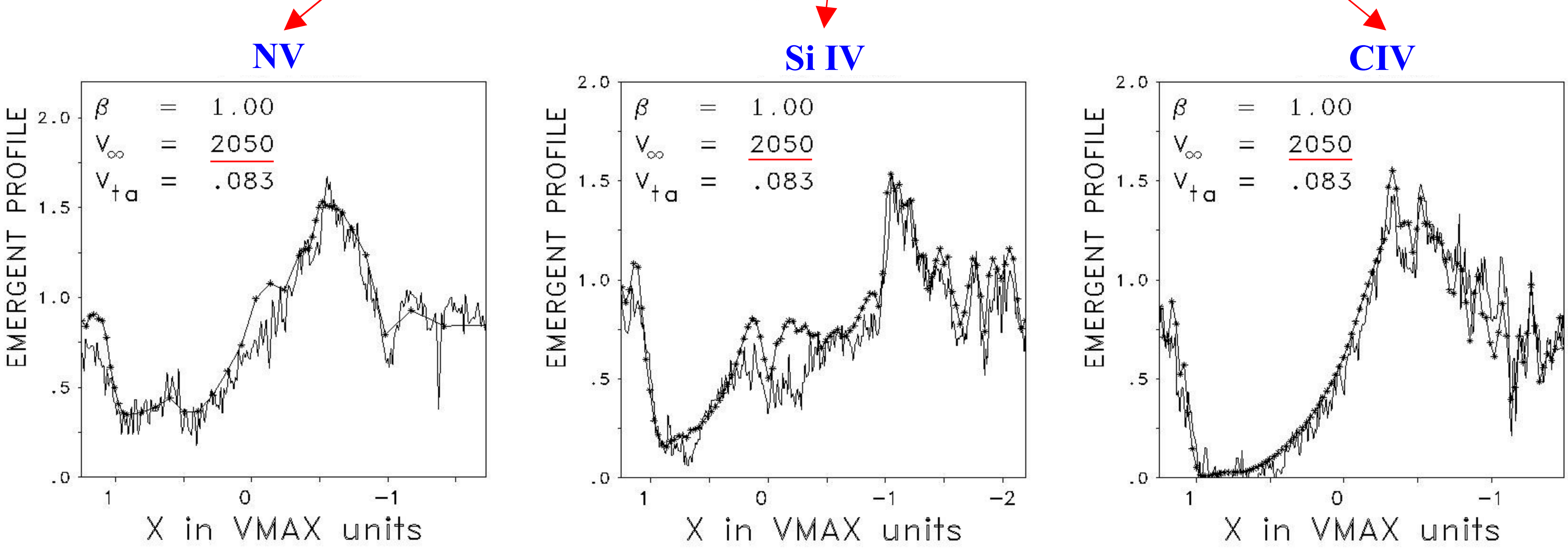
wings of Balmer lines
(Stark-broadened)
react strongly on electron-density (as a function of τ)
=> perfect gravity indicator

## Determination of terminal velocity from UV-P Cygni profiles




observation with IUE (International Ultra-violet Explorer)
no longer active

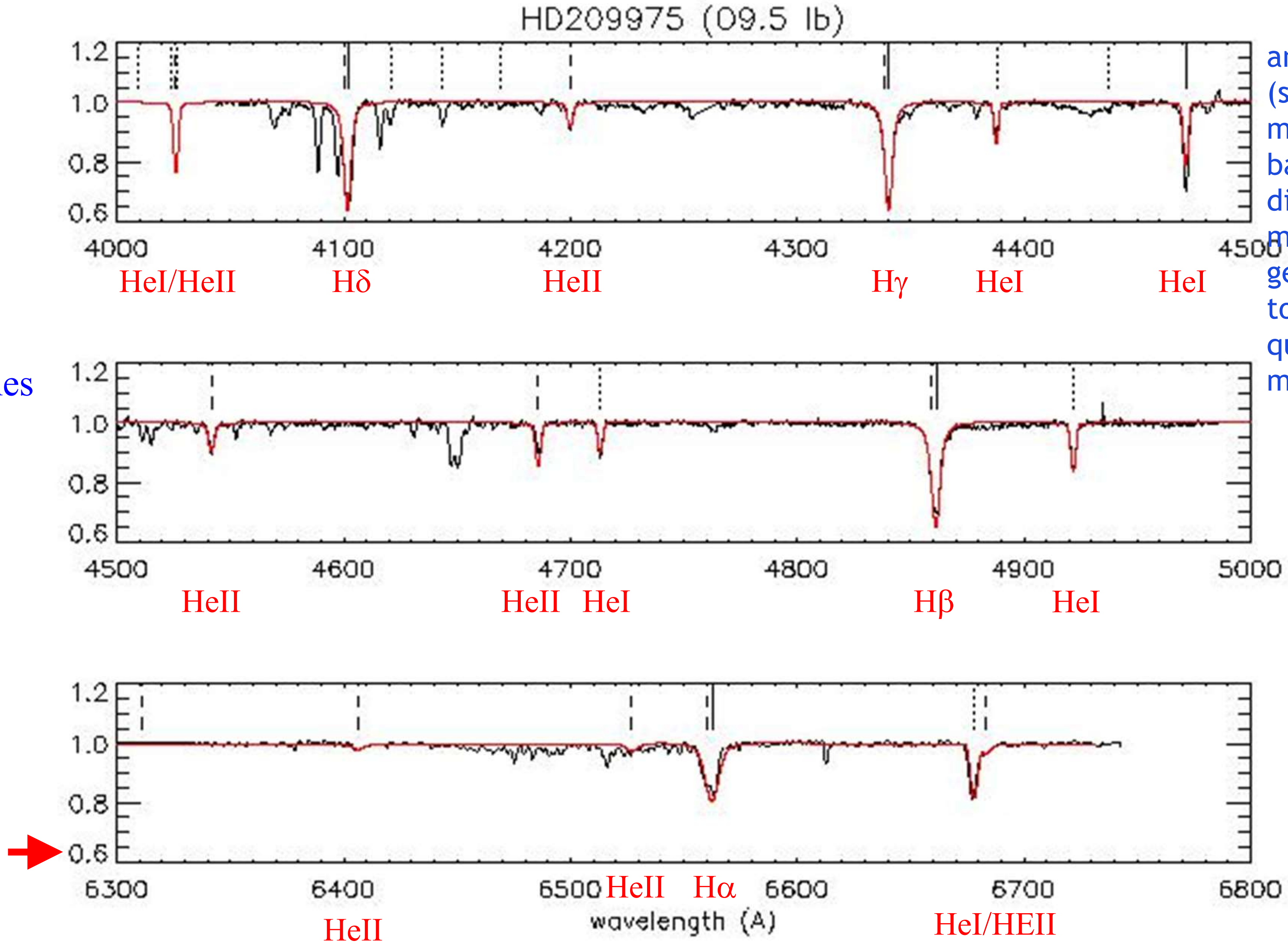
recent data (archive!) from HST ($\lambda$ >1200 Å) and FUSE ($\lambda$ > 911 Å)



# Line fitting = detailed comparison of observed and synthetic line profiles based on atmospheric models

___ Hydrogen
...... Helium I
- - -Helium II

indicated lines used for fits




analysis via (semi-) automatic methods, based on high-dimensional model grids or genetic algorithms, to optimize the fit quality for a multitude of lines

**derived parameters**

$T_{eff}$ = 31,000 K
log g = 3.17
log Q = -12.87
$Y_{He}$ = 0.10
$\beta$ = 1.0

with $v_\infty$ = 2050 km/ s we have

$$\log(\dot{M}/R_*^{1.5}) = -7.9$$

# Determination of stellar radius – if it cannot be resolved

- **IF** *you believe in stellar evolution models*
  - use **evolutionary tracks** to derive M from (measured) $T_{eff}$ and log g => **R**
  - transformation of conventional HRD into log $T_{eff}$ - log g diagram required
  - problematic for evolved massive objects, "mass discrepancy": spectroscopic masses (derived from spectroscopic analysis) and evolutionary masses often not consistent

- **IF** *you know the distance* **and** *have theoretical fluxes* (from model atmospheres), proceed as follows

- **IF** *you believe in radiation driven wind theory*
  - use **wind-momentum luminosity relation**

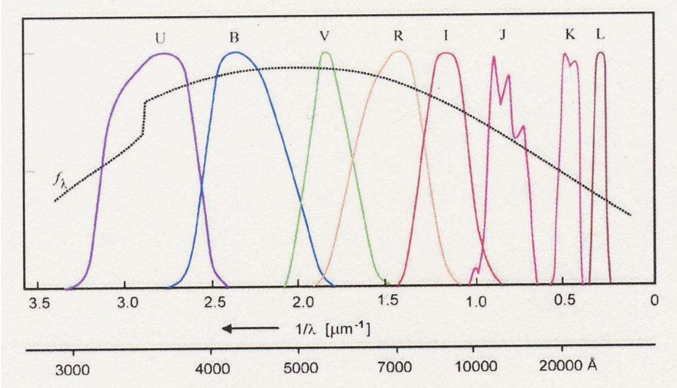
$$V = -2.5 \log \int_{\text{filter}} \mathcal{F}_\lambda S_\lambda d\lambda + \text{const}$$

$S_\lambda$ spectral response of photometric system

absolute flux calibration

$V = 0$ corresponds to $\mathcal{F}_\lambda = 3.66 \cdot 10^{-9}$ erg s$^{-1}$ cm$^{-2}$ Å$^{-1}$ at $\lambda_0 = 5{,}500$ Å outside earth's atmosphere

$\lambda_0$ *isophotal* wavelength such that $\int_{\text{filter}} \mathcal{F}_\lambda S_\lambda d\lambda \approx \mathcal{F}(\lambda_0) \int_{\text{filter}} S_\lambda d\lambda$ , $\int_{\text{filter}} S_\lambda d\lambda \approx 2895$ for Johnson V-filter




$\Rightarrow$

$$\text{const} = -2.5 \log(3.66 \cdot 10^{-9} \cdot 2895) = -12.437$$

$$M_V = -2.5 \log \left[ \left( \frac{R_* R_{\text{sun}}}{10 \text{ pc}} \right)^2 \int_{\text{filter}} \mathcal{F}_\lambda S_\lambda d\lambda \right] + \text{const}$$

$$5 \log R_* = 29.553 + (V_{\text{theo}} - M_V)$$

if $R_*$ in solar units, $M_V$ the absolute visual brightness (dereddened!) and

$V_{\text{theo}} - 2.5 \log \int_{\text{filter}} 4 H_\lambda S_\lambda d\lambda$ with $H_\lambda$ the *theoretical* Eddington flux in units of [erg s$^{-1}$ cm$^{-2}$ Å$^{-1}$]

**Alternatively,** use bolometric correction (BC)

Calibration for Galactic O-stars:

$$BC = M_{\mathrm{Bol}} - M_V \approx 27.58 - 6.8\log(T_{\mathrm{eff}}) \quad \text{(see Martins et al. 2005, A\&A 436)}$$

and definition of $M_{\mathrm{Bol}}$

$$\log\frac{L}{L_\odot} = 4\log\frac{T_{\mathrm{eff}}}{T_{\mathrm{eff},\odot}} + 2\log\frac{R_*}{R_\odot} = 0.4(M_{\mathrm{Bol},\odot} - M_{\mathrm{Bol}})$$

$$\log\frac{R_*}{R_\odot} = 0.2(4.74 - M_{\mathrm{Bol}}) - 2\log\frac{T_{\mathrm{eff}}}{5770} =$$

$$= 0.2(4.74 - M_{\mathrm{V}} - 27.58 + 6.8\log(T_{\mathrm{eff}})\ ) - 2\log\frac{T_{\mathrm{eff}}}{5770} =$$

$$= 2.954 - 0.2M_{\mathrm{V}} - 0.64\log(T_{\mathrm{eff}})$$ [valid only for O-stars with $Z \approx Z_\odot$]

remember relation between $M_V$ and V (distance modulus)

$M_V = V + 5(1 - \log d) - A_V$, $d$ distance in pc, $A_V$ reddening

d from parallaxes (if close) or cluster/ association/ galaxy membership (hot stars)
(note: clusters/ assoc. radially extended!)

For Galactic objects, use GAIA (if you believe DR2 parallaxes), or compilation by
Roberta Humphreys, 1978, ApJS 38, 309 *and/or*
Ian Howarth & Raman Prinja, 1989, ApJS 69, 527

**Back to our example**

HD 209975 (19 Cep): $M_v$ = -5.7
*check*: belongs to Cep OB2 Assoc., d ≈ 0.83 kpc (Gaia parallax: 1.165±0.15 mas = 0.85±0.11 kpc)
V = 5.11, $A_V$ = 1.17 => $M_V$ = -5.65, OK

From our final model, we calculate $V_{theo}$ = -29.08 => R = 17.4 $R_{sun}$
(Alternatively, by using BC, $M_V$ and $T_{eff}$ = 31kK, we would obtain R = 16.6 $R_{sun}$)

**Finally,** from the result of our fine fit, $\log(\dot{M}/R_*^{1.5}) = -7.9$ , we find $\dot{M} = 0.91 \cdot 10^{-6}$ $M_{sun}$ / yr

**Finished, determine metal abundances if required,**
**next star .... but end of lecture ...**